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Gravitational lensing reveals extreme dust-obscured star formation in quasar host galaxies

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ABSTRACT

We have observed 104 gravitationally lensed quasars at $z \sim 1\text{--}4$ with *Herschel*/SPIRE, the largest such sample ever studied. By targeting gravitational lenses, we probe intrinsic far-infrared (FIR) luminosities and star formation rates (SFRs) more typical of the population than the extremely luminous sources that are otherwise accessible. We detect 72 objects with *Herschel*/SPIRE and find 66 per cent (69 sources) of the sample have spectral energy distributions (SEDs) characteristic of dust emission. For 53 objects with sufficiently constrained SEDs, we find a median effective dust temperature of 38^{+12}_{-5} K. By applying the radio–infrared correlation, we find no evidence for an FIR excess that is consistent with star-formation-heated dust. We derive a median magnification-corrected FIR luminosity of $3.6^{+4.8}_{-2.4} \times 10^{11} L_{\odot}$ and median SFR of $120^{+160}_{-80} M_{\odot} \text{ yr}^{-1}$ for 94 quasars with redshifts. We find ~ 10 per cent of our sample have FIR properties similar to typical dusty star-forming galaxies at $z \sim 2\text{--}3$ and a range of SFRs $< 20\text{--}10\,000 M_{\odot} \text{ yr}^{-1}$ for our sample as a whole. These results are in line with current models of quasar evolution and suggests a coexistence of dust-obscured star formation and AGN activity is typical of most quasars. We do not find a statistically significant difference in the FIR luminosities of quasars in our sample with a radio excess relative to the radio–infrared correlation. Synchrotron emission is found to dominate at FIR wavelengths for < 15 per cent of those sources classified as powerful radio galaxies.

Key words: gravitational lensing: strong – galaxies: evolution – quasars: general – galaxies: star formation – infrared: galaxies – submillimetre: galaxies.

1 INTRODUCTION

Key to the study of galaxy formation and evolution is understanding the physical processes that drive star formation and the growth of active galactic nuclei (AGNs). The concurrence of these phenomena is thought to relate a coevolution driven by feedback from the AGN, which may quench or induce star formation in the host galaxy through interactions with the interstellar medium. The mechanism

of feedback may involve mechanical energy injection via AGN-driven jets, called ‘jet-mode’ or ‘radio-mode’ (Bicknell et al. 2000; Klammer et al. 2004), or radiative energy injection via winds, called ‘quasar-mode’, although these processes are not well understood (see Alexander & Hickox 2012, for review).

Hydrodynamical simulations of galaxy formation (Di Matteo, Springel & Hernquist 2005; Hopkins et al. 2005; Bower et al. 2006) and various observational studies (e.g. Page et al. 2004; Stevens et al. 2005; Coppin et al. 2008) support an evolutionary model, initially proposed by Sanders et al. (1988) and developed more recently by Hopkins et al. (2008), in which quasars are formed as a result

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of gas-rich major mergers. According to this scenario, luminous dusty star-forming galaxies (DSFGs) are merger-driven starbursts that represent a transition phase into dust-obscured quasars. Over time, feedback effects strip the quasar host galaxies of gas and dust, and the quasars become unobscured and ultraviolet (UV) luminous. These leave passive spheroidal galaxies when the quasar exhausts its supply of cold gas.

Quasars that are luminous in the far-infrared (FIR) to millimetre regime are therefore predicted to be in a transition phase of their evolution with high rates of dust-obscured star formation. Studying the properties of these sources can provide important information about the evolutionary process, particularly when compared to the large population of extreme starburst galaxies that were discovered through blind surveys with the Submillimetre Common-User Bolometer Array (SCUBA), *Herschel Space Observatory*, and now the Atacama Large Millimetre/submillimetre Array (ALMA).

Studies of FIR-luminous quasars, such as those in the SCUBA Bright Quasar Survey (Isaak et al. 2002; Priddey et al. 2003) and MAMBO/IRAM-30 m Survey (Omont et al. 2001, 2003), and more recent studies of quasars detected with *Herschel*/SPIRE (e.g. Pitchford et al. 2016) have found that these quasars are embedded within gas- and dust-rich starbursting galaxies, with star formation rates (SFRs) of $\sim 1000 M_{\odot} \text{ yr}^{-1}$, comparable to FIR-detected DSFGs. The low spatial density of FIR-luminous quasars, relative to DSFGs and UV-luminous quasars, has led some to argue for a quick transition from starbursting DSFG to an AGN-dominated quasar, with the FIR-luminous quasar phase being less than 100 Myr, and perhaps as short as ~ 1 Myr (e.g. Simpson et al. 2012). However, studies of individually detected quasars have mostly focused on significantly bright sources due to limitations in sensitivity or source confusion. While some recent progress has been made with the improved sensitivity and resolution of ALMA (Harrison et al. 2016; Banerji et al. 2017; Scholtz et al. 2018), resolutions of 100-pc are required to spatially resolve regions of star formation and AGN heating, which are still difficult to attain for the high-redshift Universe.

Other studies have instead used stacking to investigate the mean star formation properties of quasar host galaxies. These studies, which account for redshift and stellar mass, find no significant correlation between star formation and AGN activity, and find SFRs comparable to normal star-forming galaxies that lie on the galaxy main sequence (Rosario et al. 2013; Azadi et al. 2015; Stanley et al. 2017).

The next logical step in understanding the properties of quasar host galaxies at all luminosities requires an investigation of lower surface-brightness sources. Many of the limitations of confusion and sensitivity can be mitigated by observing quasars that have been magnified by a gravitational lens.

The advantages of observing strong gravitationally lensed quasars are threefold. The first is that magnification effects increase the apparent flux density such that a magnification factor of ~ 10 reduces integration time by a factor of ~ 100 . Sources with intrinsic flux densities below the confusion limit of field quasars can therefore be observed, probing the fainter end of the luminosity function (e.g. Impellizzeri et al. 2008). The second advantage is the increase in apparent surface area, which combined with source reconstruction methods, allow source structure to be probed on much smaller physical scales (e.g. Rybak et al. 2015a,b). A third advantage is that gravitational lensing has different systematic biases compared to field sources, while field observations tend to bias high luminosity or low-redshift sources, gravitationally lensed sources are more

biased towards compact higher redshift sources (typically $z > 1$) and less biased towards high intrinsic luminosities¹ (e.g. Swinbank et al. 2010). In combination, these methodologies allow for a more complete view of the quasar population to be constructed.

In this paper, we have targeted a sample of strong gravitationally lensed quasars with the *Herschel Space Observatory* (Pilbratt et al. 2010) and derive their dust temperatures, intrinsic FIR luminosities and dust-obscured SFRs. Previous work in this area has been undertaken by Barvainis & Ivison (2002), who detected 23 of 40 gravitationally lensed quasars and radio galaxies in their sample at 850 μm with SCUBA. They found dust emission broadly comparable to radio galaxies, in line with the AGN unification model, and no statistically significant difference AGN classified as powerful radio galaxies, as would be expected if they have the same host galaxy properties. We have observed 104 lensed quasars, including 37 of the Barvainis & Ivison sample, detecting 72 sources in at least one band with the *Herschel*/SPIRE. As our data cover shorter wavelengths, we are also able to determine the dust temperatures for the first time and infer whether the heated dust is due to star formation or AGN activity.

In Section 2, we present our sample selection, the relevant properties of the quasars in our sample, the parameters of the observations, and our data reduction methods. In Section 3, we report the results of the photometric measurements and the analysis of the radio-to-FIR spectral energy distributions (SEDs) of the sources. In Section 4, we show that the SEDs are consistent with dust heating due to star formation in the quasar host galaxies, and we compare our results with a sample of DSFGs at similar redshifts. Here, we also consider the contribution to the total radio emission from star formation processes for these quasars by considering the infrared-radio correlation. Finally, in Section 5, we present a summary of our results and discuss the future work that we will carry out with this sample.

Throughout, we assume the Planck Collaboration et al. (2016) instance of a flat Λ CDM cosmology with $H_0 = 67.8 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.31$, and $\Omega_{\Lambda} = 0.69$.

2 SAMPLE AND OBSERVATIONS

In this section, we describe our sample of gravitationally lensed quasars and present the observations that were carried out using the *Herschel Space Observatory*.

2.1 Sample selection

Our sample consists of all of the gravitationally lensed quasars that were observed with the *Herschel Space Observatory* (Pilbratt et al. 2010) using the Spectral and Photometric Imaging Receiver (SPIRE) instrument (Griffin et al. 2010). The vast majority of the observations came from our own open time project (Proposal ID: OT1_abcian_1). At the time of the proposal, these included all known quasars lensed by foreground galaxies. The majority of the sample are identified spectroscopically to be quasars, although some are identified as powerful radio galaxies without detections of prominent emission lines.² These sources are listed in the Sloan Digital Sky Survey Quasar Lens Search (SQLS) catalogue and CASTLES data base (Kochanek et al. 1999; Inada et al. 2012) and come

¹ Although these biases are dependent on whether the gravitational lens systems are selected via the lens or source populations.

² We refer to all these objects as quasars in this paper for simplicity.

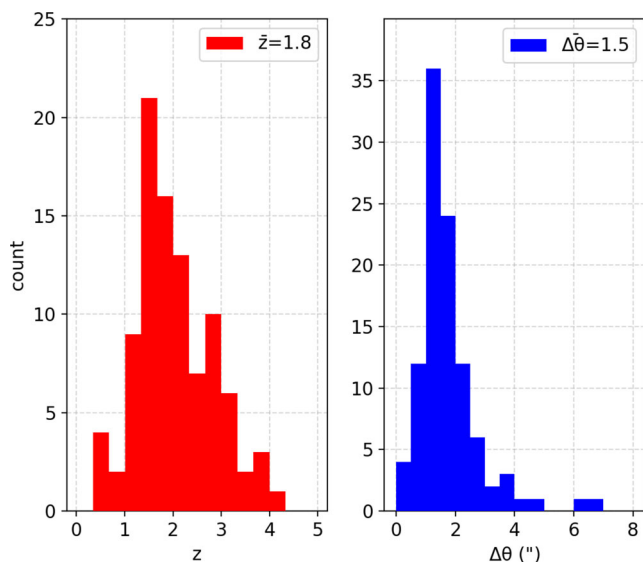


Figure 1. Left: The redshift distribution of 94 objects in our sample with known redshift, which has a median redshift of 1.8. Right: The lensed image separations, in arcsec, which have a median of 1.5 (excluding SDSS J1029+2623 which has a maximum image separation of 22.5 arcsec).

from a variety of surveys at optical and radio wavelengths. Our sample is quite heterogeneous, given the nature of the different surveys from which the targets were selected, but its size will allow us to draw representative conclusions on the relative FIR properties of jet-dominated and SF-dominated quasars, and provides a large parent sample from which further higher resolution observations of interesting individual objects can be made.

In total, there are 104 lensed quasars in our sample, the relevant properties of which are presented in Table A1 of the Appendix A. The redshift distribution and maximum image separations of the lensed quasars in our sample are presented in Fig. 1. The full width at half-maximum (FWHM) of the point spread function (PSF) in each band is 18, 24, and 35 arcsec for the 250, 350, and 500 μm bands, respectively. Therefore, all but 3 of our sample (Q0957+561, RX J0921+4529, and SDSS J1029+2623) have separations between the lensed images that are $<1/3$ of the smallest *Herschel*/SPIRE beam size, and can therefore be considered point sources for our study. The sample was observed in small map mode with one scan repetition per source, with a total integration time of 2–3 min per target, such that a source of 50 mJy will be detected at the 5σ level in the 500 μm band.

Of our quasar sample, 21 have 850 μm detections and 11 have 450 μm detections with SCUBA by Barvainis & Ivison (2002). Assuming magnifications from the literature and SEDs described by Yun & Carilli (2002) ($T_d = 58$ K, $\beta = 1.35$), nearly all of these detected sources (15 at 250 and 350 μm , 18 at 500 μm) would be below the confusion limits of *Herschel*/SPIRE were they not gravitationally lensed. It is therefore likely that the quasar population with intrinsic fluxes below those of previously detected field sources will be revealed in this study. Moreover, while SCUBA measurements lie on the Rayleigh–Jeans side of the thermal SED, the *Herschel*/SPIRE bands allow for better constraints on the peak of the SED, and thus, more accurate estimates of the characteristic FIR-luminosities and dust temperatures of the sample. We note that the previous study by Barvainis & Ivison (2002) assumed a dust temperature of 30 K for their sample, which may have biased their estimates of the FIR luminosities and inferred SFRs.

2.2 Radio properties

Radio emission from quasars may be associated with AGN (synchrotron) or star formation (synchrotron, free-free) processes. Quasars with radio jets are associated with jet-mode feedback, whereas quasars without these features are primarily radiative, so it is convenient to classify our sample based on their radio properties. There is a range of terminology and methods of classification employed in the literature to distinguish these groups, typically radio-‘loud’ and radio-‘quiet’ based on radio luminosity or radio-optical ratio. However, we find it more appropriate to group these by considering the operative feedback mechanisms. We have divided the sample into jetted (quasars with known jet-dominated radio emission) and non-jetted (quasars with star-formation-dominated radio emission and those where the dominant radio emission mechanism is unknown) dependent on whether there has been confirmation of the existence of a radio jet component with high-resolution radio data. For this, we have used the data from targeted observations for individual objects in the literature. Of the 34 quasars within the sample that we classify as jet-dominated radio sources, 31 are from the MIT-Green Bank Survey (MG; Langston et al. 1990), the Jodrell Bank-VLA Astrometric Survey (JVAS; Patnaik et al. 1992), the Cosmic Lens All-Sky Survey (CLASS; Myers et al. 2003; Browne et al. 2003), the Parkes-NRAO-MIT survey (PMN; Griffith & Wright 1993), and other radio surveys, all of which are dominated by radio-luminous AGN due to their respective flux-density limits. The remaining three sources are Q0957+561 (Garrett et al. 1994), H1413+117 (Stacey et al., in preparation), and HS 0810+2554 (Hartley et al., in preparation).

At low radio luminosities, composite AGN and star formation emission are likely, and differentiating between these possibilities is difficult. We define only two sources in our sample with established star-formation-dominated radio emission, RX J1131–1231 and IRAS F10214+5255. VLBI experiments to detect the radio core of these quasars suggest the radio emission is primarily due to star formation (Wucknitz & Volino 2008; Deane et al. 2013). In all other cases, the emission mechanism is undetermined, either because they are not detected at radio wavelengths or the detections are at too low an angular resolution to discriminate between compact (AGN) or extended (star formation) emission. We obtain the majority of these measurements from the National Radio Astronomy Observatory (NRAO) Very Large Array (VLA) Sky Survey (NVSS; Condon et al. 1998) and the Faint Images of the Radio Sky at Twenty-Centimeters (FIRST; Becker, White & Helfand 1995), both at 1.4 GHz and with beam sizes of 45 and 5 arcsec, respectively.

We show the rest-frame 1.4 GHz radio luminosities for the sample in Fig. 2. We include all quasars without evidence of jet-dominated radio emission in the non-jetted subsample for the time being, but refine these classifications using the radio–infrared correlation in Section 4.3.

2.3 Photometry

The sources have been observed with the *Herschel*/SPIRE instrument in three bands centred on 250, 350, and 500 μm , which effectively cover the rest-frame spectrum from 40 to 394 μm for the redshift range of our sample. The calibrated data were obtained from the *Herschel* Science Archive using the *Herschel* Interactive Processing Environment (HIPE; Ott 2010) version 14.0.0.

The photometry was performed using the SUSSEXtractor and Timeline Fitter algorithms within HIPE (Savage & Oliver 2007; Bendo et al. 2013) using recommendations in the SPIRE Data

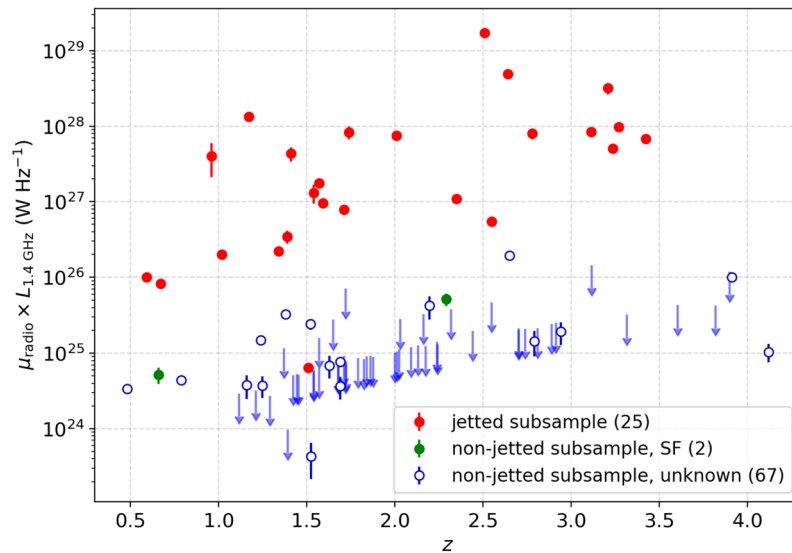


Figure 2. The rest frame 1.4 GHz radio luminosity density (interpolated or extrapolated from existing data) as a function of redshift for 92 objects in our quasar sample with radio measurements and a known redshift. Most of the upper limits are taken from FIRST or NVSS. The jetted subsample, with known jet-dominated radio emission, is shown in red. The non-jetted subsample includes two quasars with star-formation-dominated radio emission (shown in green) and 67 with unknown radio emission mechanism (shown in blue).

Reduction Guide.³ The Timeline Fitter performs point source photometry by fitting Gaussians to the baseline subtracted timeline samples, given source locations on the sky. The SUSSEXtractor method extracts point sources from the beam-smoothed, calibrated maps. We set a threshold of 3σ , where σ is the rms noise of the background around the source. While the Timeline Fitter gives more precise measurements and was the preferred method, SUSSEXtractor was occasionally more successful at extracting lower flux-density sources ($S_\nu \lesssim 30$ mJy). We place a detection limit of 3σ on the photometric measurements, where σ is the rms noise of the map, including confusion, given that we know the positions of the gravitational lens systems.

We explored the possibility of fixing the positions of source extraction with SUSSEXtractor to the ‘true’ sky positions in order to avoid an upward bias due to fitting to random noise spikes. While the effect of this is reduced as SUSSEXtractor fits to beam-smoothed maps, it has been noted to cause a bias in submillimetre measurements where the signal-to-noise ratio is low (e.g. Ivison et al. 2002; Coppin et al. 2005). As we would expect, there is a systematic upward shift in the flux densities measured when the position is left free. However, the change is generally not more than 10 per cent and within the photometric errors. We choose not to employ this method as we find there are often significant uncertainties on both the *Herschel* astrometry and the ‘true’ source position from the literature. We compared the extracted source positions of the five FIR-bright objects from our sample detected in ALMA to their ALMA positions (which have accurate and precise astrometry due to phase referencing) and find offsets up to several arcsec. This is consistent with other findings in the literature (e.g. Melbourne et al. 2012). We also find differences as much as several arcsec in the positions from optical or X-ray positions in the literature relative to the ALMA positions. There is an additional positional uncertainty as the targets are gravitationally lensed with a range of image separations (Fig. 1). Thus, the result of fixing the position

for source extraction would be a systematic down-shift of the extracted source flux densities. This bias can be more significant than the bias due to noise spikes and, as many sources are close to the detection limit, this would have a negative effect on the analysis. In any case, these uncertainties in the photometry are far lower than the uncertainties in the FIR luminosity and SFR due to SED fitting and the unknown magnification factor of the lensed systems (see Sections 3.4 and 3.3).

2.4 Source matching and confusion

Due to the sizes of the *Herschel*/SPIRE beams, we must also consider the contribution to the measured flux densities from field galaxies, including sources not associated with the target quasars or their lensing galaxies. For example, this could be due to dust-obscured star formation or AGN activity within the lensing galaxy at millimetre-wavelengths as has been seen in three gravitational lenses observed at high angular resolution with ALMA (ALMA Partnership et al. 2015; Paraficz et al. 2017; McKean et al., in preparation). We also note that at radio wavelengths, about 10 per cent of lensing galaxies have detected synchrotron emission from an AGN (McKean et al. 2005, 2007).

We take a series of steps to match the photometric data with our target quasars. We compare the extracted source position from SUSSEXtractor or the Timeline Fitter algorithm with the ‘true’ position of the lensing galaxy (where there is good astrometry, else the brightest lensed image) taken from the NASA Extragalactic Database (NED). We rejected extracted sources whose positional offsets are larger than half the FWHM of the SPIRE beam. We allow for some freedom in source fitting to allow for the combined uncertainties on the *Herschel* pointing, the ‘true’ source position, and source fitting. The extracted source positions are then cross-checked with nearby sources listed on NED to minimize the possibility of mismatching. In addition to this, we use detection in the 250 μ m band (which has the highest resolution and lowest confusion noise)

³ http://herschel.esac.esa.int/hcss-doc-14.0/print/spire_drg/spire_drg.pdf

as a prior to confirm a match at 350 and 500 μm .⁴ This strategy reduces the likelihood of contamination from field sources, but does not exclude the possibility of emission from the lensing galaxy or a nearby unknown FIR-bright field source being included in our photometric measurements.

While most of the targets appear uncontaminated, in some cases, blending is evident in the level 2 (fully calibrated) maps by visual inspection. For example, individual sources may be resolved in the higher resolution 250 μm maps, but become blended in the 500 μm band where the beam is largest. The blending results in overfitting by the source extractor and returns incorrect flux densities. We attempt to overcome this by simultaneously fitting to both the known target position and the blended source using the Timeline Fitter, where possible, then applying the same source matching criteria. Where this fails, we use SUSSEXtractor and fix the extracted position to the ‘true’ target position and blended source position (if known).

We can further identify confusion or mismatching by comparison of the *Herschel*/SPIRE data with the source SED. It is likely, based on inspection of their spectra, that we are unable to remove blended emission completely for eight sources: CLASS B0712+472, CLASS B0850+054, SDSS J0903+5028, CLASS B1152+200, Q 1208+101, CLASS B1359+154, SBS 1520+530, and Q 2237+030 (SEDs for all but CLASS B0712+472 are given in Fig. B1 of the Appendix B). In the cases of CLASS B0712+472 and Q 2237+030, there is too much blended emission to confidently measure the quasars, so we assume upper limits for all three bands by measuring the off-source rms noise of the maps. For CLASS B1152+200, this is the case at 350 and 500 μm . The remaining sources appear to have an additional contribution to their 500 μm measurement that is inconsistent with thermal dust emission or with synchrotron emission, based on their radio measurements. This may be due to errors in fitting to the blended source or further blending with nearby field sources. We identify known sources within a few arcsec of SDSS J0903+5028, Q 1208+101, Q 2237+030 that could be responsible and do not find evidence of confusion from the lensing galaxy for these objects. While SBS 1520+530 does have a star-forming lensing galaxy, the measured 500 μm flux density implies a flat spectrum that is inconsistent with the upper limits in the submillimetre/millimetre. In these cases, we assume upper limits for the 500 μm measurements that include confusion.

Almost all of the ancillary data that is used to derive the source SEDs is taken from literature, which typically consists of high resolution, targeted observations at millimetre-to-radio wavelengths, and lower resolution surveys at radio wavelengths. Where sources are detected and unresolved, we cannot be certain that they relate to a single source (the target, as opposed to a nearby companion or the lensing galaxy) without higher resolution observations on about arcsec-scales. The detections at 250 μm are matched to unique radio detections using the same matching criteria described previously, and these are assumed to relate solely to the quasar based on the assumptions that (i) the spatial density of quasars is lower than DSFGs, and so we are likely observing a single source rather than multiple sources, and (ii) as these quasars are intrinsically bright and gravitationally magnified, any companion would have to be similarly bright to contaminate our measurements, which is unlikely. Of course, further observations at higher angular resolution with millimetre-wavelength interferometers will better match the FIR emission detected here with the optical-to-radio counterparts of the

quasars. However, throughout this paper, we assume that the quasar is the sole source of the position-matched FIR emission detected with *Herschel*/SPIRE.

3 RESULTS AND ANALYSIS

In this section, we present the photometric results and describe the SED-fitting analysis used to determine the physical properties for each gravitationally lensed quasars within our sample.

3.1 *Herschel*/SPIRE measurements

The *Herschel*/SPIRE photometry for all of the sources observed in our sample is detailed in Table A1 of the Appendix A, and their SEDs, using all available data points, are shown in Fig. B1 in the Appendix B. Of the 104 sources observed, 72 are detected in at least one band down to a detection threshold of 3σ . Upper limits are given for those sources not detected at this confidence level. Of the sample, 10 targets suffer from contamination from the lensing galaxy or nearby field sources, which is apparent from their spectral properties and known properties of the lensing galaxies or nearby sources. This mostly affects the 500 μm band, due to the larger FWHM of the PSF and their rising synchrotron spectra at longer wavelengths (see Section 2.4).

The measured flux-density distribution for each of the bands, separated by their radio properties, is shown in Fig. 3 and the number of detections is given in Table 1. We use the two-sample Kolmogorov–Smirnov (K–S) test to compare whether the measured flux densities of the subsamples are consistent with the same underlying distribution. For all K–S tests in this work, we employ a Peto–Prentice Generalized Wilcoxon method⁵ for censored data using the *twosamp* task in the STSDAS statistics package within IRAF. The test returns a probability (p) for the null hypothesis, for which $p < 0.05$ we take as statistically significant. For our subsamples, the test returns probabilities of 0.44, 0.57, and 0.75 for the distributions of measured flux densities at 250, 350, and 500 μm , respectively. While the detection rates are slightly higher for the jetted quasars, the differences between the subsamples are not statistically significant.

3.2 Spectral slopes

In Figs 4 and 5, we show the spectral index between 850 and 500 μm ($\alpha_{500\mu\text{m}}^{850\mu\text{m}}$) and 500 and 250 μm ($\alpha_{250\mu\text{m}}^{500\mu\text{m}}$).⁶ 74 objects in our sample have detections in the FIR to submillimetre, including the 72 *Herschel*/SPIRE detected sources and a further 2 that have only submillimetre detections. In most cases, we find evidence for heated dust emission: of these 74 objects, we ascribe the emission in 69 cases (66 per cent of the sample) as being due to thermal dust emission from their rising or peaking spectra in FIR with frequency, relative to their submillimetre/millimetre/radio emission.

Of the five remaining sources that are detected in at least one band, there is no clear evidence for heated dust emission in the current data (see Fig. B2 for their SEDs). These sources are CLASS B1030+074, JVAS B0218+357, PKS 1830–211, PMN J1838–3427, and PMN J1632–0033. These sources do not have rising spectra in the FIR and have strong flat-spectrum synchrotron emission in the radio.

⁵ We employ this method as it is usually the most reliable and least affected by differences in the censoring patterns.

⁶ The spectral index is defined as a power law, $S_\nu \propto \nu^\alpha$, where S_ν is flux density and ν is frequency.

⁴ We make an exception for PMN J1632–0033 because of the completeness of the SED (see Fig. B2).

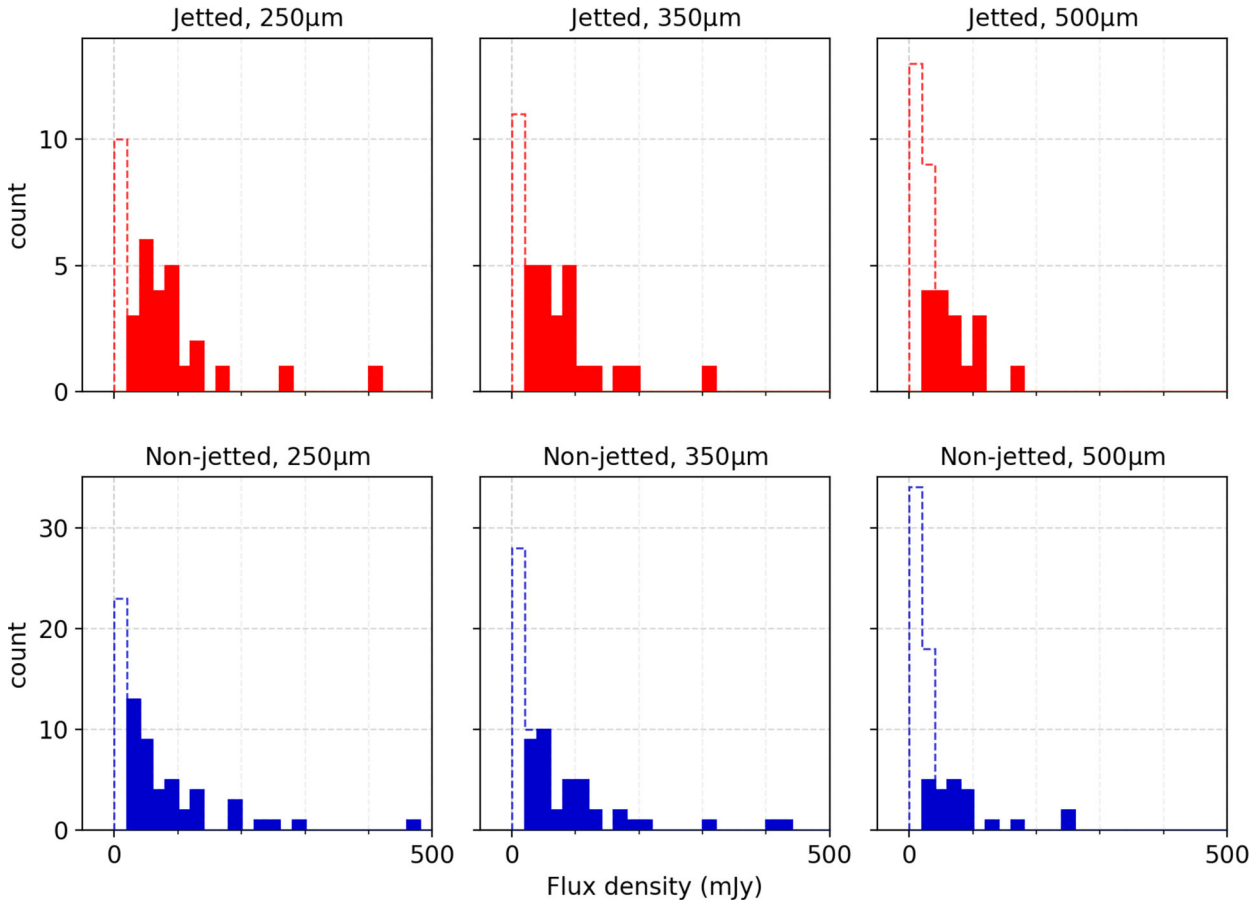


Figure 3. Number of sources binned by measured flux density for the three *Herschel*/SPIRE bands, divided into jetted (blue) and non-jetted (red) subsamples. 1.5σ limits of non-detections are stacked on top of the measured values and outlined with a dashed line. Note that PKS 1830–211 is excluded, for clarity, due to its high flux density ($S_{250\ \mu\text{m}} = 537\ \text{mJy}$, $S_{350\ \mu\text{m}} = 670\ \text{mJy}$, $S_{500\ \mu\text{m}} = 806\ \text{mJy}$).

Table 1. Number of detections in each *Herschel*/SPIRE band for the jetted and non-jetted subsamples.

	<i>N</i>	250 μm	350 μm	500 μm
Jets	34	24 (71 per cent)	23 (68 per cent)	16 (47 per cent)
No jets	70	47 (67 per cent)	41 (59 per cent)	23 (33 per cent)
Total	104	71 (68 per cent)	64 (62 per cent)	39 (38 per cent)

Unfortunately, these sources were not observed by the SCUBA and MAMBO surveys or were discovered too late to be part of the Barvainis & Ivison (2002) sample. Without measurements in the submillimetre regime, it is not clear how the synchrotron component falls off towards the FIR. In the cases of CLASS B1030+074, JVAS B0218+357, and PKS 1830–211, the flat-spectrum component continues into the millimetre regime, so it is likely there will be a significant contribution from optically thin synchrotron emission in the *Herschel*/SPIRE measurements (SED fitting of PKS 1830–211 is discussed further in C4 of the Appendix C). PMN J1838–3427 and PMN J1632–0033 do not have enough high-frequency data to extrapolate their spectra into the FIR. It is possible these sources have spectra comparable to CLASS B1127+385 or CLASS B1152+200, where submillimetre measurements or upper-limits dictate that synchrotron emission does not have a significant contribution in the FIR (see Fig. B2 of the Appendix B for their SEDs). JVAS B0218+357 and PMN J1838–3427 have measurements that appear characteristic of peaking dust emission, but this could also be explained by variability or a self-absorbed synchrotron

component. We fit thermal SEDs to the *Herschel*/SPIRE measurements for these five quasars to place upper limits on a possible contribution of heated dust to the FIR emission.

3.3 Magnifications

To derive the *intrinsic* properties of the quasars in the sample, the measurements must be corrected for their lensing magnification. Generally, these are obtained from the literature and are typically derived from an analysis of optical or radio gravitational lensing data. However, optical and radio components of quasars tend to be compact (size scales of $\leq \text{pc}$ to a few 10s of pc), and can result in very high magnification factors if the source is close to a lensing caustic. For example, JVAS B1938+666 has a radio magnification factor of 173 (Barvainis & Ivison 2002), whereas the $2.2\ \mu\text{m}$ infrared emission from the AGN host galaxy has a magnification of about 13 (Lagattuta et al. 2012). This presents a problem for accurately estimating the properties of this sample of gravitationally lensed quasars at FIR to submillimetre wavelengths, as the size

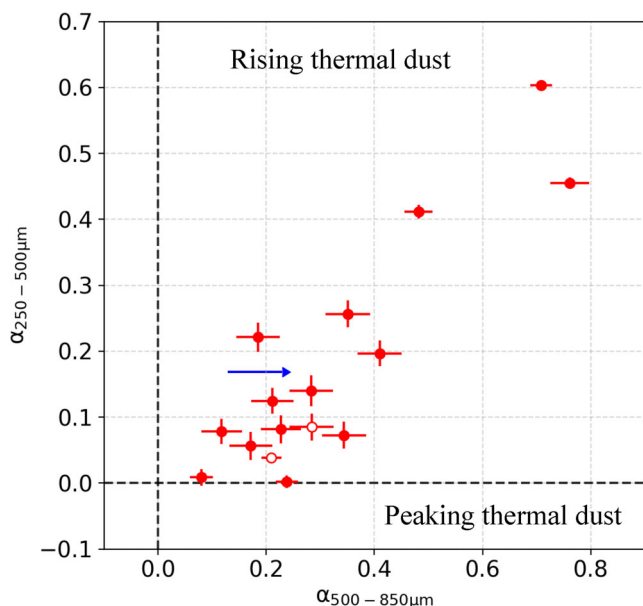


Figure 4. Spectral index with frequency of the high and low *Herschel*/SPIRE bands relative to 850 μm , for the 17 sources in the sample with previous submillimetre detections and three SPIRE detections. Open circles are measurements at the same wavelength, but not from SCUBA. Limits due to non-detections at 850 μm are shown in blue. The plot excludes PKS 1830–211, for clarity, due to its large negative spectral index ($\alpha = -0.5$).

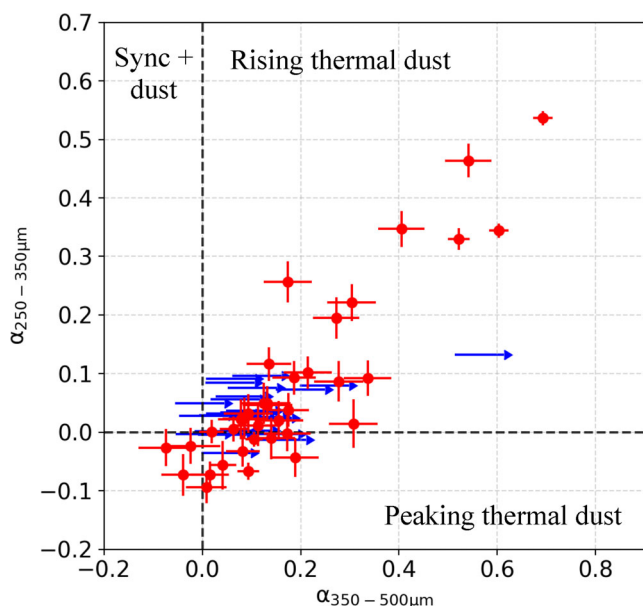


Figure 5. Spectral index of high-to-mid against mid-to-low *Herschel*/SPIRE for 63 sources with 250 and 350 μm detections. The positive quadrants contain rising spectra associated with dust. Three sources with falling spectra from 350 to 500 μm may have contamination from synchrotron emission, as discussed in Section 3.2. Lower limits due to non-detections at 500 μm are shown in blue. The plot excludes PKS 1830–211 due to its steep negative spectral index ($\alpha = -0.5$).

scales of AGN emitting regions may be anywhere from $\sim\text{pc}$ (in the case of the AGN core) to $\sim\text{kpc}$ (radio jets or a star-forming disc). Where the radio or optical magnifications are high, it is likely that the dust emission (assuming it is coincident with the quasar) will

be differentially magnified as only a small region will be close to the caustic and the overall magnification will be lower. Magnifications derived from optical or radio data are therefore unlikely to be accurate indicators of the actual dust magnification.

Only a few quasars in our sample have high-resolution observations in the FIR to submillimetre regime, thus we list the source properties given in Tables A1 and A2 uncorrected for lensing magnification. Known magnifications (μ_{SF}) in the FIR to submillimetre, based on dust or molecular gas tracers relating to star-forming regions, are given in Table 2. We assume that cold molecular gas has a similar extent, thus similar average magnification, as the star formation heated dust emission. Only two of the sources in the sample have resolved dust emission related to star formation. These are the Cloverleaf quasar and RX J0911+0551, which have magnifications of 11 and 19, respectively.⁷ For the intrinsic properties discussed below, we conservatively assume a magnification of $\mu_{\text{est.}} = 10^{+10}_{-5}$ for the sources without known magnifications. This is consistent with lens modelling of dust emission in *Herschel*-selected strongly lensed star-forming galaxies: Bussmann et al. (2013) find total magnification factors of 2–15 ($\bar{\mu} = 8$) for a sample of 20 observed with the Submillimeter Array (SMA), and Dye et al. (2018) find magnifications factors of 4–24 ($\bar{\mu} = 12.5$) in a sample of six observed with ALMA.

Magnifications of more than 20 are unlikely if the sources are extended more than $\sim 200\text{ pc}$, as discussed in the Barvainis & Ivison (2002) study. The two sources in our sample with reconstructed dust emission, RX J0911+0551 and the Cloverleaf quasar, both have dust emitting regions of $\sim 1\text{ kpc}$ in size (Tuan-Anh et al. 2017; Stacey et al., in preparation). Assuming these sizes are characteristic, our assumption of $\mu_{\text{est.}} = 10^{+10}_{-5}$ is likely representative of the magnifications of the sample, including a conservative uncertainty to account for outliers, and will provide an indication of the unlensed properties of the sample as a whole. The median values of the intrinsic properties we derive in the following analyses do not account for the factor of 2 error in the magnification because the assumption is taken for all but seven objects.

3.4 SED modelling

To constrain the physical properties of the FIR emission in each quasar host galaxy, we fit a combined non-thermal and thermal SED model to the *Herschel*/SPIRE data, along with any available data in the literature, excluding our measurements that are affected by confusion, as noted in Table A1. This model will account for any synchrotron component, in the case of the jetted targets, and any heated dust component of the SED. We use a power law with spectral index α ,

$$S_\nu \propto \nu^\alpha, \quad (1)$$

to describe the flux density (S_ν) as a function of frequency (ν) in the case of synchrotron emission. We do this only to estimate the contribution of synchrotron to the FIR spectrum, so do not attempt more complex fitting describing spectral turnovers (e.g. CLASS B1422+231 shown in Fig. B1 of the Appendix B). The SEDs of flat spectrum radio sources will likely turn down at higher frequencies and have a negligible synchrotron contribution in the FIR (e.g. CLASS B1127+385 in Fig. B1). In some cases, there is a suggestion that the synchrotron emission turns down towards the submillimetre (e.g. Q0957+561, suggested by an upper limit at

⁷ Rybak & Tuan-Anh private communication.

Table 2. Magnification values from the literature, with errors where given, and the data with which the lens modelling is performed (line or continuum). CO line emission is assumed to have a similar location and extent as star-formation-heated dust emission, thus a similar magnification.

Source	μ_{SF}	Method	Reference
APM 08279+5255	4.2	CO(1–0)	Riechers et al. (2009)
RX J0911+0551	18.7 ± 1.3	360 GHz continuum	Tuan-Anh (private communication)
Q 0957+561	7 ± 1	CO(2–1)	Krips et al. (2005)
IRAS F10214+4724	6 ± 1.5	CO(1–0)	Deane et al. (2013)
RX J1131–1231	7.3	CO(2–1)	Paraficz et al. (2017)
H1413+117	11.0	690 GHz continuum	Rybak (private communication) ^a
PSS J2322+1944	2.5	CO(2–1)	Carilli et al. (2003)

Note. ^aVenturini & Solomon (2003) also find a factor of 11 based on CO(7–6) line observations.

230 GHz), so we assume this does not contribute substantially to the FIR emission. We choose not to fit a synchrotron component where there is a single radio detection as we have no knowledge of the spectral behaviour, and, as these single measurements typically correspond to lower luminosities, the FIR contribution will be small.

We use a characteristic modified blackbody,

$$S_\nu \propto \frac{\nu^{3+\beta}}{e^{h\nu/kT_d} - 1}, \quad (2)$$

to describe the heated dust component, where h is the Planck constant, k is the Boltzmann constant, T_d is effective dust temperature, and β is the emissivity index, which determines the steepness of the Rayleigh–Jeans slope of the spectrum.

The PYTHON implementation EMCEE (Foreman-Mackey et al. 2013) was used to build a Markov Chain Monte Carlo (MCMC) analysis of the fitted SED for each data set, allowing the dust temperature and normalization as free parameters to sample the posterior probability distribution of the model. Where possible, we also leave β as a free parameter in the model, allowing for a test of the range of dust emissivities that are consistent with the data. However, fitting for β requires at least four data points to constrain the peak and Rayleigh–Jeans slope of the modified blackbody function. For many sources, our *Herschel*/SPIRE data are the only measurement in the FIR–submillimetre regime. Thus, we assume a value of $\beta = 1.5$ for these sources, as is frequently applied in the literature (e.g. Magnelli et al. 2012). Various combinations of T_d and β have been found for samples of high-redshift quasars. For example, Priddey & McMahon (2001) find an average of $T_d = 41 \pm 5$ K and $\beta = 1.95 \pm 0.3$, whereas Beelen et al. (2006) find an average of $T_d = 47 \pm 3$ K and $\beta = 1.6 \pm 0.1$ for their sample. For the sources that were not detected, or had only one detection to constrain the fit, we assume the median fitted dust temperature of the sample (38 K, see Fig. 6) and $\beta = 1.5$, and fit only for the normalization. For these sources with only one detection, we fit the 16th and 84th percentile values of the median fitted temperatures to estimate our errors on the FIR luminosity. As β is highly correlated with T_d , errors on the derived properties of sources without β -fitting may be underestimated. However, the FIR luminosity is not strongly affected by our assumptions due to the joint dependency of T_d , β and normalization, so this will not have a significant effect on the inferred values of L_{FIR} or SFR. This is unsurprising, as the luminosity is derived by the integral of the fit defined by the data points.

For the purpose of spectral fitting, the ten sources without a known redshift are assumed to have $z = 1.8$, equivalent to the median redshift of the sample. The choice of redshift significantly affects the luminosity distance; thus, these objects are not included in the overall statistics.

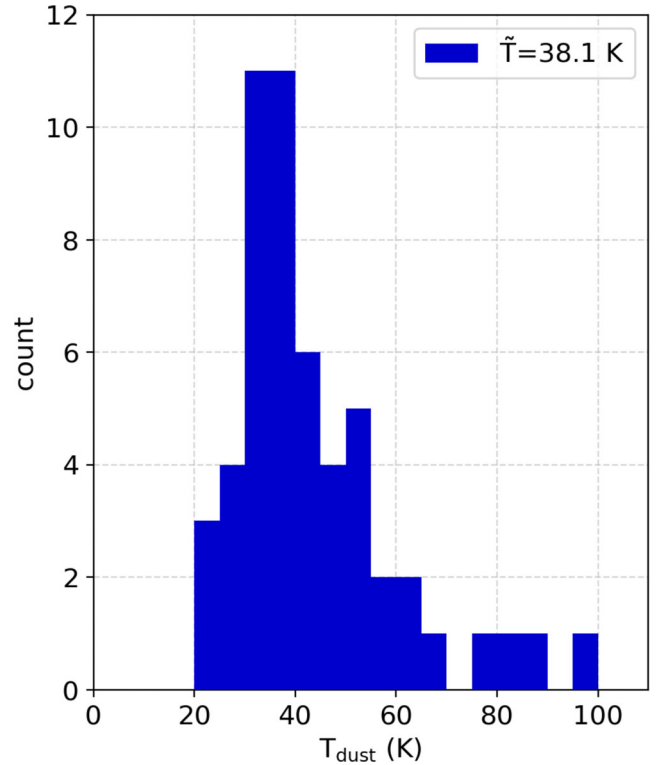


Figure 6. Histogram of effective dust temperatures for 53 quasars in the sample with temperature fitting, excluding those without a known redshift and synchrotron-dominated sources. The median dust temperature is 38^{+12}_{-5} K. Where two dust temperatures are fit, only the colder component is included here.

For three sources (APM 08279+5255, H1413+117, and IRAS F10214+4724), there are sufficient data in the mid-IR (MIR) to motivate fitting a two-temperature dust model. Table 3 shows the number of sources fitted with each combination of spectral parameters.

We have included posterior probability distributions of the MCMC output of the SED fit for three sources to show the correlation between the various fitting parameters and highlight the effect of sparse sampling of the SED. Fig. 7 shows the result for APM 08279+5255, where there is sufficient data to fit seven spectral parameters. In Fig. 8, we compare the results for two sources: PSS J2322+1944, where the peak of the dust emission and the Rayleigh–Jeans slope are both well constrained, and Q 1208+101, where the peak is poorly constrained.

Table 3. Number of sources fitted with each set of spectral parameters. The number of upper limits are given in brackets. Five synchrotron-dominated sources (JVAS B0218+357, CLASS B1030+074, PMN J1632–0033, PKS 1830–211, and PMN J1838–3427) are fitted with single temperature modified black bodies to compute upper limits on contributions from any dust emission to their FIR spectra (for PKS 1830–211 this includes a synchrotron component). Of the sources with no temperature fitting, 30 sources have no detections and 10 have only one detection.

Spectral fit	Number of sources
Two $T_{\text{dust}} + \beta$ + synchrotron	3
Single $T_{\text{dust}} + \beta$ + synchrotron	8
Single $T_{\text{dust}} + \beta$	10
Single $T_{\text{dust}} + \text{synchrotron}$	5
Single T_{dust}	33(5)
Fixed T_{dust}	10(30)
Total	69(35)

3.5 Physical properties

The dust temperature, FIR luminosity and SFR of 69 gravitationally lensed AGN in our sample are listed in Table A2. We give both the results from the MCMC analysis and those from least-squares fitting. The values from the MCMC analysis are the median, 16th and 84th percentile of the posterior probability distributions. We give upper limits for the remaining 30 sources of the sample with insufficient detections, and the further 5 that appear to be synchrotron dominated. For clarification, a summary of median values of various properties we derive and an explanation of the objects included in these statistics are given in Table 5.

A histogram of the dust temperatures derived directly from the modified blackbody fits is shown in Fig. 6. Where a model with two dust components is fit, we include only the colder component here. We find a median of $T_d = 38_{-5}^{+12}$ K for the 53 sources with fitted dust temperatures and known redshifts, 51 of which have dust temperatures < 60 K that, as we discuss in Section 4, can be reasonably attributed to be due to heating by star formation.

The FIR luminosity (L_{FIR}) is derived for all sources with fitted modified blackbody spectra by integrating the fitted modified blackbody spectra between the rest-frame wavelengths 40 and 120 μm , using the definition of the FIR regime given by Helou et al. (1988), that is,

$$L_{\text{FIR}} = \frac{4\pi D_L^2}{(1+z)} \int_{40 \mu\text{m}}^{120 \mu\text{m}} S_{\nu, \text{rest}} d\nu, \quad (3)$$

where z is the redshift and D_L is the luminosity distance. We then extrapolate to the total infrared luminosity (8–1000 μm ; rest frame), using the colour correction factor of 1.91 given by Dale et al. (2001) (i.e. $L_{\text{IR}} = 1.91 L_{\text{FIR}}$) to correct for the contribution from MIR spectral features. The methodology used to calculate the SFR is that given by Kennicutt (1998), assuming a Salpeter initial mass function,

$$\text{SFR} (\text{M}_{\odot} \text{ yr}^{-1}) = \frac{L_{\text{IR}}}{5.8 \times 10^9}, \quad (4)$$

where L_{IR} is in units of L_{\odot} .

In Figs 9 and 10, we show the FIR luminosity uncorrected for lensing magnification based on the fitted SED models as a function of redshift and dust temperature, respectively. Note that the dust temperature is invariant to the lensing magnification in the absence of strong differential magnification. The uncorrected luminosity as a function of redshift (Fig. 9) shows a clear trend in the data, from $\sim 10^{12} L_{\odot}$ at redshift 0.5–1 to $\sim 10^{13}$ – $10^{14} L_{\odot}$ at redshift 3–4.

We use the *bhkmeth* task in the STSDAS statistics package to compute the Kendall correlation test, taking into account the luminosity upper limits. The Kendall statistic τ quantifies the degree of correlation (-1 for a strong anticorrelation, 0 for no correlation, and 1 for a strong positive correlation) and the significance of this is given by the probability (p), for which < 0.05 we take as statistically significant. Our data show a correlation in temperature with redshift ($\tau = 0.64$, $p = 4 \times 10^{-4}$) and in temperature with L_{FIR} ($\tau = 0.77$, $p < 1 \times 10^{-4}$).

We find a large spread of L_{FIR} , as is clear from Fig. 10: the low luminosities are associated with low temperatures and low redshifts, and the high luminosities with high redshifts and generally higher dust temperatures. The six sources with measured luminosities $< 1.5 \times 10^{12} L_{\odot}$ (corresponding to magnification-corrected SFRs $< 50 \text{ M}_{\odot} \text{ yr}^{-1}$) are associated with dust temperatures < 25 K and/or redshifts $z < 1.5$. These trends can be explained by observational bias, given the wavelength limits of the *Herschel*/SPIRE bands and the flux limits of our observations, which approximately correspond to the luminosity detection limit shown in Fig. 9.

In Fig. 11, we present the dust emissivity as a function of dust temperature. We find a strong anticorrelation ($\tau = -1.30$, $p < 1 \times 10^{-4}$) between the parameters. This effect is expected for internally heated dust clouds (Juvela & Ysard 2012b). However, it is not clear to what extent this correlation is a reflection of the ‘true’ β – T_d relation, as the effect of source blending or observational noise may cause an artificial steepening of the anticorrelation, as noted by Juvela & Ysard (2012a). The shallow β values may be a result of fitting a composite of dust emission from star formation and AGN heating with a single grey body, when multiple components are required (see HS 0810+2554, Section C1 in the Appendix C). These problems further highlight the need for more multifrequency data to better sample the SED and to reduce fitting errors due to observational noise.

4 DISCUSSION

In this section, we investigate the global properties of the sample of gravitationally lensed quasars and compare them with samples of unlensed FIR bright quasars and star-forming galaxies.

4.1 Comparison to DSFGs

We select the unlensed DSFGs observed with *Herschel*/SPIRE by Magnelli et al. (2012) (hereafter, M12) for comparison with our lensed quasar sample, after correcting for the magnifications (see Section 3.3). The M12 objects are canonical DSFGs selected in ground-based submillimetre surveys with no evidence of a strong AGN component. We select the 46 unlensed objects with known redshifts from the M12 sample, which have redshifts $2.2_{-0.7}^{+0.5}$, FIR luminosities of $5.4_{-3.7}^{+3.1} \times 10^{12} L_{\odot}$, and SFRs of $1800_{-1200}^{+1000} \text{ M}_{\odot} \text{ yr}^{-1}$ (median, 25th and 75th percentiles). If the dust emission we detect is related to dust-obscured star formation, we expect dust temperatures comparable to DSFGs. Further to this, if DSFGs are antecedent to quasars, we would expect some fraction of our quasar sample to be FIR luminous with SFRs that are comparable to DSFGs.

We show in Figs 12 and 13 the FIR luminosity against redshift and against dust temperature, respectively, for our sample and the M12 DSFGs. The median fitted dust temperature of our sample is 38_{-5}^{+12} K (ranges are the 25th and 75th percentiles) for 53 objects with sufficiently constrained SEDs and redshifts. This is consistent with the M12 DSFGs, which have a median temperature of 36_{-9}^{+4} K, typical of star-forming galaxies at $z \sim 2$.

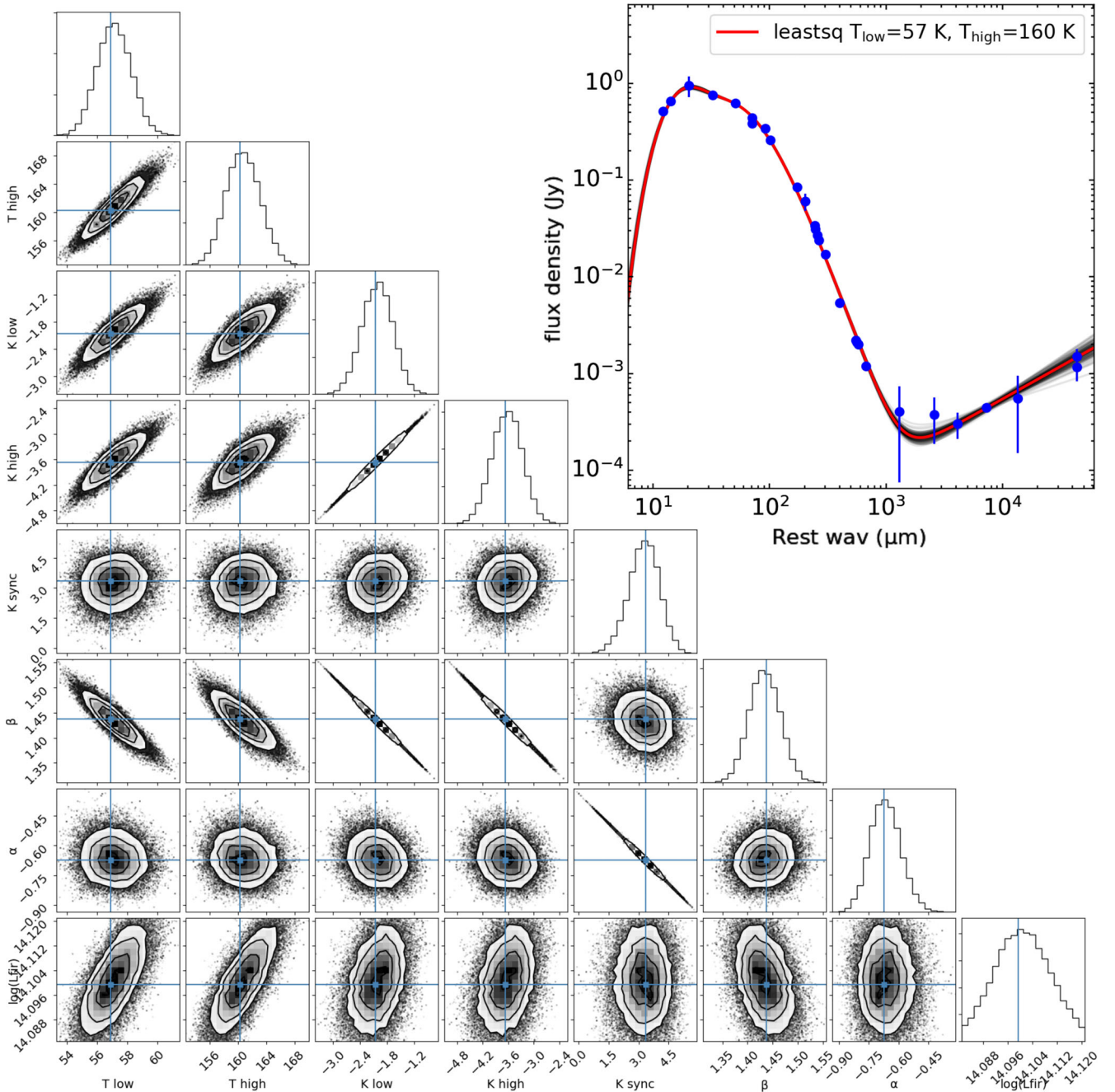


Figure 7. Two-dimensional probability densities of the MCMC output for the SED fitting of APM 08279+5255, fit with all parameters: T_1 , low dust temperature (observed); T_2 , high dust temperature (observed); K_1 , K_2 , K_3 , lognormalizations of the dust and synchrotron fits; β , the emissivity index; α , the synchrotron power-law index. Also shown is L_{FIR} . The blue points on the corner plot show the least-squares parameters. The SED is shown above, with 100 random samples of the MCMC in black and the least-squares model in red.

We apply the Kaplan–Meier (K–M) method to estimate the underlying distribution of FIR luminosities, taking into account the upper limits, using the task *kmestimate* in the *STSDAS* statistics package. This method assumes a randomly censored distribution: while this seems counter-intuitive as we have a fixed flux-density limit, our redshift range spans several orders of magnitude in luminosity distance so the sample is effectively randomly censored. We find a K–M estimated median, 25th and 75th percentiles of $3.6^{+4.8}_{-2.4} \times 10^{11} L_{\odot}$ for the intrinsic luminosities for 94 objects with redshifts, including 63 detections and 31 upper limits, compared to $5.8^{+7.1}_{-2.7} \times 10^{11} L_{\odot}$ for just the 63 objects with detected dust emission and

known redshifts. The K–M estimated median of SFRs in our sample is $120^{+160}_{-80} M_{\odot} \text{yr}^{-1}$, with $190^{+230}_{-90} M_{\odot} \text{yr}^{-1}$ for just the objects with detected dust emission.

The SEDs determined here clearly demonstrate that 69 objects (66 per cent) of our sample show evidence for heated dust emission at FIR to submillimetre wavelengths (these SEDs are shown in Fig. B1 of the Appendix B). Also, given the similar dust temperatures of our lensed quasar sample and the DSFGs studied by M12, there is at least circumstantial evidence that this dust heating is due to star formation activity. Approximately 10 per cent of the sample have extreme SFRs $>1000 M_{\odot} \text{yr}^{-1}$ comparable to

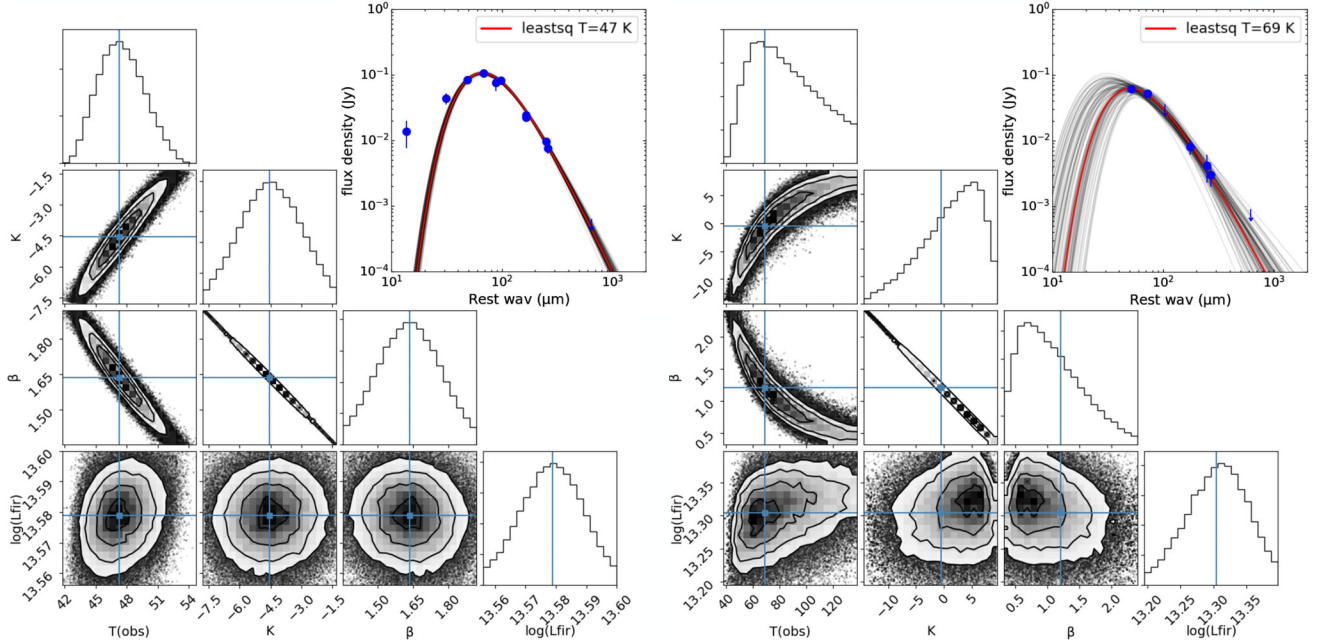


Figure 8. Two-dimensional probability densities of the MCMC results of SED fitting for PSS J2322+1944 (left) and Q 1208+101 (right), showing the correlations between the spectral parameters: T , observed dust temperature; K , normalization; β , the emissivity index. Also shown is L_{FIR} . The blue points on the corner plots show the least-squares parameters. The SEDs are shown above the corner plots, with 100 random samples of the MCMC in black and the least-squares model in red.

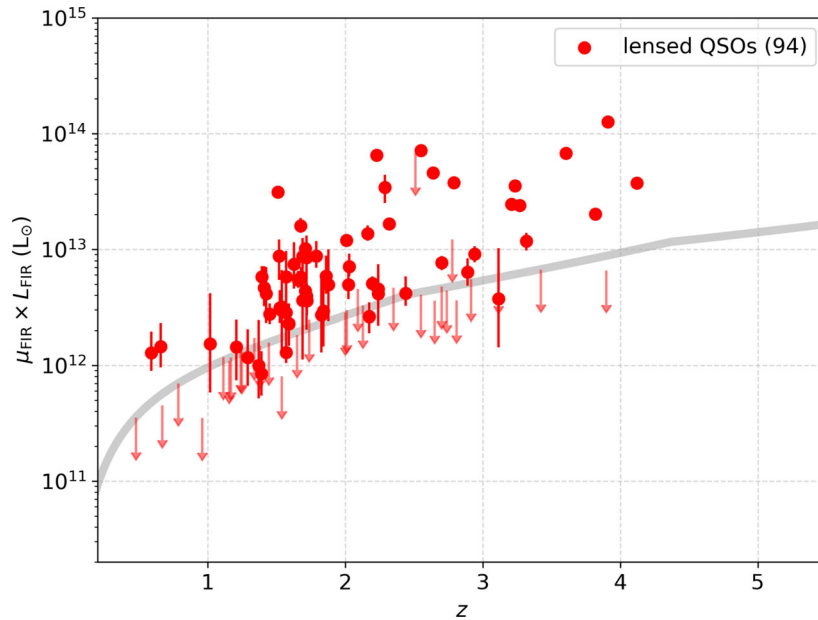


Figure 9. FIR luminosity (40–120 μm) against redshift for the 94 objects in our sample with known redshift. This includes 53 with fitted dust temperatures, 10 with fixed dust temperatures, and 31 upper limits. The measured luminosities are shown in red, with no magnification correction. The grey line shows the estimated luminosity detection limit for a source with $T_d = 38\text{ K}$ and $\beta = 1.5$, assuming a 3σ detection limit based on the mean rms noise in each *Herschel*/SPIRE band. This is an overestimate at low redshift, as sources with lower dust temperatures will be preferentially detected; likewise, this is an overestimate at high redshift where there will be bias towards higher temperature sources.

typical, unlensed DSFGs at $z \sim 2\text{--}4$ detected in *Herschel*/SPIRE. The SFRs of these lensed quasars are consistent with sources that are transitioning from DSFGs to UV-bright quasars according to the Sanders et al. (1988) evolutionary model. The rest of the detected sample have still extreme SFRs similar to the lower luminosity DSFGs selected at $z < 1.5$, but there is no clear cut-off at low SFR.

The range of SFRs we find ($< 20\text{--}10\,000\text{ M}_\odot\text{ yr}^{-1}$) is consistent with sources at different stages of evolution, and is not too surprising given the heterogeneous nature of our sample. Nevertheless, the high detection rate in the *Herschel*/SPIRE bands implies that most quasars are FIR-luminous sources with a strong coexistence of extreme dust-obscured star formation and AGN activity. This

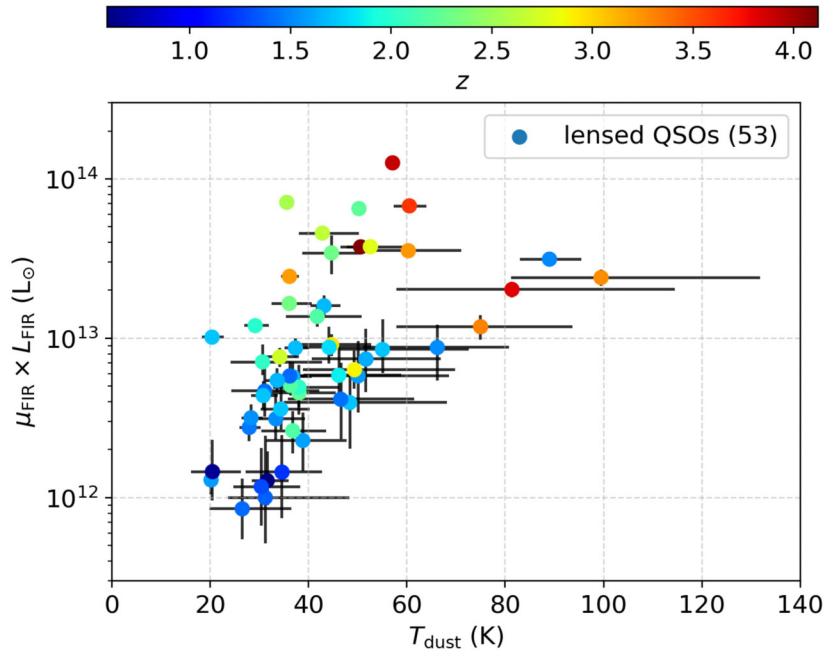


Figure 10. FIR luminosity against fitted dust temperature for the 53 objects in our sample with fitted dust temperature and known redshift. The colour scale indicates source redshift. The luminosities are not corrected for lensing magnification.

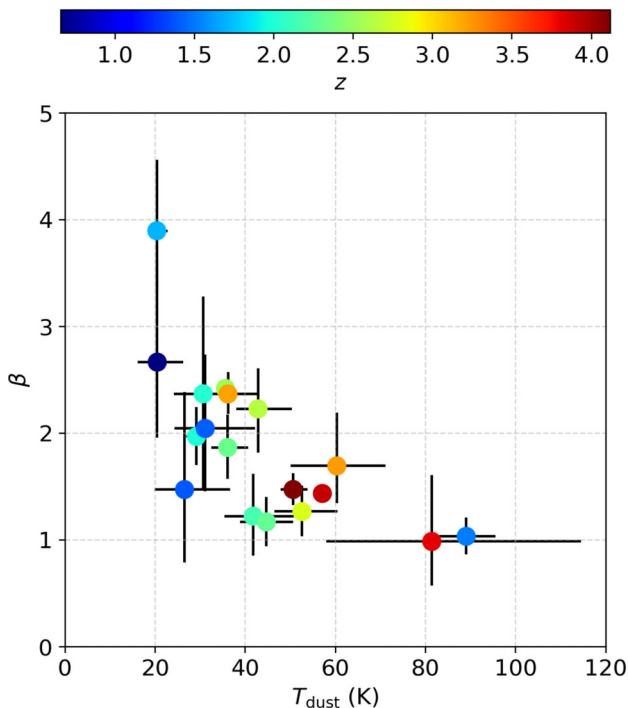


Figure 11. β against effective dust temperature for 18 objects in our sample with fitted β and known redshift. The colour scale indicates source redshift.

result implies a transition time from quasar-starburst to unobscured, gas-poor system of the order of the lifetime of the quasar (i.e. $\lesssim 100$ Myr), rather than much shorter time-scales of ~ 1 Myr, as has been suggested by Simpson et al. (2012). Further studies, including spectral line data of the molecular gas in these systems, are required to understand anything of the gas reservoirs and depletion times,

or make any conclusions regarding possible implications for AGN and stellar feedback.

4.2 Comparison by radio properties

Using the two-sample K–S test (described in Section 3.1), we compare the derived L_{FIR} distributions of the jetted subsample with the remaining quasars in the sample. These results are compiled in Table 4. We do not correct for the lensing magnification here to prevent any bias due to our assumptions about the magnification factor of the heated dust. Including all measured luminosities and upper limits (excluding those without redshifts), the test returns a probability of 0.23 that these subsamples are drawn from the same underlying distribution. The K–M estimated median, 25th and 75th percentiles of the FIR luminosity is $1.6^{+10}_{-1.5} \times 10^{12} L_{\odot}$ for the jetted subsample and $3.7^{+3.5}_{-2.4} \times 10^{12} L_{\odot}$ for the remaining quasars. While the jetted subsample has a larger distribution of luminosities, the test suggests that the difference in the luminosity distributions is not statistically significant.

A combination of systematic biases and the smaller size of the jetted subsample may affect our findings. We note that quasars with radio jets tend to be hosted in more massive galaxies (Mandelbaum et al. 2009), thus our jetted subsample may be biased towards larger FIR luminosities and hence larger SFRs. At present, we do not have data to account for the stellar mass of the galaxies in our sample.

The median, 25th and 75th percentiles of the redshifts are $1.7^{+0.9}_{-0.3}$ for the jetted subsample and $1.8^{+0.6}_{-0.3}$ for the non-jetted. Despite the fact that the sample selection between these groups is different, with the jetted quasars generally selected in the radio by source properties and the non-jetted objects typically selected by lens population, the redshift distributions of the two groups are similar and therefore not a substantial source of systematic bias.

Our result is consistent with the conclusions of Barvainis & Ivison (2002), who found no statistically significant difference in 850 μm luminosity between their samples of quasars and radio galaxies.

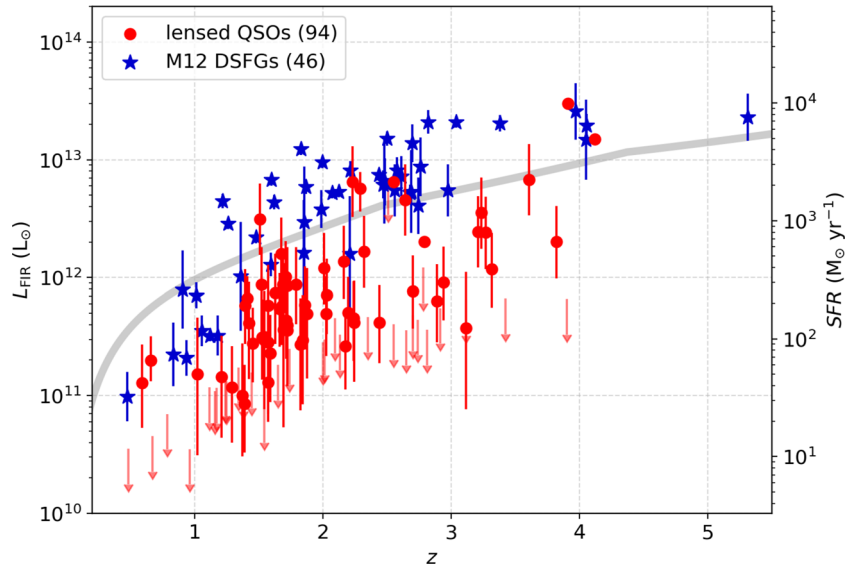


Figure 12. FIR luminosity (40–120 μm) and equivalent SFR against redshift for lensed quasars in this sample (94 objects, excluding those without known redshifts) and for the M12 DSFGs (46 objects). The quasar luminosities are magnification corrected (see Section 3.3). Magnification factors for seven sources are given in Table 2. Where the lensing magnification of the FIR emission is unknown, we assume a value of 10_{-5}^{+10} . The grey line shows the luminosity detection limit for a source with $T_d = 38$ K and $\beta = 1.5$, assuming a 3σ detection limit.

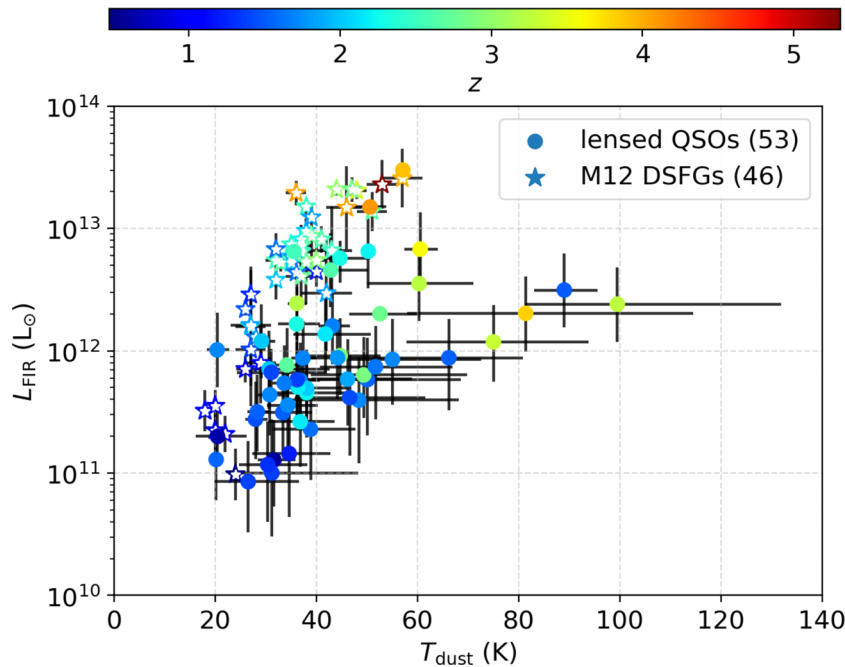


Figure 13. FIR luminosity against dust temperature for our lensed quasar sample (53 objects, with fitted dust temperature and known redshift) and for the M12 DSFGs (46 objects). Quasar luminosities are magnification corrected (see Section 3.3). The colour scale indicates source redshift.

Other studies have found no significant differences in the star-forming properties of quasars by radio mode. Harris et al. (2016) analysed a sample of optically luminous quasars at redshifts between 2 and 3 through stacking of which 95 per cent are undetected individually with *Herschel*/SPIRE. They find a mean SFR of $300 \pm 100 M_{\odot} \text{yr}^{-1}$, consistent with our overall result, but find no correlation with black hole accretion. A recent study by Pitchford et al. (2016) of higher luminosity quasars with *Herschel*/SPIRE also find no relation between AGN accretion/outflows and the FIR properties of their host galaxies. Alternatively, Kalfountzou et al. (2014) studied a stacked sample of quasars and do find a positive corre-

lation between jet activity and FIR luminosity for jetted quasars, defined by a $5 \text{ GHz}/4000 \text{ \AA}$ ratio > 10 . However, they find average SFRs to be comparable for jetted and non-jetted quasars, except at low optical luminosities.

Overall, our results do not point towards there being an enhancement in the FIR luminosity of jetted quasars and radio galaxies, relative to the non-jetted quasars in our sample. However, a more complete understanding of this result, particularly given the contradictory studies discussed above, will require detailed observations of individual objects. In this respect, our investigation of lensed quasars from within our sample will again be important since it will

Table 4. Kaplan–Meier (K–M) estimated 50th, 25th, and 75th percentile ranges of the FIR luminosity distributions of the jetted and non-jetted subsamples. This includes the 92 objects in our sample with radio measurements and known redshifts. We give the number of values, N , and in brackets the number of upper limits. We also give the results when the subsamples are selected by q_{IR} value, and the K–S test probability that these samples are drawn from the same underlying distribution.

		$N(\text{lims})$	K–M L_{FIR} ($10^{12} L_{\odot}$)	K–S test
Uncorrected	Jetted	15(10)	$1.6^{+10}_{-1.5}$	0.23
	non-jetted	48(19)	$3.7^{+3.5}_{-2.4}$	
Corrected	Jetted	19(17)	$1.3^{+7.8}_{-1.3}$	0.06
	non-jetted	43(13)	$4.1^{+3.3}_{-1.4}$	

allow the radio-jets, (stellar) host galaxy, and the heated dust to be mapped on small angular-scales.

4.3 Radio–infrared correlation

The radio–infrared luminosity correlation for star-forming galaxies has been well established for several decades. The relation is described by the parameter q_{IR} , the ratio between the total infrared luminosity (8–1000 μm ; rest frame) and the 1.4 GHz rest-frame luminosity, defined by Condon, Anderson & Helou (1991) as

$$q_{\text{IR}} = \log_{10} \left(\frac{L_{\text{IR}}}{3.75 \times 10^{12} L_{1.4\text{GHz}}} \right). \quad (5)$$

We explore the radio–infrared correlation for our sample to evaluate the contributions from star formation to the radio and FIR emission; for example, those sources above the correlation would have an excess of non-thermal synchrotron emission and those below the correlation would have an excess of thermal dust emission in the FIR, both of which could be related to the presence of a significant AGN contribution to the respective wavelength regimes (Sopp & Alexander 1991). Ivison et al. (2010) find $q_{\text{IR}} = 2.40 \pm 0.24$ for a flux-limited sample of sources selected from *Herschel*/SPIRE at 250 μm with VLA flux densities at 1.4 GHz. We plot the rest-frame radio and IR-luminosity for the sample in Fig. 14 and the median q_{IR} from Ivison et al. for reference. We interpolate or extrapolate to the rest frame 1.4 GHz (depending on the low-frequency data available) by fitting the radio SEDs with a power law, as given by equation (1). A spectral index of $\alpha = -0.70 \pm 0.14$ is assumed for those objects with a single radio measurement, which is typically at 1.4 GHz from NVSS or FIRST. As above, we do not account for magnification to prevent bias due to our assumptions about the magnification factor of the dust emission. The q_{IR} values for the radio-detected quasars are shown in Fig. 15.

We find that almost all of the jetted quasars lie significantly above the radio–infrared correlation for star-forming galaxies, by up to 3–

4 orders of magnitude in rest frame 1.4 GHz luminosity, and thus have q_{IR} values below that obtained by Ivison et al. (2010). This is to be expected as these are all powerful radio sources that are known to be dominated by synchrotron emission associated with AGN activity; the core and jet components of many sources are well studied as part of the CLASS, MG, and PMN gravitational lens surveys. We note that the average magnification factors of jetted radio sources will likely be higher than that of the dust, as the radio emission comes from a more compact region, although how much higher will be dependent on where the radio source lies relative to the lens and the lensing caustics. In such cases, the inferred q_{IR} values may be lower if the radio component is boosted relative to the dust heated by star formation. However, we do not expect this to alter our conclusions as the effect will only be significant for sources with extremely compact radio emission associated with jets and, in almost all cases, this will not produce the several orders of magnitude difference needed to account for the offset from the correlation seen in Fig. 15.

Quasars whose radio emission is associated with star formation are not expected to have significantly different magnifications between the radio and FIR, and so should remain close to the radio–infrared correlation for star-forming galaxies. Only 19 quasars in our sample classified as non-jetted have radio detections, of which only 2 are confirmed to have radio emission that is dominated by star formation (IRAS F10214+5255 and RX J1131–1231), the rest are currently undetermined. We observe a scatter around the radio–infrared relation for the non-jetted quasars. The scatter above the Ivison et al. (2010) relation may be due to contributions to their radio emission from low-power radio jets, or possibly additional radio emission from the foreground lensing galaxy. As the radio components of these sources have not yet been observed at a high enough angular resolution, it is not clear whether the apparent radio power dichotomy represents a true bi-modality in emission mechanism.

Recent studies point towards synchrotron emission from star formation as the dominant source of radio emission in non-jetted quasars (Padovani 2016, for review). However, evidence of a milliarcsec-scale jet in the classically radio-‘quiet’-lensed quasar HS 0810+2554 (part of our sample) suggests that it is not correct to assume that star formation is the primary radio emission mechanism in all cases (Hartley et al., in preparation). There may be a composite of emission processes, or a further subpopulation of quasars with low-power radio jets. Most of the detected radio-‘quiet’ quasars in our sample have q_{IR} values around the Ivison et al. (2010) relation, within the expected scatter. Deeper, higher spatial resolution observations of these sources are required to determine whether they hold to the relation, indicating whether the radio and FIR emission are indeed dominated by star formation. The radio upper limits in Fig. 14, due to non-detections in FIRST and NVSS, indicate that this population of quasars would be found in the $\sim \mu\text{Jy}$ regime as proposed by White et al. (2007), if star formation dominates.

Table 5. A summary of the objects used for our statistics. We give the median, 25th and 75th percentiles of certain properties, the number of measurements (N) this includes (number of limits, in brackets, where included), relevant figures within the paper, and an explanation of which objects are included in the selection. We give magnification-corrected luminosities here, but the selection is the same for the uncorrected luminosities and SFRs.

Property	Median	$N(\text{lims})$	Fig.	Comment
z	$1.8^{+0.7}_{-0.3}$	94	1	Objects with known redshift.
T_{dust}	38^{+12}_{-5} K	53	6, 10	Objects with known redshift, and dust temperature fitted as a free parameter in the SED.
β	$2.0^{+0.4}_{-0.5}$	21	11	Objects with β fitted as a free parameter in the SED. 18 of these have known redshift.
L_{FIR}	$3.6^{+4.8}_{-2.4} \times 10^{11} L_{\odot}$	63(31)	9, 12	Objects with known redshifts, of which 63 have measured dust emission and 31 have upper limits.

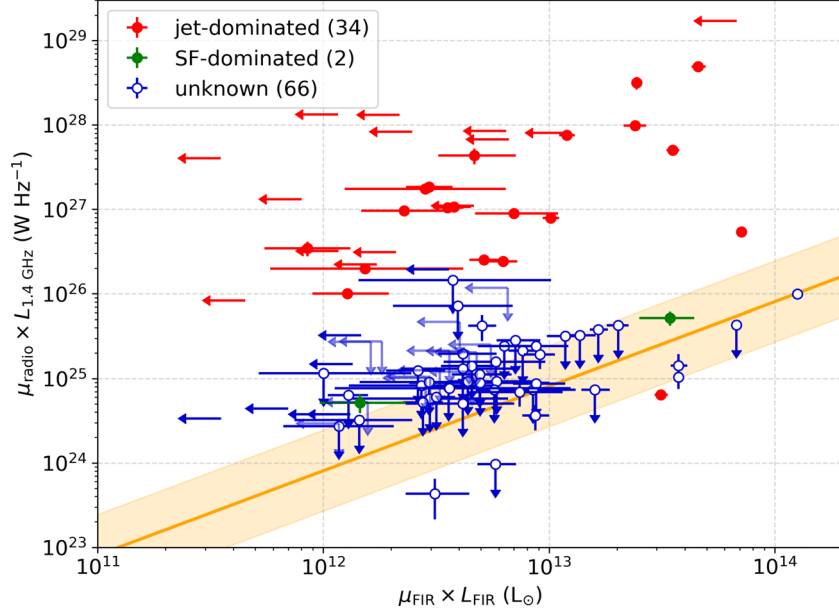


Figure 14. The radio–infrared correlation for 102 quasars in our sample, excluding WFI J2026–4536 and WFI J2033–4723 for which there are no radio data available. The median q_{IR} for star-forming galaxies from Ivison et al. (2010) is shown in yellow; the shaded region is $2\sigma_{q_{\text{IR}}}$.

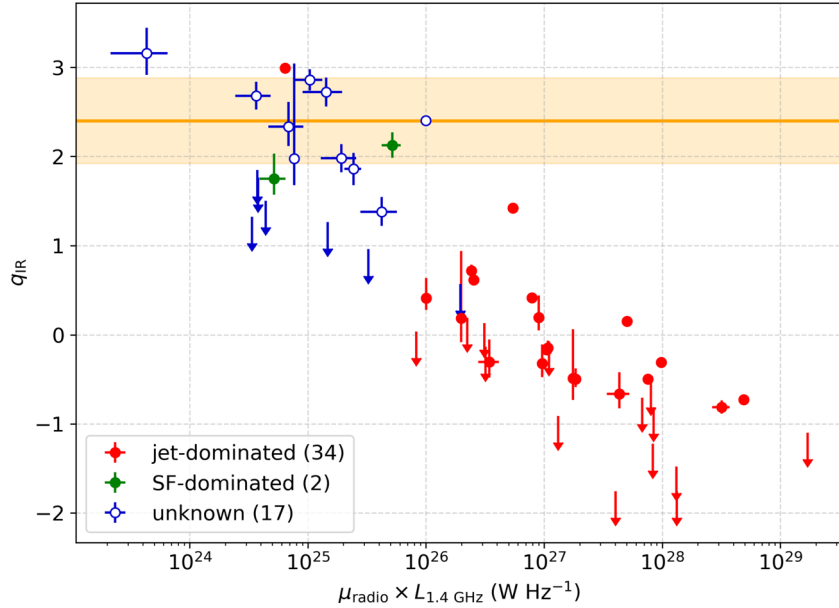


Figure 15. The radio–infrared factor, q_{IR} , for 53 quasars in the sample with radio detections. The median q_{IR} for star-forming galaxies from Ivison et al. (2010) is shown in yellow; the shaded region is $2\sigma_{q_{\text{IR}}}$.

We add 11 of the non-jetted quasars with radio detections more than 2σ above the radio–infrared correlation to our jetted subsample to refine our subsamples,⁸ under the assumption that the radio excess is due to AGN activity within the background object and not from the foreground lensing galaxy.

We again perform the K–S test on the FIR luminosity distribution of the samples with measured L_{FIR} and radio emission, as before, and find the probability that they are drawn from the same sample decreases from 0.23 to 0.06, but is still not statistically significant.

⁸ We exclude WFI 2026–4536 and WFI 2033–4723 in our radio–infrared correlation analysis, as there are no radio data for these sources.

The K–M-estimated median is higher for the objects which do not have a radio excess: $1.3^{+7.8}_{-1.2} \times 10^{12} L_{\odot}$ for the new jetted subsample and $4.1^{+3.3}_{-1.4} \times 10^{12} L_{\odot}$ for the remaining sources (as above, the uncertainties are at the 25th and 75th percentiles). It is possible that different FIR properties are simply caused by the small number of detections in the jetted subsample.

Notably, HS 0810+2554 lies below the radio–infrared relation with a q_{IR} value of 2.90 despite having radio jets that dominate its radio emission. In Section C1 of the Appendix C, we explore the possibility that this is caused by fitting a single temperature dust model to a composite of AGN and star-formation-heated dust. It is possible this is also the case for another seven

sources with larger fitted dust temperatures ($T_d > 60$ K) that appear to be outliers from the bulk of the sample (see Fig. 13) and would explain some of the observed anticorrelation between β and T_d (as shown in Fig. 11). Interestingly, when we apply a two-temperature model to HS 0810+2554 (described in Section C1), the corrected q_{IR} lies within the 2σ range of the radio–infrared correlation amidst the non-jetted sources. This suggests that the correlation may not be a reliable method of isolating jet-dominated quasars in the case of radio-weak AGN. If indeed quasars with radio jets are misidentified as non-jetted, this may cause or mask differences in the FIR properties between the subsamples.

The fact that no sources lie below the expected q_{IR} range in Fig. 14 implies that we do not significantly overestimate the FIR luminosity due to an additional un-modelled AGN component in the dust emission, with the exception of HS 0810+2554. We note that APM 08279+5255, which has the most well-defined SED and is fitted here with a two-temperature dust model, falls exactly on the radio–infrared correlation with a q_{IR} value of 2.40. Therefore, with sufficient data in the MIR and submillimetre it will be possible to better isolate the star-forming contribution to the SEDs. However, the radio and FIR luminosities derived for our current *Herschel*/SPIRE data set, like in the case of the fitted dust temperatures, appear to be consistent with star formation being the dominant mechanism for heating the dust.

5 CONCLUSIONS

According to the current paradigm of galaxy evolution, some fraction of the quasar population is expected to be transitional sources from DSFGs to UV-luminous quasars. These transition sources will be FIR-luminous, with clear evidence of ongoing star formation. However, only a handful of extremely FIR-luminous sources have been studied thus far due to the limitations in observational sensitivity, and then at wavelengths relatively insensitive to T_d and L_{FIR} . In order to study the link between DSFGs and quasars, we observed 104 gravitationally lensed quasars with the *Herschel*/SPIRE instrument at 250, 350, and 500 μm to determine the fraction that are FIR-luminous. Due to the magnification effects of gravitational lensing, we are able to measure the star-forming properties of individual quasars at lower intrinsic luminosities than those previously studied and thus more representative of the quasar population. We find most sources in our sample have magnification-corrected FIR luminosities below the estimated detection limit, a factor of 10 lower on average.

From our study, we detected 72 (69 per cent) of the gravitationally lensed quasars in at least one band with *Herschel*/SPIRE and find evidence for heated dust emission in 69 (66 per cent) of the objects, given the shape of their SEDs. By fitting modified black bodies to our new measurements and to data in the literature, we derive a median $\mu_{\text{FIR}} \times L_{\text{FIR}}$ of $3.6^{+4.7}_{-2.4} \times 10^{12} L_{\odot}$ and an implied magnification-corrected L_{FIR} of $3.6^{+4.8}_{-2.4} \times 10^{11} L_{\odot}$ and SFR of $120^{+160}_{-80} M_{\odot} \text{ yr}^{-1}$ for 94 objects with redshifts, under the assumption that for the vast majority of the targets the FIR magnification is $\mu_{\text{FIR}} = 10^{+10}_{-5}$. The range of fitted dust temperatures of the sample is 38^{+12}_{-5} K for 53 objects with redshifts and sufficiently constrained SEDs to fit for the dust temperature, which is characteristic of ongoing dust-obscured star formation. We compare our sample of gravitationally lensed quasars to a sample of DSFGs observed with *Herschel*/SPIRE at similar redshifts and find similar dust temperatures, which gives circumstantial evidence of star formation being the dominant mechanism for heating the dust. We find ~ 10 per cent have similar SFRs to DSFGs, suggestive of sources transitioning

from DSFGs to UV-luminous quasar, and a large range of SFRs $< 20\text{--}10\,000 M_{\odot} \text{ yr}^{-1}$.

Finally, using the radio–infrared correlation for star-forming galaxies, we find that the jetted quasars, selected by measured jet emission based on high resolution radio data, show an excess of radio luminosity by of up to four orders of magnitude. Non-jetted quasars lie close to the correlation expected for star-forming galaxies; however, the scatter above the correlation suggests at least some of these sources have contributions to their radio emission from AGN activity. We do not find evidence for an excess of FIR emission in the sample, given their radio luminosities, which again supports the view that the dust heating is dominated by star formation.

We find no significant difference in the L_{FIR} distribution of our jetted and non-jetted subsamples. Non-jetted quasars are three times as luminous on average when we select jetted sources by their radio excess relative to the radio–infrared correlation; however, the difference is not statistically significant. We caution that further data are required to eliminate the possibility of a systematic bias due to stellar mass or small sample sizes. Radio investigations are also required to identify the radio emission mechanism (and hence AGN feedback mechanisms) associated with non-jetted quasars, as our results suggest that the radio–infrared correlation may not be sufficient to identify sources with jet-dominated radio emission.

Our results reveal a strong co-existence of AGN activity and host galaxy star formation in quasars, as proposed in the Sanders et al. (1988) model. We find that this is true for the majority of quasars, suggesting that the FIR-luminous quasar phase is not distinct from the unobscured quasar phase and there is no sharp transition between them. However, our analysis is limited by the available observational data, which prevents us from making definitive statements about the implications for AGN and stellar feedback models. More photometric data points in the millimetre and submillimetre are needed to reduce the errors and assumptions in SED fitting, and thus the derived host galaxy properties. With additional MIR data, we can fit the AGN contribution to the dust heating and derive better constraints on the SFR for these sources. Much progress can also be made with the advent of ALMA, where the extent of the heated dust emission can be mapped on 50–500 mas-scales, which given the magnifications provided by the gravitational lenses, will also allow structures < 50 pc in size to be resolved. In addition, through high-resolution imaging of the millimetre emission from these sources, it will also be possible to determine robust wavelength-dependent magnifications from lens modelling, which will further reduce the uncertainties in our analysis for individual objects.

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REFERENCES

- Agudo I., Thum C., Wiesemeyer H., Krichbaum T. P., 2010, *ApJS*, 189, 1
- Agudo I., Thum C., Gómez J. L., Wiesemeyer H., 2014, *A&A*, 566, A59
- Alexander D. M., Hickox R. C., 2012, *New A Rev.*, 56, 93
- Alloin D., Kneib J.-P., Guilloteau S., Bremer M., 2007, *A&A*, 470, 53
- ALMA Partnership et al., 2015, *ApJ*, 808, L4
- Andreani P., Franceschini A., Granato G., 1999, *MNRAS*, 306, 161
- Augusto P. et al., 2001, *MNRAS*, 326, 1007
- Azadi M. et al., 2015, *ApJ*, 806, 187
- Banerji M., Carilli C. L., Jones G., Wagg J., McMahon R. G., Hewett P. C., Alaghband-Zadeh S., Feruglio C., 2017, *MNRAS*, 465, 4390
- Barvainis R., Ivison R., 2002, *ApJ*, 571, 712
- Barvainis R., Lonsdale C., 1997, *AJ*, 113, 144
- Barvainis R., Antonucci R., Coleman P., 1992, *ApJ*, 399, L19
- Barvainis R., Antonucci R., Hurt T., Coleman P., Reuter H.-P., 1995, *ApJ*, 451, L9
- Barvainis R., Alloin D., Bremer M., 2002, *A&A*, 385, 399
- Becker R. H., White R. L., Edwards A. L., 1991, *ApJS*, 75, 1
- Becker R. H., White R. L., Helfand D. J., 1995, *ApJ*, 450, 559
- Beelen A., Cox P., Benford D. J., Dowell C. D., Kovács A., Bertoldi F., Omont A., Carilli C. L., 2006, *ApJ*, 642, 694
- Bendo G. J. et al., 2013, *MNRAS*, 433, 3062
- Benford D. J., Cox P., Omont A., Phillips T. G., McMahon R. G., 1999, *ApJ*, 518, L65
- Bennett C. L., Lawrence C. R., Burke B. F., Hewitt J. N., Mahoney J., 1986, *ApJS*, 61, 1
- Bicknell G. V., Sutherland R. S., van Breugel W. J. M., Dopita M. A., Dey A., Miley G. K., 2000, *ApJ*, 540, 678
- Biggs A. D. et al., 2003, *MNRAS*, 338, 1084
- Bower R. G., Benson A. J., Malbon R., Helly J. C., Frenk C. S., Baugh C. M., Cole S., Lacey C. G., 2006, *MNRAS*, 370, 645
- Browne I. W. A. et al., 2003, *MNRAS*, 341, 13
- Bussmann R. S. et al., 2013, *ApJ*, 779, 25
- Carilli C. L., Bertoldi F., Omont A., Cox P., McMahon R. G., Isaak K. G., 2001, *AJ*, 122, 1679
- Carilli C. L., Lewis G. F., Djorgovski S. G., Mahabal A., Cox P., Bertoldi F., Omont A., 2003, *Science*, 300, 773
- Carilli C. L. et al., 2005, *ApJ*, 618, 586
- Castangia P. et al., 2011, *A&A*, 529, A150
- Chapman S. C., Scott D., Lewisi G. F., Borys C., Fahlman G. G., 1999, *A&A*, 352, 406
- Condon J. J., Condon M. A., Broderick J. J., Davis M. M., 1983, *AJ*, 88, 20
- Condon J. J., Anderson M. L., Helou G., 1991, *ApJ*, 376, 95
- Condon J. J., Cotton W. D., Greisen E. W., Yin Q. F., Perley R. A., Taylor G. B., Broderick J. J., 1998, *AJ*, 115, 1693
- Coppin K., Halpern M., Scott D., Borys C., Chapman S., 2005, *MNRAS*, 357, 1022
- Coppin K. E. K. et al., 2008, *MNRAS*, 389, 45
- Cox P. et al., 2002, *A&A*, 387, 406
- Dale D. A., Helou G., Contursi A., Silbermann N. A., Kolhatkar S., 2001, *ApJ*, 549, 215
- Davies M. L. et al., 2009, *MNRAS*, 400, 984
- Deane R. P., Rawlings S., Garrett M. A., Heywood I., Jarvis M. J., Klöckner H.-R., Marshall P. J., McKean J. P., 2013, *MNRAS*, 434, 3322
- Di Matteo T., Springel V., Hernquist L., 2005, *Nature*, 433, 604
- Downes D., Neri R., Wiklind T., Wilner D. J., Shaver P. A., 1999, *ApJ*, 513, L1
- Dye S. et al., 2018, *MNRAS*, 476, 4383
- Ellingsen S. P., Voronkov M. A., Breen S. L., Lovell J. E. J., 2012, *ApJ*, 747, L7
- Falco E. E., Lehar J., Perley R. A., Wambsganss J., Gorenstein M. V., 1996, *AJ*, 112, 897
- Foreman-Mackey D., Hogg D. W., Lang D., Goodman J., 2013, *PASP*, 125, 306
- García-Burillo S. et al., 2006, *ApJ*, 645, L17
- Garrett M. A., Calder R. J., Porcas R. W., King L. J., Walsh D., Wilkinson P. N., 1994, *MNRAS*, 270, 457
- Giommi P. et al., 2012, *A&A*, 541, A160
- Gregory P. C., Condon J. J., 1991, *ApJS*, 75, 1011
- Griffin M. J. et al., 2010, *A&A*, 518, L3
- Griffith M. R., Wright A. E., 1993, *AJ*, 105, 1666
- Griffith M., Langston G., Heflin M., Conner S., Lehar J., Burke B., 1990, *ApJS*, 74, 129
- Griffith M. R., Wright A. E., Burke B. F., Ekers R. D., 1994, *ApJS*, 90, 179
- Griffith M. R., Wright A. E., Burke B. F., Ekers R. D., 1995, *ApJS*, 97, 347
- Hainline L. J., Scoville N. Z., Yun M. S., Hawkins D. W., Frayer D. T., Isaak K. G., 2004, *ApJ*, 609, 61
- Harris K. et al., 2016, *MNRAS*, 457, 4179
- Harrison C. M. et al., 2016, *MNRAS*, 457, L122
- Helou G., Khan I. R., Malek L., Boehmer L., 1988, *ApJS*, 68, 151
- Henkel C., Braatz J. A., Menten K. M., Ott J., 2008, *A&A*, 485, 451
- Hewitt J. N., Turner E. L., Lawrence C. R., Schneider D. P., Brody J. P., 1992, *AJ*, 104, 968
- Hopkins P. F., Hernquist L., Cox T. J., Di Matteo T., Martini P., Robertson B., Springel V., 2005, *ApJ*, 630, 705
- Hopkins P. F., Hernquist L., Cox T. J., Kereš D., 2008, *ApJS*, 175, 356
- Hughes D. H., Dunlop J. S., Rawlings S., 1997, *MNRAS*, 289, 766
- Impellizzeri C. M. V., McKean J. P., Castangia P., Roy A. L., Henkel C., Brunthaler A., Wucknitz O., 2008, *Nature*, 456, 927
- Inada N. et al., 2003, *AJ*, 126, 666
- Inada N. et al., 2006, *AJ*, 131, 1934
- Inada N. et al., 2007, *AJ*, 133, 206
- Inada N. et al., 2009, *AJ*, 137, 4118
- Inada N. et al., 2012, *AJ*, 143, 119
- Irwin M. J., Ibatá R. A., Lewis G. F., Totten E. J., 1998, *ApJ*, 505, 529
- Isaak K. G., Priddey R. S., McMahon R. G., Omont A., Peroux C., Sharp R. G., Withington S., 2002, *MNRAS*, 329, 149
- Ivison R. J. et al., 2002, *MNRAS*, 337, 1
- Ivison R. J. et al., 2010, *A&A*, 518, L35
- Jackson N. et al., 1995, *MNRAS*, 274, L25
- Jackson N., Ofek E. O., Oguri M., 2008, *MNRAS*, 387, 741
- Jackson N., Ofek E. O., Oguri M., 2009, *MNRAS*, 398, 1423
- Jackson N., Tagore A. S., Roberts C., Sluse D., Stacey H., Vives-Arias H., Wucknitz O., Volino F., 2015, *MNRAS*, 454, 287
- Jethava N., Henkel C., Menten K. M., Carilli C. L., Reid M. J., 2007, *A&A*, 472, 435
- Johnston D. E. et al., 2003, *AJ*, 126, 2281
- Juvela M., Ysard N., 2012a, *A&A*, 539, A71
- Juvela M., Ysard N., 2012b, *A&A*, 541, A33

- Kalfountzou E. et al., 2014, *MNRAS*, 442, 1181
- Kayo I., Inada N., Oguri M., Morokuma T., Hall P. B., Kochanek C. S., Schneider D. P., 2010, *AJ*, 139, 1614
- Kellermann K. I., Pauliny-Toth I. I. K., 1973, *AJ*, 78, 828
- Kennicutt R. C., Jr, 1998, *ApJ*, 498, 541
- King L. J., Browne I. W. A., Muxlow T. W. B., Narasimha D., Patnaik A. R., Porcas R. W., Wilkinson P. N., 1997, *MNRAS*, 289, 450
- Klamer I. J., Ekers R. D., Sadler E. M., Hunstead R. W., 2004, *ApJ*, 612, L97
- Kochanek C. S., Falco E. E., Impey C., Lehar J., McLeod B., Rix H.-W., 1999, The CASTLES Survey, available at: <https://www.cfa.harvard.edu/castles/>
- Koopmans L. V. E. et al., 1999, *MNRAS*, 303, 727
- Kratzer R. M., Richards G. T., Goldberg D. M., Oguri M., Kochanek C. S., Hodge J. A., Becker R. H., Inada N., 2011, *ApJ*, 728, L18
- Krips M., Neri R., Eckart A., Downes D., Martín-Pintado J., Planesas P., 2005, *A&A*, 431, 879
- Krips M., Peck A. B., Sakamoto K., Petitpas G. B., Wilner D. J., Matsushita S., Iono D., 2007, *ApJ*, 671, L5
- Kühr H., Witzel A., Pauliny-Toth I. I. K., Nauber U., 1981, *A&AS*, 45, 367
- Lagattuta D. J., Vegetti S., Fassnacht C. D., Auger M. W., Koopmans L. V. E., McKean J. P., 2012, *MNRAS*, 424, 2800
- Langston G. I., Heflin M. B., Conner S. R., Lehar J., Carilli C. L., Burke B. F., 1990, *ApJS*, 72, 621
- Larionov M. G., Parijskij Y. N., Zhuravlev V. I., Sidorenkov V. N., Berlin A. B., Nizhel'skii N. A., 1994, *A&AS*, 106, 119
- Lawrence A. et al., 1993, *MNRAS*, 260, 28
- Lawrence C. R., Elston R., Januzzi B. T., Turner E. L., 1995, *AJ*, 110, 2570
- Lee J. A., Sohn B. W., Jung T., Byun D.-Y., Lee J. W., 2017, *ApJS*, 228, 22
- Lehar J. et al., 1997, *AJ*, 114, 48
- Leung T. K. D., Riechers D. A., Pavesi R., 2017, *ApJ*, 836, 180
- Lewis G. F., Chapman S. C., Ibata R. A., Irwin M. J., Totten E. J., 1998, *ApJ*, 505, L1
- Lis D. C., Neufeld D. A., Phillips T. G., Gerin M., Neri R., 2011, *ApJ*, 738, L6
- Lowe S. R., Gawroński M. P., Wilkinson P. N., Kus A. J., Browne I. W. A., Pazderski E., Feiler R., Kettle D., 2007, *A&A*, 474, 1093
- Magnelli B. et al., 2012, *A&A*, 539, A155 (M12)
- Mandelbaum R., Li C., Kauffmann G., White S. D. M., 2009, *MNRAS*, 393, 377
- Marlow D. R. et al., 2001, *AJ*, 121, 619
- Martí-Vidal I. et al., 2013, *A&A*, 558, A123
- Marton G. et al., 2017, preprint ([arXiv:1705.05693](https://arxiv.org/abs/1705.05693))
- Massardi M. et al., 2008, *MNRAS*, 384, 775
- Massardi M., López-Caniego M., González-Nuevo J., Herranz D., de Zotti G., Sanz J. L., 2009, *MNRAS*, 392, 733
- McKean J. P. et al., 2005, *MNRAS*, 356, 1009
- McKean J. P. et al., 2007, *MNRAS*, 378, 109
- Melbourne J. et al., 2012, *AJ*, 143, 125
- Morgan N. D., Dressler A., Maza J., Schechter P. L., Winn J. N., 1999, *AJ*, 118, 1444
- Moshir M. et al., 1990, *Bull. Am. Astron. Soc.*, 22, 1325
- Muller S., Guélin M., Dumke M., Lucas R., Combes F., 2006, *A&A*, 458, 417
- Murphy T. et al., 2010, *MNRAS*, 402, 2403
- Myers S. T. et al., 1995, *ApJ*, 447, L5
- Myers S. T. et al., 1999, *AJ*, 117, 2565
- Myers S. T. et al., 2003, *MNRAS*, 341, 1
- Nieppola E., Tornikoski M., Valtaja E., León-Tavares J., Hovatta T., Lähteenmäki A., Tammi J., 2011, *A&A*, 535, A69
- Ofek E. O., Oguri M., Jackson N., Inada N., Kayo I., 2007, *MNRAS*, 382, 412
- Oguri M. et al., 2008a, *AJ*, 135, 520
- Oguri M., Inada N., Blackburne J. A., Shin M.-S., Kayo I., Strauss M. A., Schneider D. P., York D. G., 2008b, *MNRAS*, 391, 1973
- Omout A., Cox P., Bertoldi F., McMahon R. G., Carilli C., Isaak K. G., 2001, *A&A*, 374, 371
- Omout A., Beelen A., Bertoldi F., Cox P., Carilli C. L., Priddey R. S., McMahon R. G., Isaak K. G., 2003, *A&A*, 398, 857
- Ott S., 2010, in Mizumoto Y., Morita K.-I., Ohishi M., eds, ASP Conf. Ser. Vol. 434, Astronomical Data Analysis Software and Systems XIX. Astron. Soc. Pac., San Francisco, p. 139
- Padovani P., 2016, *A&A Rev.*, 24, 13
- Page M. J., Stevens J. A., Ivison R. J., Carrera F. J., 2004, *ApJ*, 611, L85
- Paraficz D. et al., 2017, *A&A*, preprint ([arXiv:1705.09931](https://arxiv.org/abs/1705.09931))
- Patnaik A. R., Browne I. W. A., Wilkinson P. N., Wrobel J. M., 1992, *MNRAS*, 254, 655
- Pilbratt G. L. et al., 2010, *A&A*, 518, L1
- Pitchford L. K. et al., 2016, *MNRAS*, 462, 4067
- Planck Collaboration XIII, 2016, *A&A*, 594, A13
- Planesas P., Martín-Pintado J., Neri R., Colina L., 1999, *Science*, 286, 2493
- Priddey R. S., McMahon R. G., 2001, *MNRAS*, 324, L17
- Priddey R. S., Isaak K. G., McMahon R. G., Omout A., 2003, *MNRAS*, 339, 1183
- Richards J. L. et al., 2011, *ApJS*, 194, 29
- Riechers D. A., 2011, *ApJ*, 730, 108
- Riechers D. A., Walter F., Carilli C. L., Lewis G. F., 2009, *ApJ*, 690, 463
- Riechers D. A., Weiß A., Walter F., Wagg J., 2010, *ApJ*, 725, 1032
- Riechers D. A. et al., 2011, *ApJ*, 739, L32
- Rosario D. J. et al., 2013, *A&A*, 560, A72
- Rowan-Robinson M., 2000, *MNRAS*, 316, 885
- Rowan-Robinson M. et al., 1993, *MNRAS*, 261, 513
- Rusin D. et al., 2001, *AJ*, 122, 591
- Rybak M., McKean J. P., Vegetti S., Andreani P., White S. D. M., 2015a, *MNRAS*, 451, L40
- Rybak M., Vegetti S., McKean J. P., Andreani P., White S. D. M., 2015b, *MNRAS*, 453, L26
- Sanders D. B., Soifer B. T., Elias J. H., Madore B. F., Matthews K., Neugebauer G., Scoville N. Z., 1988, *ApJ*, 325, 74
- Savage R. S., Oliver S., 2007, *ApJ*, 661, 1339
- Scholtz J. et al., 2018, *MNRAS*, 475, 1288
- Simpson J. M. et al., 2012, *MNRAS*, 426, 3201
- Snellen I. A. G., de Bruyn A. G., Schilizzi R. T., Miley G. K., Myers S. T., 1995, *ApJ*, 447, L9
- Solomon P., Vanden Bout P., Carilli C., Guélin M., 2003, *Nature*, 426, 636
- Sopp H. M., Alexander P., 1991, *MNRAS*, 251, 14P
- Stacey H. R., 2015, Msc thesis, Univ. Manchester
- Stanley F. et al., 2017, *MNRAS*, 472, 2221
- Stevens J. A., Page M. J., Ivison R. J., Carrera F. J., Mittaz J. P. D., Smail I., McHardy I. M., 2005, *MNRAS*, 360, 610
- Swinbank A. M. et al., 2010, *Nature*, 464, 733
- Sykes C. M. et al., 1998, *MNRAS*, 301, 310
- Tingay S. J., Jauncey D. L., King E. A., Tzioumis A. K., Lovell J. E. J., Edwards P. G., 2003, *PASJ*, 55, 351
- Tinti S., Dallacasa D., de Zotti G., Celotti A., Stanghellini C., 2005, *A&A*, 432, 31
- Tuan-Anh P., Boone F., Hoai D. T., Nhung P. T., Weiß A., Kneib J.-P., Beelen A., Salomé P., 2013, *A&A*, 552, L12
- Tuan-Anh P., Hoai D. T., Nhung P. T., Diep P. N., Phuong N. T., Thao N. T., Darriulat P., 2017, *MNRAS*, 467, 3513
- van der Werf P. P. et al., 2011, *ApJ*, 741, L38
- Venturini S., Solomon P. M., 2003, *ApJ*, 590, 740
- Vollmer B., Krichbaum T. P., Angelakis E., Kovalev Y. Y., 2008, *A&A*, 489, 49
- Waldram E. M., Pooley G. G., Grainge K. J. B., Jones M. E., Saunders R. D. E., Scott P. F., Taylor A. C., 2003, *MNRAS*, 342, 915
- Weiß A., Henkel C., Downes D., Walter F., 2003, *A&A*, 409, L41
- Weiß A., Downes D., Neri R., Walter F., Henkel C., Wilner D. J., Wagg J., Wiklind T., 2007, *A&A*, 467, 955
- White R. L., Becker R. H., 1992, *ApJS*, 79, 331
- White R. L., Helfand D. J., Becker R. H., Glikman E., de Vries W., 2007, *ApJ*, 654, 99
- Winn J. N. et al., 2000, *AJ*, 120, 2868
- Winn J. N., Hewitt J. N., Patnaik A. R., Schechter P. L., Schommer R. A., López S., Maza J., Wachter S., 2001, *AJ*, 121, 1223
- Winn J. N. et al., 2002, *AJ*, 123, 10
- Winn J. N., Rusin D., Kochanek C. S., 2003, *ApJ*, 587, 80
- Winn J. N., Rusin D., Kochanek C. S., 2004, *Nature*, 427, 613

Wright A. E., Wark R. M., Troup E., Otrupcek R., Hunt A., Cooke D. J., 1990, *Proc. Astron. Soc. Aust.*, 8, 261
 Wright E. L. et al., 2009, *ApJS*, 180, 283
 Wu J., Vanden Bout P. A., Evans N. J., II, Dunham M. M., 2009, *ApJ*, 707, 988
 Wucknitz O., Volino F., 2008, in Proceedings of the 9th European VLBI Network Symposium on the Role of VLBI in the Golden Age for Radio Astronomy. Bologna, Italy, p. 102

Xanthopoulos E. et al., 1998, *MNRAS*, 300, 649
 Xanthopoulos E., Combes F., Wiklind T., 2001, *MNRAS*, 325, 273
 York T. et al., 2005, *MNRAS*, 361, 259
 Yun M. S., Carilli C. L., 2002, *ApJ*, 568, 88
 Zeiger B., Darling J., 2010, *ApJ*, 709, 386

APPENDIX A: SOURCE TABLES

Table A1. *Herschel*/SPIRE measurements of the lensed quasar sample. We give the lens name, whether the object is jetted (J) or non-jetted (N), the maximum separation of the lensed images ($\Delta\theta$), the source redshift (z_s), the measured flux densities at 250, 350, and 500 μm , and in brackets, the uncertainty from the source extraction. We comment on notable features of the lens system or with a † denote any issues with the flux-density measurements not discussed in Section 2.4. Redshifts and image separations are from the CASTLES catalogue (Kochanek et al. 1999), unless another reference is given.

Lens name	Type	$\Delta\theta$ (arcsec)	z_s	$S_{250\mu\text{m}}$ (mJy)	$S_{350\mu\text{m}}$ (mJy)	$S_{500\mu\text{m}}$ (mJy)	Comments	References
HE 0047–1756	N	1.44	1.676	197(9)	130(8)	60(9)		
CLASS B0128+437	J	0.55	3.114	<33	<33	<36		
Q J0158–4325	N	1.22	1.29	39(9)	38(8)	<27		
JVAS B0218+357	J	0.34	0.96	89(7)	122(7)	120(8)		
HE 0230–2130	N	2.05	2.162	126(9)	109(8)	77(9)		
SDSS J0246–0825	N	1.19	1.68	88(7)	75(7)	30(8)		
CFRS 03.1077	N	2.1	2.941	43(7)	44(7)	<37		
MG J0414+0534	J	2.4	2.64	266(7)	190(8)	112(10)		
HE 0435–1223	N	2.42	1.689	133(7)	101(7)	53(9)		
CLASS B0445+123	J	1.35	–	<38	<50	<42		
HE 0512–3329	N	0.65	1.57	60(7)	39(9)	<32		
CLASS B0631+519	J	1.16	–	63(12)	82(7)	71(12)		
CLASS B0712+472	J	1.46	1.34	<33	<28	<36		
CLASS B0739+366	J	0.53	–	69(6)	61(10)	<40		
SDSS J0746+4403	N	1.11	2.0	<30	<27	<27		Inada et al. (2007)
MG J0751+2716	J	0.7	3.21	102(4)	105(3)	78(4)		
SDSS J0806+2006	N	1.4	1.54	40(9)	<21	<26		Inada et al. (2006)
HS 0810+2554	N	0.96	1.5	186(9)	98(8)	53(10)†	†Possible blending	
HS 0818+1227	N	2.83	3.115	<30	<28	<30		
SDSS J0819+5356	N	4.04	2.239	40(9)	40(8)	<40		Inada et al. (2009)
SDSS J0820+0812	N	2.2	2.024	49(9)	54(8)	<32		Jackson, Ofek & Oguri (2009)
APM 08279+5255	N	0.38	3.91	621(3)	437(3)	259(4)		Downes et al. (1999)
SDSS J0832+0404	N	1.98	1.116	<30	<30	<32		Oguri et al. (2008a)
CLASS B0850+054	J	0.68	–	57(9)	45(7)	<32		
SDSS J0903+5028	N	2.99	3.605	227(11)	182(8)	<50		Johnston et al. (2003)
SDSS J0904+1512	N	1.13	1.826	27(7)	<30	<47		Kayo et al. (2010)
SBS J0909+523	J	1.17	1.38	<27	<30	<47		
RX J0911+0551	N	2.47	2.79	181(11)	176(9)	97(9)		
RX J0921+4529	N	6.97	1.65	<26	<30	<30		
SDSS J0924+0219	N	1.75	1.524	67(8)	56(8)	32(10)		Inada et al. (2003)
FBQS J0951+2635	J	1.11	1.24	<30	<30	<41		
Q 0957+561	J	6.26	1.41	108(10)	81(7)	<30		
SDSS J1001+5027	N	2.82	1.84	30(9)	<27	<30		
SDSS J1004+1229	J	1.54	2.65	<26	<26	<41		
LBQS J1009–0252	N	1.54	2.74	<25	<27	<27		
SDSS J1011+0143	N	3.67	2.701	<28	<30	<40		
Q 1017–207	N	0.85	2.55	<27	<30	<40		
SDSS J1021+4913	N	1.14	1.72	38(9)	28(8)	<30		
IRAS F10214+4724	J	1.59	2.29	416(5)	303(4)	169(4)		
SDSS J1029+2623	N	22.5	2.197	49(5)	47(6)	30(5)		
JVAS B1030+074	J	1.65	1.54	35(6)	44(9)	63(11)		
SDSS J1054+2733	N	1.27	1.452	88(9)	84(8)	55(9)		Kayo et al. (2010)
SDSS J1055+4628	N	1.15	1.249	<28	<30	<33		Kayo et al. (2010)
HE 1104–1805	N	3.19	2.32	139(7)	122(7)	88(10)		
PG 1115+080	N	2.32	1.72	69(17)	40(7)	<39		
CLASS B1127+385	J	0.74	–	90(9)	105(7)	56(10)		
RX J1131–1231	N	3.8	0.658	285(7)	166(8)	62(9)		
MG J1131+0456	J	2.1	–	<25	<32	<32		
SDSS J1131+1915	N	1.46	2.915	<28	<30	<30		Kayo et al. (2010)
SDSS J1138+0314	N	1.34	2.44	33(7)	<27	<38		

Table A1 – *continued*

Lens name	Type	$\Delta\theta$ (arcsec)	z_s	$S_{250\mu\text{m}}$ (mJy)	$S_{350\mu\text{m}}$ (mJy)	$S_{500\mu\text{m}}$ (mJy)	Comments	References
SDSS J1155+6346	N	1.95	2.89	29(7)	30(6)	<30		
CLASS B1152+200	J	1.59	1.019	37(10)	<30	<32		
SDSS J1206+4332	N	2.9	1.79	95(7)	70(8)	<31		
Q 1208+101	N	0.48	3.82	61(9)	52(7)	<36		
SDSS J1216+3529	N	1.49	2.013	<32	<32	<30		Oguri et al. (2008a)
HST 12531–2914	N	1.23	–	<28	<26	<32		
SDSS J1258+1657	N	1.28	2.702	46(7)	54(8)	60(13)		Inada et al. (2009)
SDSS J1304+2001	N	1.87	2.175	27(6)	26(7)	<30		Kayo et al. (2010)
SDSS J1313+5151	N	1.24	1.877	49(6)	<30	<36		Ofek et al. (2007)
SDSS J1322+1052	N	2.0	1.711	87(10)	79(7)	57(9)		Oguri et al. (2008a)
SDSS J1330+1810	N	1.76	1.393	126(7)	91(7)	36(11)		Oguri et al. (2008b)
SDSS J1332+0347	N	1.14	1.445	<27	<30	<32		
LBQS J1333+0113	N	1.63	1.57	90(9)	80(7)	55(10)		
SDSS J1339+1310	N	1.69	2.241	27(8)	<30	<32		Inada et al. (2009)
SDSS J1349+1227	N	3.0	1.722	59(8)	52(8)	32(9)		Kayo et al. (2010)
SDSS J1353+1138	N	1.41	1.629	69(9)	43(7)	<30		
Q 1355–2257	N	1.23	1.37	28(7)	30(8)	<32		
CLASS B1359+154	J	1.71	3.235	139(7)	99(7)	64(9)†	†Possible blending	
SDSS J1400+3134	N	1.74	3.317	45(7)	28(9)	<42		Inada et al. (2009)
SDSS J1402+6321	J	1.35	0.48	<27	<25	<30		
SDSS J1406+6126	N	1.98	2.13	<35	<36	<41		
HST 14113+5211	N	1.8	2.81	<28	<30	<34		
H 1413+117	J	1.35	2.55	521(3)	403(3)	247(4)		
JVAS B1422+231	J	1.68	3.62	36(11)	<30	<45		
SDSS J1455+1447	N	1.73	1.424	54(7)	34(7)	<30		Kayo et al. (2010)
SBS J1520+530	N	1.59	1.86	54(9)	43(8)	<33		
SDSS J1524+4409	N	1.67	1.21	43(10)	32(7)	<30		Oguri et al. (2008a)
HST 15433+5352	N	1.18	2.092	<30	<30	<47		
MG J1549+3047	J	1.7	1.17	<28	<32	<36	Lensed radio lobe	
CLASS B1555+375	J	0.42	–	<29	<30	<30		
CLASS B1600+434	J	1.4	1.59	36(6)	27(6)	<39	Star-forming lens galaxy	
CLASS B1608+656	J	2.27	1.39	33(6)	33(4)	28(5)		
SDSS J1620+1203	N	2.77	1.158	<26	<28	<40		Kayo et al. (2010)
PMN J1632–0033	J	1.47	3.424	<24	35(6)	56(9)		
FBQS J1633+3134	J	0.75	1.52	62(9)	30(8)	<30		
SDSS J1650+4251	N	1.23	1.54	86(9)	89(8)	53(9)		
MG J1654+1346	J	2.1	1.74	<34	<30	<43	Lensed radio lobe	
PKS J1830–211	J	0.99	2.51	537(9)	670(9)	806(11)		
PMN J1838–3427	J	0.99	2.78	65(5)	90(6)	86(8)		
CLASS B1933+503	J	1.0	1.71	243(7)	212(8)	125(10)		
JVAS B1938+666	J	1.0	2.01	163(7)	164(7)	120(10)		
PMN J2004–1349	J	1.18	–	38(7)†	49(6)†	28(7)†	†Diffuse galactic emission	
MG J2016+112	J	3.52	3.27	81(7)	48(6)	<32		
WFI J2026–4536	N	1.34	2.23	460(7)	301(7)	162(10)		
WFI J2033–4723	N	2.33	1.66	100(7)	84(7)	50(11)		
CLASS B2045+265	J	2.74	2.35*	<36	<33	<43		*C. Fassnacht, priv. comm.
CLASS B2108+213	J	4.57	0.67	<32	<32	<36		
JVAS B2114+022	J	1.31	0.59	137(7)	108(9)	36(9)	Star-forming lens galaxy	
HE 2149–2745	N	1.7	2.03	84(7)	77(9)	37(9)		
CY 2201–3201	N	0.83	3.9	<37	<30	<44		
Q 2237+030	N	1.78	1.69	<30	<28	<40		
CLASS B2319+052	J	1.36	–	56(6)	57(7)	<33		
PSS J2322+1944	N	1.49	4.12	83(4)	105(3)	81(5)		
SDSS J2343–0050	N	1.51	0.787	<26	<30	<34		Jackson, Ofek & Oguri (2008)

Table A2. We give the FIR luminosities (40–120 μm) and SFRs of the quasar sample, derived from both the least-squares SED fit and the median, 16th and 84th percentiles from the MCMC analysis. The values are not corrected for lensing magnification (μ_{FIR}). Where objects are fit with two-temperature dust models, both temperatures are reported but the lower temperature is used for estimation of L_{FIR} (i.e. luminosity due to star formation) and SFR. Square brackets denote temperature fits where the source is synchrotron-dominated (noted in the comments), which are used only to derive upper limits on the dust emission.

Lens Name	T_{dust} (K)	Least squares			T_{dust} (K)	MCMC			Comments
		β	$\log \mu L_{\text{FIR}}$ (L_{\odot})	$\log \mu \text{SFR}$ ($M_{\odot} \text{yr}^{-1}$)		β	$\log \mu L_{\text{FIR}}$ (L_{\odot})	$\log \mu \text{SFR}$ ($M_{\odot} \text{yr}^{-1}$)	
HE 0047–1756	42.6	–	13.2	3.7	43.2 $^{+3.4}_{-2.8}$	–	13.2 $^{+0.1}_{-0.1}$	3.7 $^{+0.1}_{-0.1}$	
CLASS B0128+437	–	–	<12.8	<3.3	–	–	–	–	
Q J0158–4325	27.7	–	12.0	2.5	30.4 $^{+7.9}_{-5.7}$	–	12.1 $^{+0.3}_{-0.2}$	2.6 $^{+0.3}_{-0.2}$	
JVAS B0218+357	[16.8]	–	<11.5	<2.1	–	–	–	–	synchrotron dominated
HE 0230–2130	39.4	1.3	13.1	3.6	41.8 $^{+9.0}_{-6.3}$	1.2 $^{+0.4}_{-0.4}$	13.1 $^{+0.1}_{-0.1}$	3.65 $^{+0.08}_{-0.06}$	
SDSS J0246–0825	36.0	–	12.7	3.3	36.8 $^{+3.9}_{-3.1}$	–	12.8 $^{+0.1}_{-0.1}$	3.3 $^{+0.1}_{-0.1}$	
CFRS03.1077	42.7	–	13.0	3.5	44.7 $^{+8.0}_{-6.1}$	–	13.0 $^{+0.1}_{-0.1}$	3.48 $^{+0.07}_{-0.07}$	
MG J0414+0534	41.2	2.3	13.7	4.2	42.8 $^{+7.4}_{-4.8}$	2.2 $^{+0.4}_{-0.4}$	13.66 $^{+0.04}_{-0.03}$	4.18 $^{+0.04}_{-0.03}$	
HE 0435–1223	37.0	–	12.9	3.4	37.3 $^{+2.7}_{-2.3}$	–	12.9 $^{+0.1}_{-0.1}$	3.5 $^{+0.1}_{-0.1}$	
CLASS B0445+123	–	–	<12.1	<2.6	–	–	–	–	
HE 0512–3329	41.8	–	12.6	3.1	50.1 $^{+18.4}_{-12.2}$	–	12.8 $^{+0.3}_{-0.2}$	3.3 $^{+0.3}_{-0.2}$	
CLASS B0631+519	26.3	–	12.6	3.1	26.7 $^{+1.4}_{-1.3}$	–	12.6 $^{+0.1}_{-0.1}$	3.1 $^{+0.1}_{-0.1}$	
CLASS B0712+472	–	–	<12.2	<2.8	–	–	–	–	
CLASS B0739+366	25.7	2.9	12.7	3.2	27.0 $^{+6.2}_{-4.2}$	2.7 $^{+0.7}_{-0.7}$	12.7 $^{+0.1}_{-0.1}$	3.2 $^{+0.1}_{-0.1}$	
SDSS J0746+4403	–	–	<12.5	<3.0	–	–	–	–	
MG J0751+2716	36.2	2.4	13.4	3.9	36.2 $^{+1.9}_{-1.7}$	2.4 $^{+0.2}_{-0.2}$	13.39 $^{+0.01}_{-0.01}$	3.91 $^{+0.01}_{-0.01}$	
SDSS J0806+2006	–	–	12.4 $^{+0.4}_{-0.2}$	2.9 $^{+0.4}_{-0.2}$	–	–	–	–	
HS 0810+2554	89.1	1.0	13.5	4.0	89.0 $^{+6.5}_{-6.0}$	1.0 $^{+0.2}_{-0.2}$	13.50 $^{+0.03}_{-0.03}$	4.01 $^{+0.03}_{-0.03}$	
SDSS J0819+5356	36.2	–	12.6	3.2	38.1 $^{+7.5}_{-6.3}$	–	12.7 $^{+0.1}_{-0.1}$	3.2 $^{+0.1}_{-0.1}$	
SDSS J0820+0812	35.8	–	12.7	3.2	38.1 $^{+8.1}_{-5.2}$	–	12.7 $^{+0.2}_{-0.1}$	3.2 $^{+0.2}_{-0.1}$	
HS 0818+1227	–	–	12.6 $^{+0.4}_{-0.4}$	3.1 $^{+0.4}_{-0.4}$	–	–	–	–	
APM 08279+5255	56.9, 160.3	1.4	14.1	4.6	57.1 $^{+1.2}_{-1.2}$, 160.7 $^{+2.5}_{-2.5}$	1.43 $^{+0.03}_{-0.03}$	14.10 $^{+0.01}_{-0.01}$	4.62 $^{+0.01}_{-0.01}$	
SDSS J0832+0404	–	–	<12.1	<2.6	–	–	–	–	
CLASS B0850+054	42.4	–	12.8	3.3	48.8 $^{+17.6}_{-10.0}$	–	12.8 $^{+0.3}_{-0.1}$	3.4 $^{+0.3}_{-0.1}$	
SDSS J0903+5028	60.2	–	13.8	4.3	60.5 $^{+3.5}_{-3.1}$	–	13.83 $^{+0.02}_{-0.02}$	4.35 $^{+0.02}_{-0.02}$	
SDSS J0904+1512	–	–	12.4 $^{+0.2}_{-0.2}$	3.0 $^{+0.2}_{-0.2}$	–	–	–	–	
SBS 0909+523	–	–	<12.2	<2.7	–	–	–	–	
RX J0911+0551	51.0	1.3	13.6	4.1	52.6 $^{+7.8}_{-6.1}$	1.3 $^{+0.2}_{-0.2}$	13.58 $^{+0.04}_{-0.03}$	4.09 $^{+0.04}_{-0.03}$	
RX J0921+4529	–	–	<12.3	<2.8	–	–	–	–	
SDSS J0924+0219	31.8	–	12.5	3.0	33.4 $^{+5.9}_{-4.0}$	–	12.5 $^{+0.2}_{-0.1}$	3.0 $^{+0.2}_{-0.1}$	
FBQS 0951+2635	–	–	<12.1	<2.6	–	–	–	–	
Q 0957+561	27.9	2.3	12.6	3.1	31.1 $^{+11.1}_{-6.8}$	2.0 $^{+0.7}_{-0.6}$	12.7 $^{+0.2}_{-0.1}$	3.2 $^{+0.2}_{-0.1}$	
SDSS J1001+5027	–	–	12.5 $^{+0.3}_{-0.1}$	3.0 $^{+0.3}_{-0.1}$	–	–	–	–	
SDSS J1004+1229	–	–	<12.6	<3.1	–	–	–	–	
LBQS 1009–0252	–	–	<12.6	<3.2	–	–	–	–	
SDSS J1011+0143	–	–	<12.7	<3.2	–	–	–	–	
Q 1017–207	–	–	<12.6	<3.1	–	–	–	–	
SDSS J1021+4913	38.8	–	12.4	3.0	48.4 $^{+19.8}_{-13.4}$	–	12.6 $^{+0.3}_{-0.2}$	3.1 $^{+0.3}_{-0.2}$	
IRAS F10214+4724	40.6, 104.3	1.3	13.5	4.0	44.7 $^{+5.9}_{-5.9}$, 90.2 $^{+7.8}_{-6.0}$	1.2 $^{+0.2}_{-0.2}$	13.5 $^{+0.1}_{-0.1}$	4.1 $^{+0.1}_{-0.1}$	
SDSS J1029+2623	35.9	–	12.7	3.2	36.4 $^{+3.1}_{-2.5}$	–	12.7 $^{+0.1}_{-0.1}$	3.2 $^{+0.1}_{-0.1}$	
JVAS B1030+074	[20.2]	–	<11.9	<2.4	–	–	–	–	synchrotron dominated
SDSS J1054+2733	27.6	–	12.4	2.9	28.0 $^{+2.3}_{-2.0}$	–	12.4 $^{+0.1}_{-0.1}$	3.0 $^{+0.1}_{-0.1}$	
SDSS J1055+4628	–	–	<12.1	<2.6	–	–	–	–	
HE 1104–1805	35.5	1.9	13.2	3.7	36.1 $^{+4.6}_{-3.6}$	1.9 $^{+0.3}_{-0.3}$	13.22 $^{+0.04}_{-0.04}$	3.74 $^{+0.04}_{-0.04}$	
PG 1115+080	51.7	–	13.2	3.7	55.1 $^{+17.5}_{-8.9}$	–	12.9 $^{+0.2}_{-0.1}$	3.4 $^{+0.2}_{-0.1}$	
CLASS B1127+385	23.7	2.9	12.8	3.3	24.2 $^{+4.3}_{-3.1}$	2.8 $^{+0.6}_{-0.6}$	12.8 $^{+0.1}_{-0.1}$	3.3 $^{+0.1}_{-0.1}$	
RX J1131–1231	19.4	2.9	12.1	2.6	20.5 $^{+5.7}_{-4.3}$	2.7 $^{+1.0}_{-0.7}$	12.2 $^{+0.3}_{-0.1}$	2.7 $^{+0.3}_{-0.2}$	
MG J1131+0456	–	–	<12.3	<2.9	–	–	–	–	

Table A2 – *continued*

Lens Name	T_{dust} (K)	Least squares			T_{dust} (K)	MCMC			Comments
		β	$\log \mu L_{\text{FIR}}$ (L_{\odot})	$\log \mu \text{SFR}$ ($M_{\odot} \text{yr}^{-1}$)		β	$\log \mu L_{\text{FIR}}$ (L_{\odot})	$\log \mu \text{SFR}$ ($M_{\odot} \text{yr}^{-1}$)	
SDSS J1131+1915	–	–	<12.7	<3.2	–	–	–	–	
SDSS J1138+0314	–	–	$12.6^{+0.2}_{-0.1}$	$3.1^{+0.2}_{-0.1}$	–	–	–	–	
SDSS J1155+6346	43.0	–	12.8	3.3	$49.4^{+20.5}_{-10.5}$	–	$12.8^{+0.1}_{-0.1}$	$3.3^{+0.1}_{-0.1}$	
CLASS B1152+200	–	–	$12.2^{+0.8}_{-0.3}$	$2.7^{+0.8}_{-0.3}$	–	–	–	–	
SDSS J1206+4332	42.0	–	12.9	3.4	$44.2^{+8.5}_{-5.5}$	–	$12.9^{+0.2}_{-0.1}$	$3.5^{+0.2}_{-0.1}$	
Q 1208+101	68.9	1.2	13.3	3.8	$81.4^{+33.1}_{-23.5}$	$1.0^{+0.6}_{-0.4}$	$13.3^{+0.1}_{-0.1}$	$3.8^{+0.1}_{-0.1}$	
SDSS J1216+3529	–	–	<12.5	<3.0	–	–	–	–	
HST 12531–2914	–	–	<12.2	<2.7	–	–	–	–	
SDSS J1258+1657	33.3	–	12.9	3.4	$34.2^{+3.9}_{-3.2}$	–	$12.9^{+0.1}_{-0.1}$	$3.4^{+0.1}_{-0.1}$	
SDSS J1304+2001	35.5	–	12.4	2.9	$31.0^{+7.3}_{-4.6}$	–	$12.4^{+0.2}_{-0.1}$	$2.9^{+0.2}_{-0.1}$	
SDSS J1313+5151	–	–	$12.7^{+0.4}_{-0.2}$	$3.2^{+0.4}_{-0.2}$	–	–	–	–	
SDSS J1322+1052	30.4	–	12.6	3.2	$30.8^{+2.9}_{-2.4}$	–	$12.6^{+0.1}_{-0.1}$	$3.2^{+0.1}_{-0.1}$	
SDSS J1330+1810	35.6	–	12.7	3.3	$36.2^{+3.5}_{-2.7}$	–	$12.8^{+0.1}_{-0.1}$	$3.3^{+0.1}_{-0.1}$	
SDSS J1332+0347	–	–	<12.2	<2.7	–	–	–	–	
LBQS 1333+0133	20.1	–	12.1	2.6	$20.2^{+1.4}_{-1.3}$	–	$12.1^{+0.1}_{-0.1}$	$2.6^{+0.1}_{-0.1}$	
SDSS J1339+1310	–	–	$12.6^{+0.3}_{-0.2}$	$3.1^{+0.3}_{-0.2}$	–	–	–	–	
SDSS J1349+1227	32.9	–	12.5	3.0	$34.4^{+5.9}_{-4.1}$	–	$12.6^{+0.2}_{-0.1}$	$3.1^{+0.2}_{-0.1}$	
SDSS J1353+1138	45.7	–	12.8	3.3	$51.7^{+15.2}_{-11.0}$	–	$12.9^{+0.2}_{-0.2}$	$3.4^{+0.2}_{-0.2}$	
Q 1355–2257	25.7	–	11.8	2.3	$31.2^{+17.1}_{-7.6}$	–	$12.0^{+0.6}_{-0.2}$	$2.5^{+0.6}_{-0.2}$	
CLASS B1359+154	58.8	1.7	13.5	4.1	$60.3^{+10.8}_{-10.2}$	$1.7^{+0.5}_{-0.4}$	$13.55^{+0.03}_{-0.03}$	$4.06^{+0.03}_{-0.03}$	
SDSS J1400+3134	69.6	–	13.4	3.9	$75.0^{+18.7}_{-17.2}$	–	$13.1^{+0.1}_{-0.1}$	$3.6^{+0.1}_{-0.1}$	
SDSS J1402+6321	–	–	<11.5	<2.0	–	–	–	–	
SDSS J1406+6126	–	–	<12.5	<3.0	–	–	–	–	
HST 14113+52116	–	–	<12.6	<3.1	–	–	–	–	
H 1413+117	35.5, 124.1	2.4	13.85	4.37	$35.6^{+0.6}_{-0.6}$, $125.6^{+10.6}_{-8.9}$	$2.4^{+0.1}_{-0.1}$	$13.85^{+0.01}_{-0.01}$	$4.37^{+0.01}_{-0.01}$	
JVAS B1422+231	–	–	12.4	2.9	–	–	$12.36^{+0.40}_{-0.18}$	$2.87^{+0.40}_{-0.18}$	
SDSS J1455+1447	39.7	–	12.5	3.0	$46.6^{+14.9}_{-10.8}$	–	$12.6^{+0.3}_{-0.2}$	$3.1^{+0.3}_{-0.2}$	
SBS 1520+530	42.1	–	12.7	3.2	$46.2^{+12.7}_{-7.2}$	–	$12.8^{+0.2}_{-0.1}$	$3.3^{+0.2}_{-0.1}$	
SDSS J1524+4409	32.0	–	12.4	2.9	$34.6^{+8.1}_{-7.4}$	–	$12.2^{+0.3}_{-0.2}$	$2.7^{+0.3}_{-0.2}$	
HST 15433+5352	–	–	<12.7	<3.2	–	–	–	–	
MG J1549+3047	–	–	<12.1	<2.6	–	–	–	–	
CLASS B1555+375	–	–	<12.3	<2.8	–	–	–	–	
CLASS B1600+434	35.8	–	12.3	2.8	$38.9^{+8.9}_{-7.5}$	–	$12.4^{+0.2}_{-0.2}$	$2.9^{+0.2}_{-0.2}$	
CLASS B1608+656	31.9	1.8	11.9	2.4	$36.3^{+13.7}_{-9.2}$	$1.5^{+0.9}_{-0.7}$	$11.9^{+0.2}_{-0.1}$	$2.4^{+0.2}_{-0.1}$	
SDSS J1620+1203	–	–	<12.0	<2.6	–	–	–	–	
PMN J1632–0033	[26.8]	–	<12.8	<3.3	–	–	–	–	synchrotron dominated
FBQS 1633+3134	65.1	–	12.9	3.5	$66.2^{+14.6}_{-15.1}$	–	$12.9^{+0.2}_{-0.2}$	$3.5^{+0.2}_{-0.2}$	
SDSS J1650+4251	28.1	–	12.5	3.0	$28.3^{+2.3}_{-1.9}$	–	$12.5^{+0.1}_{-0.1}$	$3.0^{+0.1}_{-0.1}$	
MG J1654+1346	–	–	<12.4	<2.9	–	–	–	–	
PKS J1830–211	[29.7]	–	<13.8	<4.3	–	–	–	–	synchrotron dominated
PMN J1838–3427	[32.8]	–	<13.1	<3.6	–	–	–	–	synchrotron dominated
CLASS B1933+503	20.4	3.9	13.0	3.5	$20.4^{+2.3}_{-2.1}$	$3.9^{+0.7}_{-0.6}$	$13.01^{+0.04}_{-0.03}$	$3.52^{+0.04}_{-0.03}$	
JVAS B1938+666	28.9	2.0	13.1	3.6	$29.2^{+2.7}_{-2.3}$	$2.0^{+0.3}_{-0.3}$	$13.08^{+0.04}_{-0.03}$	$3.59^{+0.04}_{-0.03}$	
PMN J2004–1349	30.7	–	12.5	3.0	$31.8^{+4.0}_{-3.1}$	–	$12.5^{+0.1}_{-0.1}$	$3.0^{+0.1}_{-0.1}$	
MG J2016+112	88.6	–	13.4	3.9	$99.5^{+32.4}_{-18.3}$	–	$13.38^{+0.05}_{-0.05}$	$3.90^{+0.05}_{-0.05}$	
WFI J2026–4536	50.1	–	13.8	4.3	$50.3^{+1.3}_{-1.2}$	–	$13.81^{+0.02}_{-0.02}$	$4.33^{+0.02}_{-0.02}$	
WFI J2033–4723	33.2	–	12.7	3.2	$33.7^{+3.0}_{-2.4}$	–	$12.7^{+0.1}_{-0.1}$	$3.3^{+0.1}_{-0.1}$	
CLASS B2045+265	–	–	<12.7	<3.2	–	–	–	–	
CLASS B2108+213	–	–	<11.7	<2.2	–	–	–	–	

Table A2 – continued

Lens Name	T_{dust} (K)	β	Least squares		T_{dust} (K)	β	MCMC		Comments
			$\log \mu L_{\text{FIR}}$ (L_{\odot})	$\log \mu \text{ SFR}$ ($M_{\odot} \text{ yr}^{-1}$)			$\log \mu L_{\text{FIR}}$ (L_{\odot})	$\log \mu \text{ SFR}$ ($M_{\odot} \text{ yr}^{-1}$)	
JVAS B2114+022	30.4	–	12.1	2.6	31.6 $^{+4.3}_{-3.1}$	–	12.1 $^{+0.2}_{-0.1}$	2.6 $^{+0.2}_{-0.1}$	
HE 2149–2745	26.9	2.8	12.8	3.3	30.7 $^{+12.0}_{-6.5}$	2.4 $^{+0.9}_{-0.9}$	12.9 $^{+0.1}_{-0.2}$	3.4 $^{+0.1}_{-0.2}$	
CY 2201–3201	–	–	<12.8	<3.3	–	–	–	–	
Q 2237+030	–	–	12.5 $^{+0.3}_{-0.4}$	3.1 $^{+0.3}_{-0.4}$	–	–	–	–	
CLASS B2319+052	19.1	4.5	12.6	3.1	19.0 $^{+4.1}_{-3.0}$	4.6 $^{+1.2}_{-1.1}$	12.6 $^{+0.1}_{-0.1}$	3.1 $^{+0.1}_{-0.1}$	
PSS J2322+1944	47.2	1.65	13.6	4.1	47.4 $^{+2.5}_{-2.3}$	1.6 $^{+0.1}_{-0.1}$	13.58 $^{+0.01}_{-0.01}$	4.10 $^{+0.01}_{-0.01}$	
SDSS J2343–0050	–	–	<11.8	<2.4	–	–	–	–	

APPENDIX B: SEDs AND ANCILLARY DATA

Table B1. Data from the literature shown in our SEDs in Fig. B1. Data points in the FIR–submillimetre frequency range that have been excluded from our SED fitting are denoted by †. Any radio data not given here are upper limits from FIRST (<1 mJy; Becker et al. 1995) or NVSS (<2.5 mJy; Condon et al. 1998) at 1.4 GHz. B102 refers to Barvainis & Ivison (2002). If errors are not given in the literature, we assume a flux calibration error of 10 per cent.

ν (GHz)	S_{ν} (Jy)	Reference
<i>QJ0158–4325</i>		
8.46	<0.0002	Morgan et al. (1999)
<i>B0218+357</i>		
229	0.57 ± 0.03	Agudo et al. (2014)
94	1.3 ± 0.3	Wright et al. (2009)
90	0.67 ± 0.07	Kühr et al. (1981)
89.3	0.52 ± 0.08	Jethava et al. (2007)
86.4	0.48 ± 0.07	Jethava et al. (2007)
86.2	0.70 ± 0.03	Agudo et al. (2014)
83.6	0.58 ± 0.09	Jethava et al. (2007)
61.1	1.0 ± 0.2	Wright et al. (2009)
43.2	0.34 ± 0.03	Jethava et al. (2007)
41	1.0 ± 0.1	Wright et al. (2009)
37	0.91 ± 0.09	Nieppola et al. (2011)
33	1.170 ± 0.216	Wright et al. (2009)
31.4	0.64 ± 0.18	Kühr et al. (1981)
31.4	0.43 ± 0.08	Kühr et al. (1981)
30	0.846 ± 0.043	Lowe et al. (2007)
23	1.104 ± 0.175	Wright et al. (2009)
22.4	1.253 ± 0.172	Patnaik et al. (1992)
16.1	1.27 ± 0.13	Davies et al. (2009)
15.064	0.74 ± 0.07	Kühr et al. (1981)
15.0	1.445 ± 0.145	Richards et al. (2011)
14.1	1.18 ± 0.12	Jethava et al. (2007)
10.695	0.90 ± 0.04	Kühr et al. (1981)
8.6	1.30 ± 0.13	Jethava et al. (2007)
8.6	0.663 ± 0.022	Zeiger & Darling (2010)
8.4	1.236 ± 0.046	Patnaik et al. (1992)
8.085	1.02 ± 0.050	Kühr et al. (1981)
4.9	1.16 ± 0.01	Kühr et al. (1981)
4.85	1.498 ± 0.188	Gregory & Condon (1991)
4.83	1.74 ± 0.17	Griffith et al. (1990)
4.585	1.04 ± 0.05	Kühr et al. (1981)
4.84	1.383 ± 0.083	Patnaik et al. (1992)
2.9	0.69 ± 0.06	Jethava et al. (2007)
2.867	0.650 ± 0.16	Zeiger & Darling (2010)
2.695	1.03 ± 0.02	Kühr et al. (1981)

Table B1 – continued

ν (GHz)	S_{ν} (Jy)	Reference
2.695	1.14 ± 0.06	Kühr et al. (1981)
1.415	0.770 ± 0.077	Patnaik et al. (1992)
1.41	1.25 ± 0.01	Kühr et al. (1981)
1.41	1.24 ± 0.06	Kühr et al. (1981)
<i>HE0230–2130</i>		
667.0	0.077 ± 0.013	BI02
353.0	0.0210 ± 0.0017	BI02
109.0	<0.0022	Riechers (2011)
1.4	<0.0025	NVSS
<i>CFRS03.1077</i>		
1.4	0.00038 ± 0.00012	FIRST
<i>B0414+054</i>		
5000	0.18 ± 0.04	Lawrence et al. (1995)
3000	<0.786	Lawrence et al. (1995)
667	0.066 ± 0.016†	BI02
353	0.0253 ± 0.0018	BI02
231	0.0207 ± 0.0013	BI02
100	0.040 ± 0.002	BI02
15.0	0.381 ± 0.006	Hewitt et al. (1992)
10.7	0.38 ± 0.02	Kellermann & Pauliny-Toth (1973)
8.6	0.340 ± 0.034	Tingay et al. (2003)
6.1	0.71 ± 0.02	Castangia et al. (2011)
5.0	0.864 ± 0.015	Hewitt et al. (1992)
5.0	0.710 ± 0.071	Wright et al. (1990)
4.8	1.110 ± 0.154	Gregory & Condon (1991)
4.8	0.959 ± 0.051	Griffith et al. (1995)
2.7	1.130 ± 0.113	Wright et al. (1990)
2.5	1.200 ± 0.012	Tingay et al. (2003)
1.4	2.087 ± 0.074	NVSS
<i>HE0435–1223</i>		
5.0	1.13 ± 0.04 × 10 ^{−4}	Jackson et al. (2015)
<i>B0631+519</i>		
15.0	0.034 ± 0.003	York et al. (2005)
8.4	0.042 ± 0.002	York et al. (2005)
4.8	0.089 ± 0.010	Gregory & Condon (1991)
1.4	0.0966 ± 0.0029	NVSS
<i>B0739+366</i>		
667	0.036 ± 0.008	BI02
353	0.0066 ± 0.0013	BI02
231	<0.0028	BI02
15	0.0219 ± 0.0019	Marlow et al. (2001)
8.4	0.0251 ± 0.0022	”
8.4	0.0241 ± 0.0021	”
5.0	0.0317 ± 0.0027	”
1.4	0.0279 ± 0.0010	NVSS

Table B1 – continued

ν (GHz)	S_ν (Jy)	Reference
<i>MGJ0751+2716</i>		
667	0.071 ± 0.015	BI02
353	0.0258 ± 0.0030	BI02
246	0.0067 ± 0.0013	Barvainis, Alloin & Bremer (2002)
220	0.0043 ± 0.0008	Alloin et al. (2007)
110	0.0030 ± 0.0005	Alloin et al. (2007)
100	0.0041 ± 0.0005	BI02
82	0.0051 ± 0.0004	Alloin et al. (2007)
42	0.0132 ± 0.0010	Carilli et al. (2005)
15	0.048 ± 0.004	Lehar et al. (1997)
8.1	0.12 ± 0.01	Condon et al. (1983)
8.3	0.104 ± 0.001	Lehar et al. (1997)
4.8	0.191 ± 0.001	Lehar et al. (1997)
4.8	0.214 ± 0.028	Gregory & Condon (1991)
2.7	0.32 ± 0.02	Condon et al. (1983)
1.4	0.595 ± 0.018	NVSS
<i>HS0810+2554</i>		
5000.0	0.284 ± 0.040	Moshir et al. (1990)
3000.0	0.538 ± 0.156	Moshir et al. (1990)
353.0	0.0076 ± 0.0018	Priddey et al. (2003)
9.0	0.000278 ± 0.000007	Jackson et al. (2015)
1.4	0.000679 ± 0.000058	Jackson et al. (2015)
<i>HS0818+1227</i>		
667.0	<0.083	BI02
353.0	0.0046 ± 0.0017	BI02
1.4	<0.001	FIRST
<i>APM08279+5255</i>		
5000	0.511 ± 0.051	Irwin et al. (1998)
4300	0.654 ± 0.009	Marton et al. (2017)
3000	0.951 ± 0.228	Irwin et al. (1998)
1875	0.759 ± 0.010	Marton et al. (2017)
857	0.386 ± 0.032	Beelen et al. (2006)
667	0.342 ± 0.026	Beelen et al. (2006)
353	0.084 ± 0.003	Lewis et al. (1998)
302	0.060 ± 0.012	Krips et al. (2007)
250	0.034 ± 0.001	Lis et al. (2011)
246	0.031 ± 0.002	van der Werf et al. (2011)
237	0.027 ± 0.001	van der Werf et al. (2011)
231	0.024 ± 0.002	Lewis et al. (1998)
201	0.017 ± 0.001	van der Werf et al. (2011)
153	0.0054 ± 0.0003	van der Werf et al. (2011)
111	0.0022 ± 0.0002	Riechers et al. (2010)
109	0.0021 ± 0.0002	"
105	0.0020 ± 0.0001	"
90.8	0.00120 ± 0.00013	García-Burillo et al. (2006)
46.9	0.000405 ± 0.000330	Riechers et al. (2009)
23.5	0.000376 ± 0.000190	"
14.9	0.000303 ± 0.000093	"
8.4	0.000446 ± 0.000020	"
4.5	0.000551 ± 0.000400	"
1.4	0.00116 ± 0.00033	"
1.4	0.0015 ± 0.0003	NVSS
<i>B0850+054</i>		
15.0	0.031 ± 0.001	Biggs et al. (2003)
8.5	0.047 ± 0.001	"
5.0	0.064 ± 0.002	"
<i>RXJ0911+0551</i>		
857	0.150 ± 0.021	Wu et al. (2009)
667	0.065 ± 0.019	BI02
353	0.027 ± 0.004	Hainline et al. (2004)
353	0.0267 ± 0.0014	BI02

Table B1 – continued

ν (GHz)	S_ν (Jy)	Reference
230	0.0102 ± 0.0018	BI02
212.5	0.0047 ± 0.0010	Tuan-Anh et al. (2013)
100.0	0.0017 ± 0.0003	BI02
30.4	<0.0017	Riechers et al. (2011)
5.0	$1.28(\pm 0.05) \times 10^{-4}$	Jackson et al. (2015)
1.4	<0.0002	Jackson et al. (2015)
<i>SDSSJ0924+0219</i>		
5.0	$1.5(\pm 0.4) \times 10^{-5}$	Jackson et al. (2015)
<i>Q0957+561</i>		
353	0.0075 ± 0.0014	BI02
231	<0.004	BI02
100	<0.0284	Planesas et al. (1999)
4.8	0.205 ± 0.022	Gregory & Condon (1991)
1.4	0.552 ± 0.017	NVSS
<i>IRAS F10214+4724</i>		
5000	0.2 ± 0.045	Moshir et al. (1990)
3000	0.57 ± 0.14	Moshir et al. (1990)
857	0.38 ± 0.05	Benford et al. (1999)
667	0.27 ± 0.05	Rowan-Robinson et al. (1993)
375	0.050 ± 0.005	"
272	0.024 ± 0.005	"
8.4	0.00027 ± 0.00005	Lawrence et al. (1993)
4.86	0.00036 ± 0.00006	"
1.49	0.00118 ± 0.0001	"
1.4	0.0018 ± 0.0001	FIRST
<i>SDSSJ1029+2623</i>		
5.0	$6.43(\pm 0.23) \times 10^{-4}$	Kratzer et al. (2011)
<i>CLASS B1030+074</i>		
216	0.114 ± 0.010	Xanthopoulos, Combes & Wiklind (2001)
144	0.225 ± 0.012	"
112	0.246 ± 0.012	"
100	0.184 ± 0.002	Barvainis et al. (2002)
22	0.320 ± 0.030	Lee et al. (2017)
22	0.231 ± 0.012	Xanthopoulos et al. (1998)
15	0.319 ± 0.016	"
15	0.223 ± 0.015	"
15	0.290 ± 0.004	Richards et al. (2011)
8.4	0.210 ± 0.011	Xanthopoulos et al. (1998)
8.4	0.218 ± 0.011	Xanthopoulos et al. (1998)
8.4	0.203 ± 0.0003	Xanthopoulos et al. (2001)
5.0	0.353 ± 0.018	Xanthopoulos et al. (1998)
4.85	0.364 ± 0.051	Gregory & Condon (1991)
4.85	0.356 ± 0.052	Becker, White & Edwards (1991)
4.85	0.242 ± 0.016	Griffith et al. (1995)
4.85	0.341 ± 0.005	Xanthopoulos et al. (2001)
4.775	0.219 ± 0.20	Bennett et al. (1986)
1.7	0.258 ± 0.013	Xanthopoulos et al. (1998)
1.4	0.163 ± 0.004	White & Becker (1992)
1.4	0.155 ± 0.004	NVSS
1.4	0.156 ± 0.005	Xanthopoulos et al. (2001)
<i>HE1104-1805</i>		
353	0.015 ± 0.003	BI02
231	0.0053 ± 0.0009	BI02
100	<0.0022	BI02
1.4	<0.0025	NVSS
<i>PG1115+080</i>		
353	0.0037 ± 0.0013	BI02
231	<0.003	BI02
104	<0.0015	Riechers (2011)
100	<0.005	BI02
1.4	<0.00094	FIRST

Table B1 – continued

ν (GHz)	S_ν (Jy)	Reference
<i>B1127+385</i>		
667	<0.065	BI02
353.0	0.014 ± 0.002	BI02
231	<0.0028	BI02
8.4	0.027 ± 0.002	Koopmans et al. (1999)
5.0	0.027 ± 0.002	Koopmans et al. (1999)
4.85	0.041 ± 0.007	Gregory & Condon (1991)
4.85	0.037 ± 0.006	Becker et al. (1991)
1.7	0.030 ± 0.002	Koopmans et al. (1999)
1.4	0.029 ± 0.001	Becker et al. (1995)
<i>RX J1131–1231</i>		
216.0	<0.0025	Leung, Riechers & Pavesi (2017)
144.1	$0.00195 \pm 0.00020^\dagger$	Paraficz et al. (2017)
139.3	$0.00039 \pm 0.00012^\dagger$	Leung et al. (2017)
4.9	0.00127 ± 0.00004	Leung et al. (2017)
1.4	<0.028	FIRST
<i>B1152+200</i>		
667	<0.07	BI02
353	<0.0065	BI02
14.94	0.0581 ± 0.0004	Myers et al. (1999)
8.46	0.0695 ± 0.0005	Myers et al. (1999)
4.85	0.070 ± 0.0011	Gregory & Condon (1991)
1.4	0.0774 ± 0.008	NVSS
<i>Q1208+101</i>		
353	0.0081 ± 0.0020	BI02
250	0.0042 ± 0.0019	Andreani, Franceschini & Granato (1999)
231	0.003 ± 0.001	BI02
100	<0.0009	BI02
1.4	<0.00095	FIRST
<i>SDSSJ1330+1810</i>		
1.4	<0.0002	Stacey (2015)
<i>SDSSJ1353+1138</i>		
1.4	0.0005 ± 0.00015	FIRST
<i>B1359+154</i>		
667	$0.039 \pm 0.010^\dagger$	BI02
353	0.012 ± 0.002	BI02
96.4	0.00360 ± 0.00018	Riechers (2011)
14.94	0.01625 ± 0.0009	Myers et al. (1999)
8.46	0.0279 ± 0.0014	Myers et al. (1999)
4.8	0.066 ± 0.010	Becker et al. (1991)
1.4	0.115 ± 0.035	NVSS
<i>H1413+117</i>		
5080	0.207 ± 0.051	Weiβ et al. (2003)
5000	0.230 ± 0.078	Barvainis et al. (1995)
3261	0.266 ± 0.050	Weiβ et al. (2003)
3000	0.370 ± 0.078	Barvainis et al. (1995)
2520	0.356 ± 0.048	Weiβ et al. (2003)
1630	0.240 ± 0.095	Weiβ et al. (2003)
1500	0.280 ± 0.014	Rowan-Robinson (2000)
870	0.189 ± 0.056	Barvainis, Antonucci & Coleman (1992)
857	0.376 ± 0.015	Weiβ et al. (2003)
857	0.293 ± 0.014	Benford et al. (1999)
685	0.224 ± 0.038	Barvainis et al. (1992)
394	0.044 ± 0.008	Barvainis et al. (1992)
375	0.066 ± 0.007	Hughes, Dunlop & Rawlings (1997)
353	0.059 ± 0.008	BI02
250	0.016 ± 0.002	Weiβ et al. (2003)
240	0.018 ± 0.002	Barvainis et al. (1995)
230	0.0075 ± 0.0006	Weiβ et al. (2003)

Table B1 – continued

ν (GHz)	S_ν (Jy)	Reference
100	<0.0015	Weiβ et al. (2003)
92.9	0.0003 ± 0.0001	Stacey et al., in preparation
24.0	0.00026 ± 0.00003	Solomon et al. (2003)
14.9	0.00056 ± 0.00018	Barvainis & Lonsdale (1997)
8.5	0.00098 ± 0.00008	"
4.9	0.00195 ± 0.00013	"
1.5	0.00768 ± 0.00050	"
1.4	0.0082 ± 0.0006	NVSS
<i>B1422+231</i>		
22.5	0.145 ± 0.015	Tinti et al. (2005)
15.0	0.251 ± 0.013	"
8.5	0.460 ± 0.016	"
8.1	0.479 ± 0.016	"
4.9	0.669 ± 0.020	"
4.5	0.686 ± 0.021	"
1.7	0.414 ± 0.012	"
1.4	0.352 ± 0.011	"
<i>SBS1520+530</i>		
353	<0.0078	BI02
250	<0.0042	BI02
100	<0.0006	BI02
1.4	<0.001	FIRST
<i>B1600+434</i>		
353	0.0073 ± 0.0018	BI02
300	0.0126 ± 0.0023	BI02
100	$0.0250 \pm 0.0003^\dagger$	BI02
15	0.042 ± 0.002	WalDRAM et al. (2003)
8.4	0.132 ± 0.013	Jackson et al. (1995)
4.8	0.037 ± 0.007	Gregory & Condon (1991)
1.4	0.079920 ± 0.000145	FIRST
<i>B1608+656</i>		
353.0	0.0081 ± 0.0017	BI02
300.0	<0.0066	Xanthopoulos et al. (2001)
231.0	0.0056 ± 0.0017	BI02
100.0	0.0081 ± 0.0004	BI02
15	0.081 ± 0.004	Snellen et al. (1995)
8.4	0.083 ± 0.004	Snellen et al. (1995)
8.4	0.0732 ± 0.0020	Myers et al. (1995)
4.8	0.0920 ± 0.0090	Gregory & Condon (1991)
1.4	0.063 ± 0.003	Snellen et al. (1995)
<i>PMN J1632–0033</i>		
43.34	0.089 ± 0.014	Winn et al. (2002)
43.34	0.112 ± 0.006	Winn, Rusin & Kochanek (2004)
22.46	0.135 ± 0.014	Winn et al. (2002)
22.46	0.161 ± 0.008	Winn et al. (2004)
22.46	0.158 ± 0.008	Winn et al. (2004)
14.94	0.153 ± 0.007	Winn et al. (2002)
14.94	0.187 ± 0.009	Winn et al. (2004)
14.94	0.195 ± 0.010	Winn et al. (2004)
8.64	0.177 ± 0.005	Winn et al. (2002)
8.46	0.220 ± 0.011	Winn et al. (2004)
8.46	0.227 ± 0.011	"
8.46	0.211 ± 0.011	"
8.46	0.152 ± 0.008	McKean et al. (2007)
8.45	0.165 ± 0.030	Winn et al. (2002)
8.44	0.160 ± 0.040	"
6.1	0.201 ± 0.006	"
5.0	0.191 ± 0.020	"
5.0	0.222 ± 0.011	Winn, Rusin & Kochanek (2003)
4.8	0.233 ± 0.007	Winn et al. (2002)
4.86	0.223 ± 0.011	Winn et al. (2002)

Table B1 – continued

ν (GHz)	S_ν (Jy)	Reference
4.86	0.204 ± 0.010	McKean et al. (2007)
4.85	0.270 ± 0.040	Becker et al. (1991)
3.9	0.235 ± 0.047	Larionov et al. (1994)
2.7	0.200 ± 0.040	Wright et al. (1990)
1.4	0.230 ± 0.012	Winn et al. (2002)
1.4	0.236 ± 0.047	White & Becker (1992)
<i>FBQS1633+3134</i>		
353.0	<0.0035	BI02
1.4	0.00177 ± 0.00014	FIRST
<i>PKS 1830–211</i>		
353	0.79 ± 0.17	Giommi et al. (2012)
300	0.8 ± 0.1	Martí-Vidal et al. (2013)
250	0.9 ± 0.1	Martí-Vidal et al. (2013)
229	1.23 ± 0.06	Agudo et al. (2010)
217	1.18 ± 0.08	Giommi et al. (2012)
143	1.64 ± 0.07	Giommi et al. (2012)
100	2.0 ± 0.02	Muller et al. (2006)
100	2.47 ± 0.11	Giommi et al. (2012)
86	1.76 ± 0.09	Agudo et al. (2010)
70	2.44 ± 0.22	Giommi et al. (2012)
44	3.42 ± 0.36	Giommi et al. (2012)
33	4.12 ± 0.24	Massardi et al. (2009)
30	4.11 ± 0.33	Giommi et al. (2012)
23	5.08 ± 0.22	Massardi et al. (2009)
20	5.50 ± 0.36	Massardi et al. (2008)
14	6.45 ± 0.65	Henkel et al. (2008)
12.8	8.12 ± 0.81	Henkel et al. (2008)
8.4	6.59 ± 0.66	Wright et al. (1990)
6.55	9.74 ± 0.97	Ellingsen et al. (2012)
5	8.9 ± 0.9	Wright et al. (1990)
4.9	7.92 ± 0.10	Griffith et al. (1994)
2.7	9.3 ± 0.93	Wright et al. (1990)
1.4	10.90 ± 0.33	NVSS
<i>PMN J1838–3427</i>		
20	0.234 ± 0.012	Murphy et al. (2010)
14.94	0.181 ± 0.009	Winn et al. (2000)
8.64	0.264 ± 0.009	"
8.46	0.206 ± 0.006	"
8.46	0.181 ± 0.005	"
8	0.284 ± 0.014	Murphy et al. (2010)
5	0.299 ± 0.015	Murphy et al. (2010)
4.86	0.214 ± 0.006	Winn et al. (2000)
4.85	0.258 ± 0.021	Wright et al. (1990)
4.80	0.219 ± 0.007	Winn et al. (2000)
2.7	0.240 ± 0.048	Wright et al. (1990)
1.4	0.280 ± 0.008	NVSS
<i>B1933+503</i>		
3000	<0.443	Chapman et al. (1999)
667	0.114 ± 0.017	"
353	0.0240 ± 0.0026	"
231	0.030 ± 0.007	"
15	0.0371 ± 0.0011	Sykes et al. (1998)
8.4	0.0410 ± 0.0011	"
5.0	0.0575 ± 0.0057	"
1.7	0.0759 ± 0.0075	"
<i>B1938+666</i>		
667	0.126 ± 0.022	BI02
353	0.0346 ± 0.0020	BI02
231	0.0147 ± 0.0020	BI02
97	0.0200 ± 0.0014	Riechers (2011)
22	0.088 ± 0.009	King et al. (1997)
15	0.141 ± 0.14	King et al. (1997)

Table B1 – continued

ν (GHz)	S_ν (Jy)	Reference
8.4	0.224 ± 0.022	Patnaik et al. (1992)
4.8	0.314 ± 0.047	Becker et al. (1991)
4.8	0.316 ± 0.028	Gregory & Condon (1991)
1.4	0.5768 ± 0.0173	NVSS
1.4	0.634 ± 0.063	FIRST
<i>PMNJ2004–1349</i>		
22.46	0.0164 ± 0.002	Winn et al. (2001)
14.96	0.0205 ± 0.0010	"
8.46	0.0294 ± 0.0090	"
5.0	0.030 ± 0.002	"
4.8	0.073 ± 0.011	Griffith et al. (1994)
1.4	0.079 ± 0.08	NVSS
<i>MGJ2016+112</i>		
353.0	<0.0048	Barvainis et al. (2002)
231.0	<0.0025	"
100.0	0.0018 ± 0.0002	"
4.8	0.098 ± 0.010	Bennett et al. (1986)
1.4	0.1911 ± 0.0058	NVSS
<i>B2114+022</i>		
353.0	<0.0043	BI02
87	<0.030	Xanthopoulos et al. (2001)
15.0	0.051 ± 0.002	Augusto et al. (2001)
8.4	0.100 ± 0.002	"
5.0	0.156 ± 0.016	"
5.0	0.230 ± 0.023	Wright et al. (1990)
4.85		Griffith et al. (1990)
4.85	0.136 ± 0.003	Vollmer et al. (2008)
4.775		Bennett et al. (1986)
2.6		Wright et al. (1990)
1.7	0.148 ± 0.015	Augusto et al. (2001)
1.4	0.137 ± 0.004	NVSS
<i>HE2149–2745</i>		
353.0	0.008 ± 0.002	BI02
100.0	<0.0064	BI02
1.4	<0.0025	NVSS
<i>Q2237+030</i>		
667	<0.017	BI02
353	0.0039 ± 0.0012	BI02
250	<0.0064	BI02
100	<0.0008	BI02
8.4	0.000593 ± 0.000088	Falco et al. (1996)
1.5	0.000832 ± 0.000087	Falco et al. (1996)
<i>B2319+052</i>		
667.0	0.040 ± 0.008	BI02
353.0	0.0039 ± 0.0012	BI02
231.0	<0.003	BI02
15.0	0.0182 ± 0.0006	Rusin et al. (2001)
8.4	0.0308 ± 0.0001	"
5.0	0.0666 ± 0.0001	"
1.4	0.0853 ± 0.0004	"
<i>PSSJ2322+1944</i>		
4300	$0.0137 \pm 0.0061^\dagger$	Marton et al. (2017)
1875	$0.0434 \pm 0.0084^\dagger$	"
660	0.075 ± 0.019	Cox et al. (2002)
353	0.0225 ± 0.0025	Isaak et al. (2002)
353	0.024 ± 0.002	Cox et al. (2002)
231	0.0096 ± 0.0005	Omont et al. (2001)
225	0.0075 ± 0.0013	Cox et al. (2002)
90	<0.00064	Cox et al. (2002)
5.0	$<9 \times 10^{-5}$	Carilli et al. (2001)
1.4	$9.8(\pm 1.5) \times 10^{-5}$	Carilli et al. (2001)

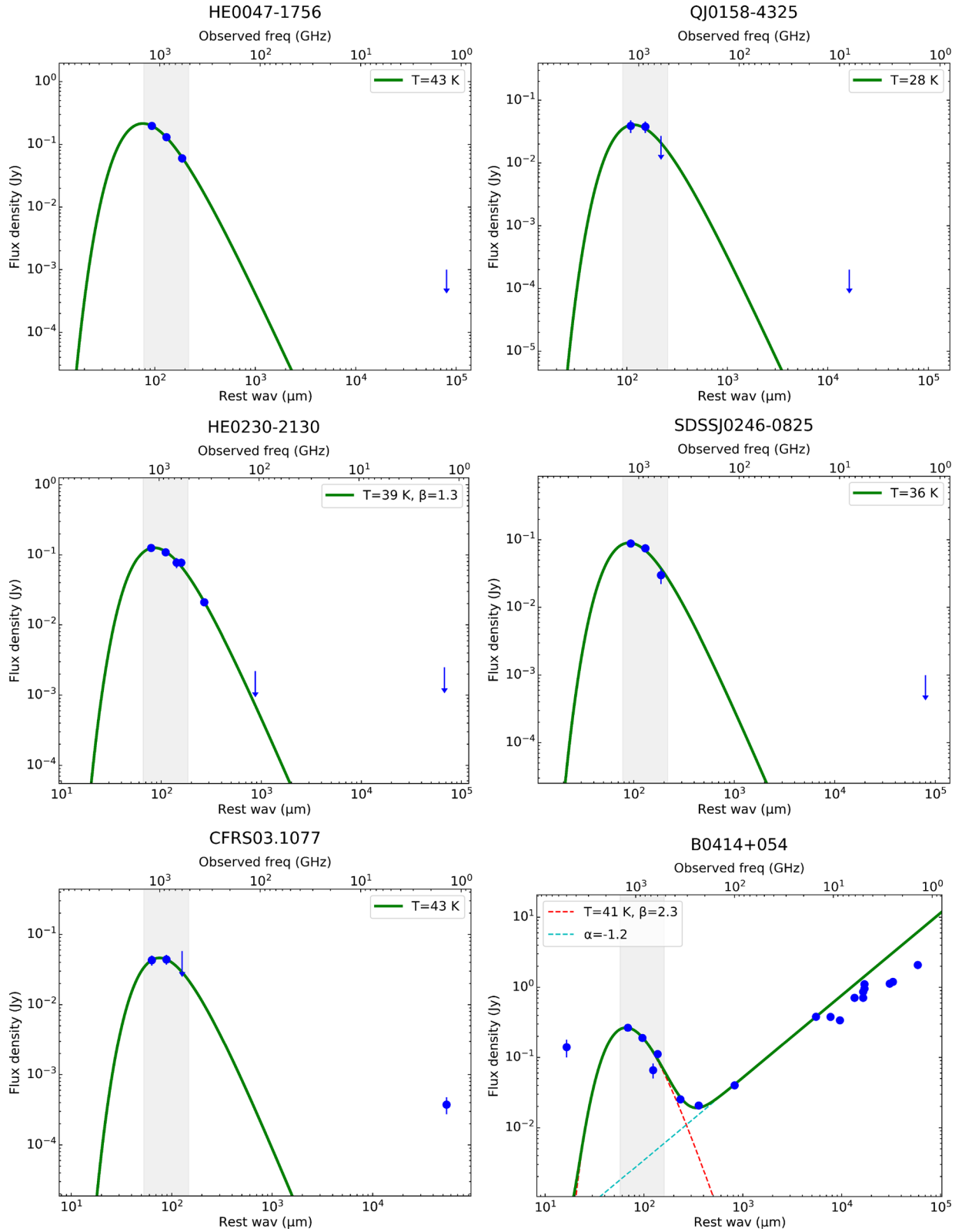


Figure B1. Included here are 69 SEDs for the quasars in our sample with fitted SEDs, excluding synchrotron-dominated sources (see Fig. B2). The legend details the free parameters of the model and their least-squares values, excluding the normalization (FIR luminosities are given in Table A2). Where β is not given, it is fixed to 1.5. For objects with T fixed, the dust temperature is set to 38 K, the median of the sample. For composite spectra, modified blackbody fits are in red, synchrotron in cyan, and the total spectrum in green. The *Herschel*/SPIRE bandwidth is in grey. As we discuss in Section 3.4, we do not attempt to fit complex synchrotron components in cases where there is a suggestion that the synchrotron emission is falling off towards the submillimetre or if there is only one detection.

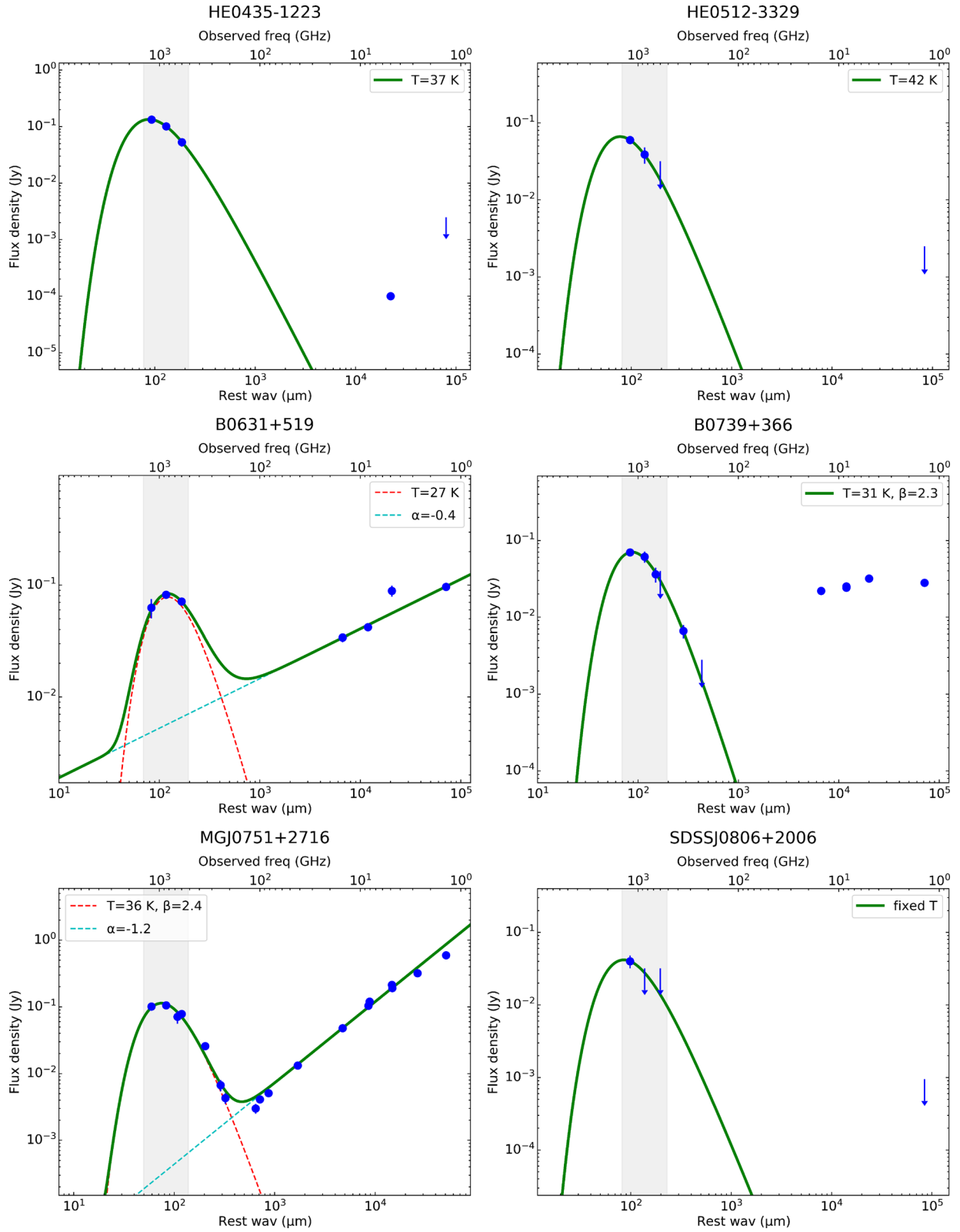


Figure B1 – continued

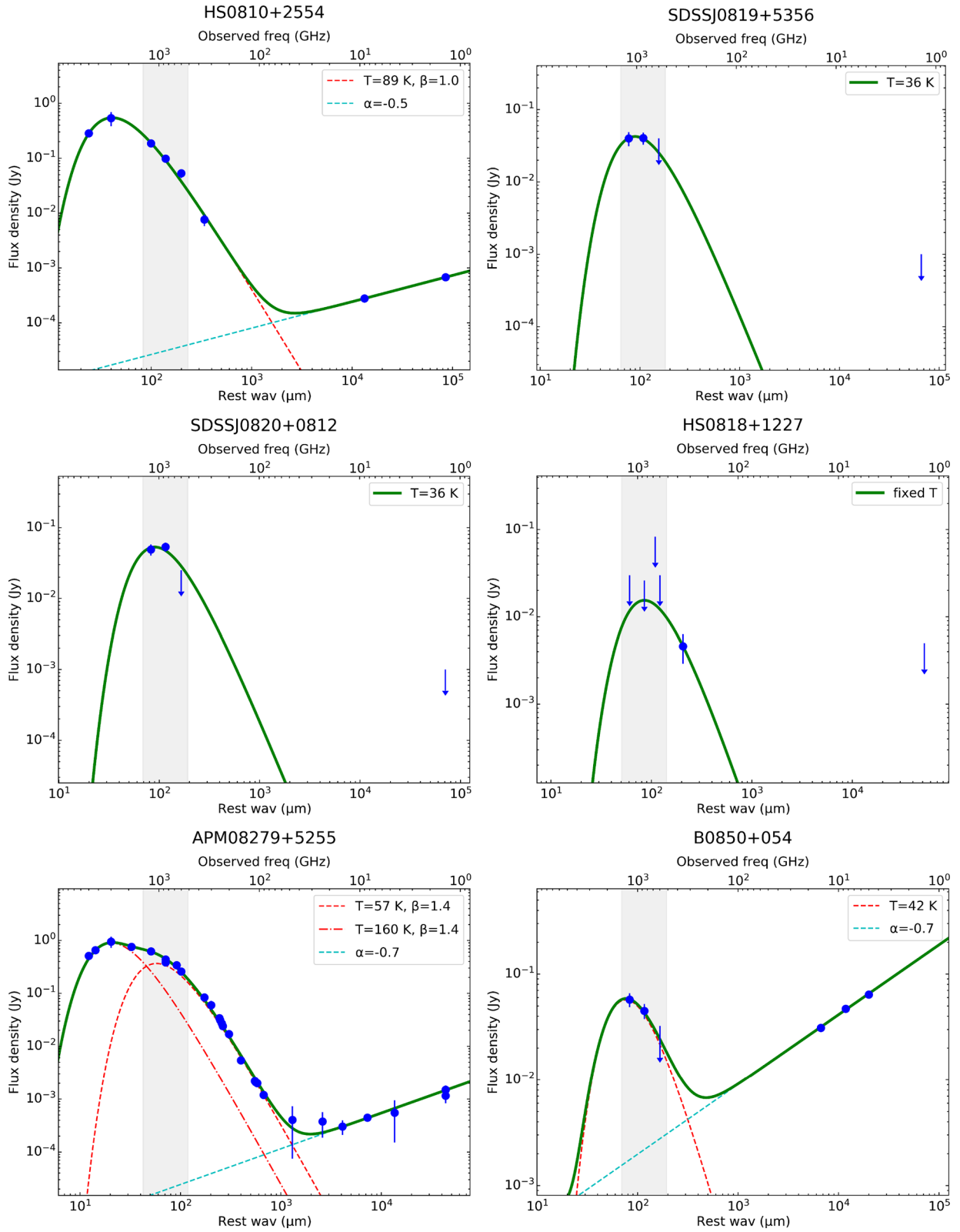
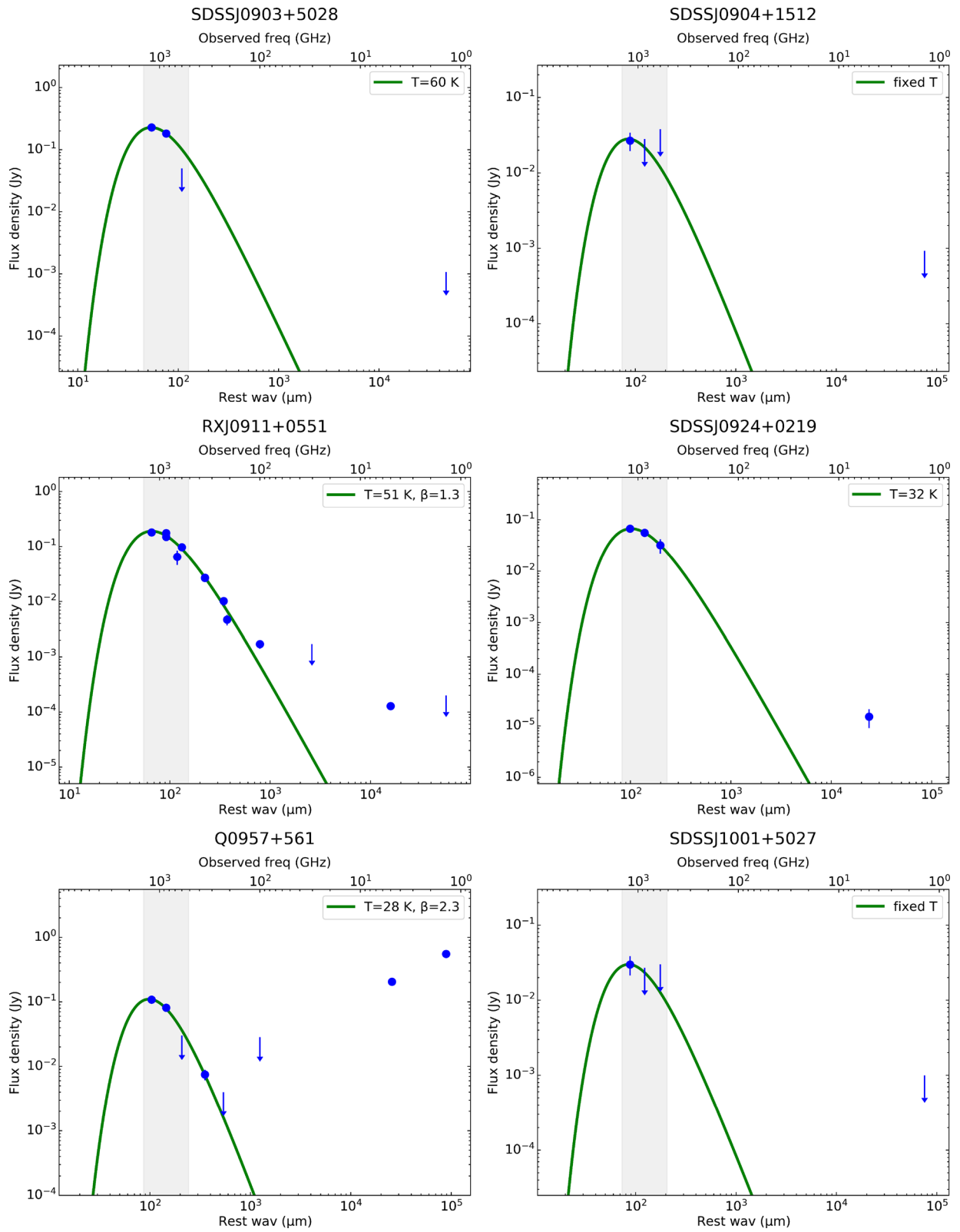


Figure B1 – continued

Figure B1 – *continued*

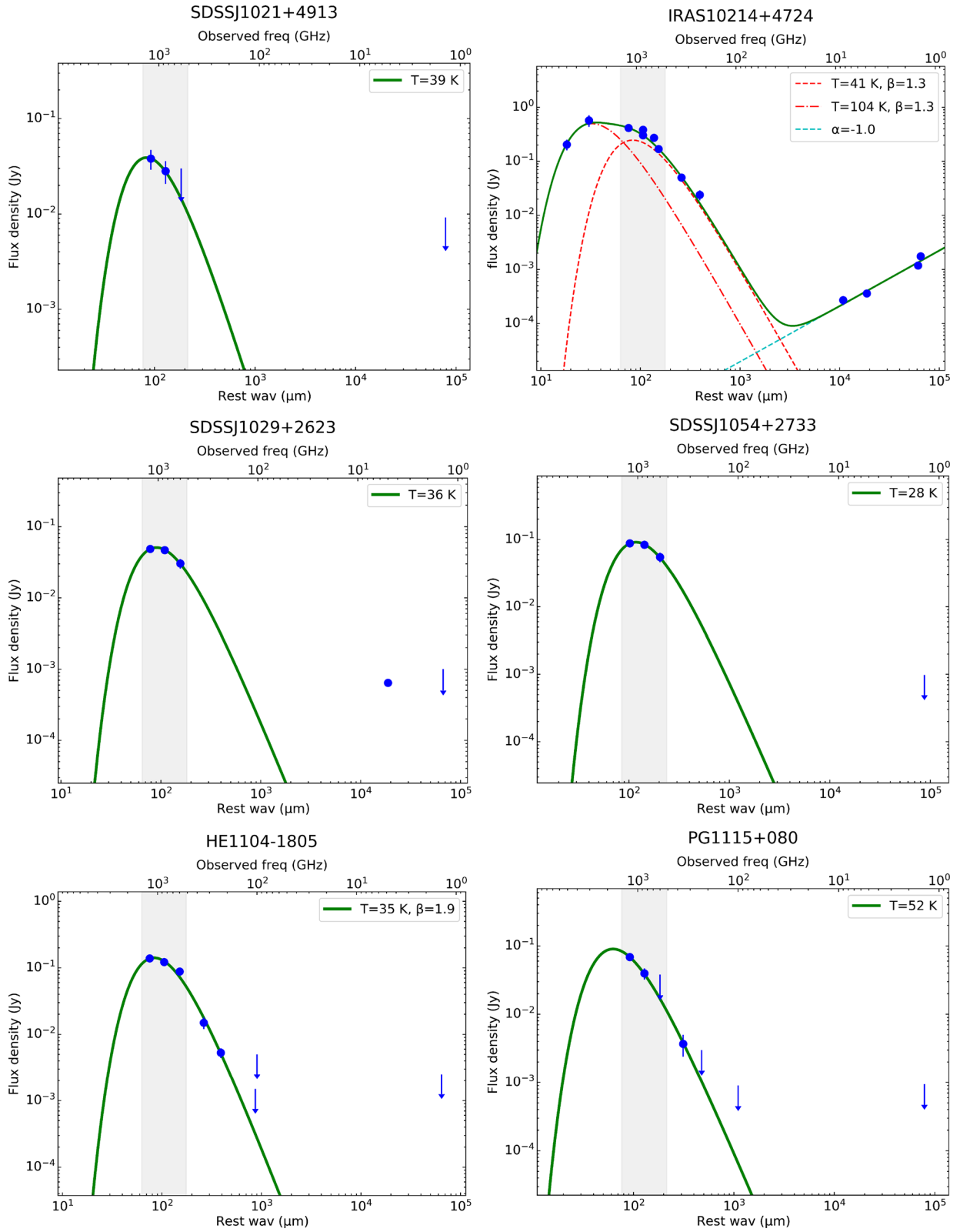
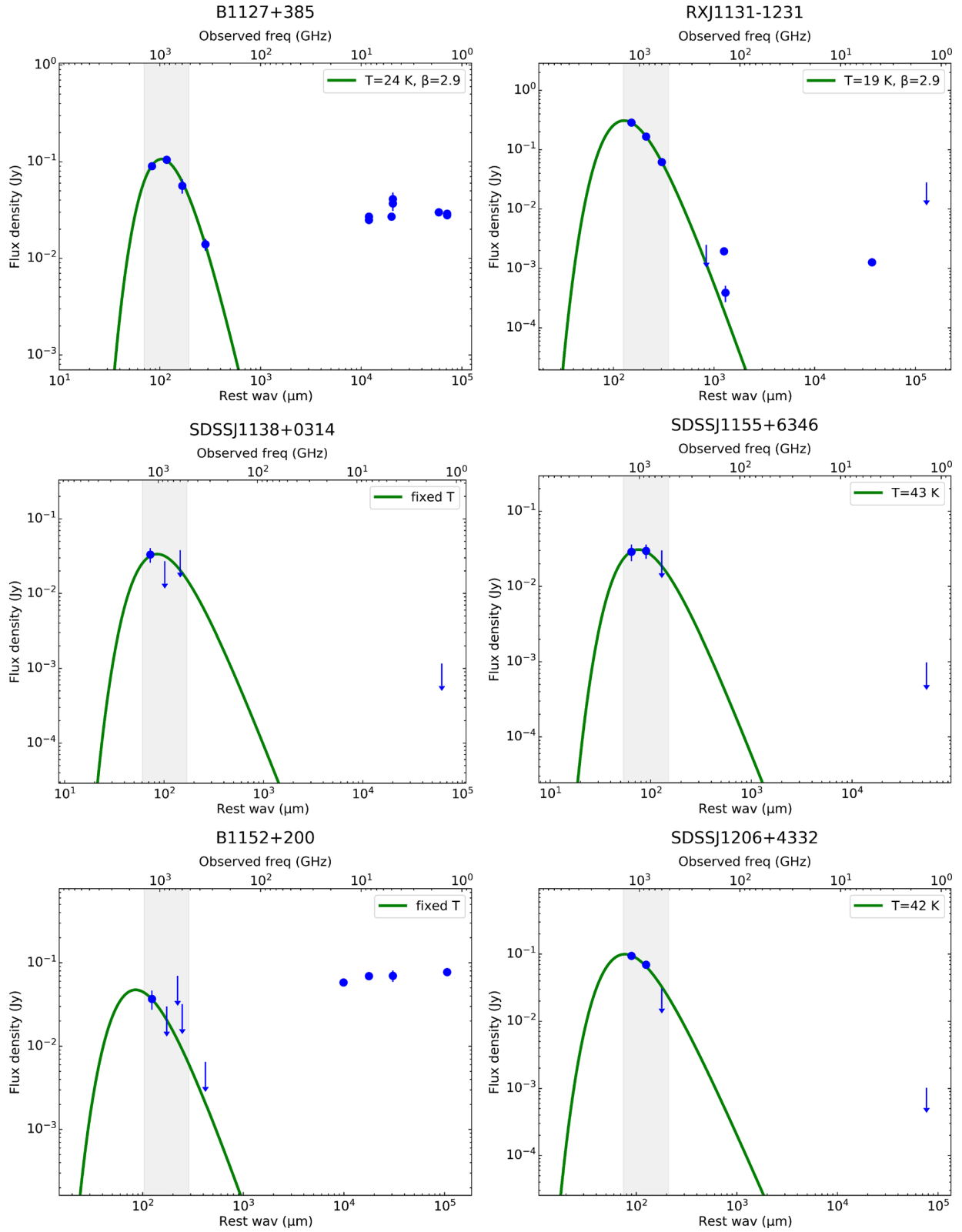


Figure B1 – continued

Figure B1 – *continued*

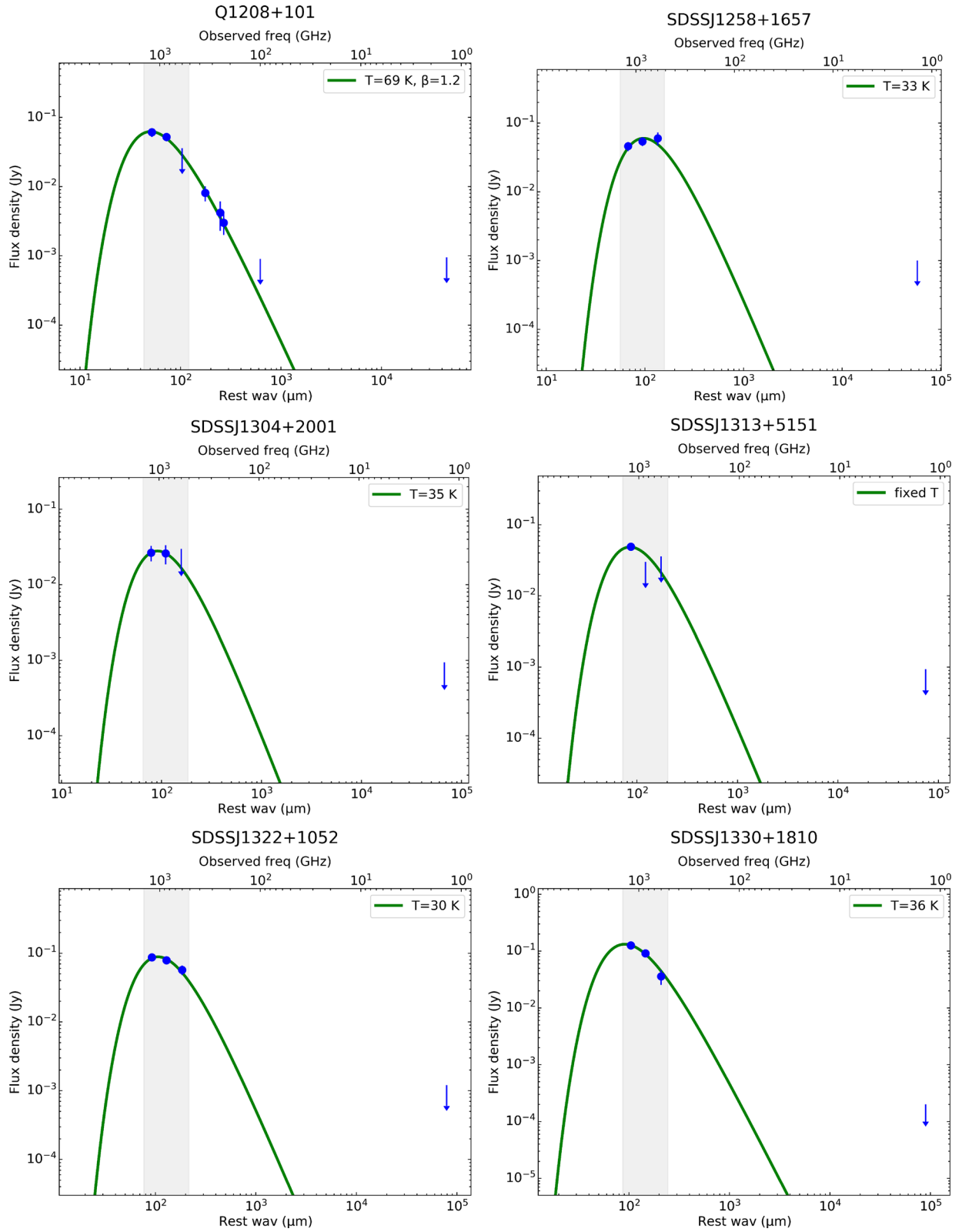
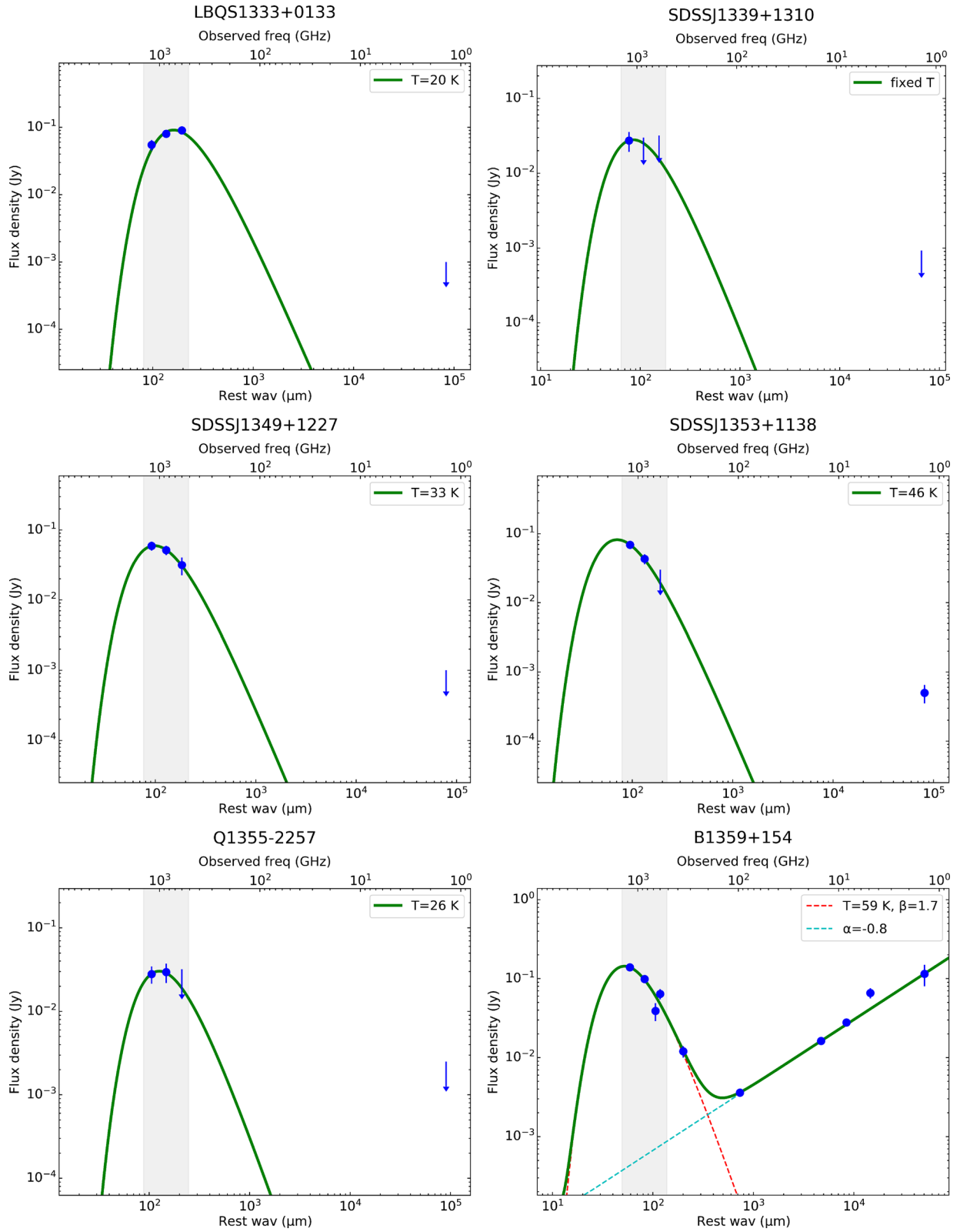


Figure B1 – continued

Figure B1 – *continued*

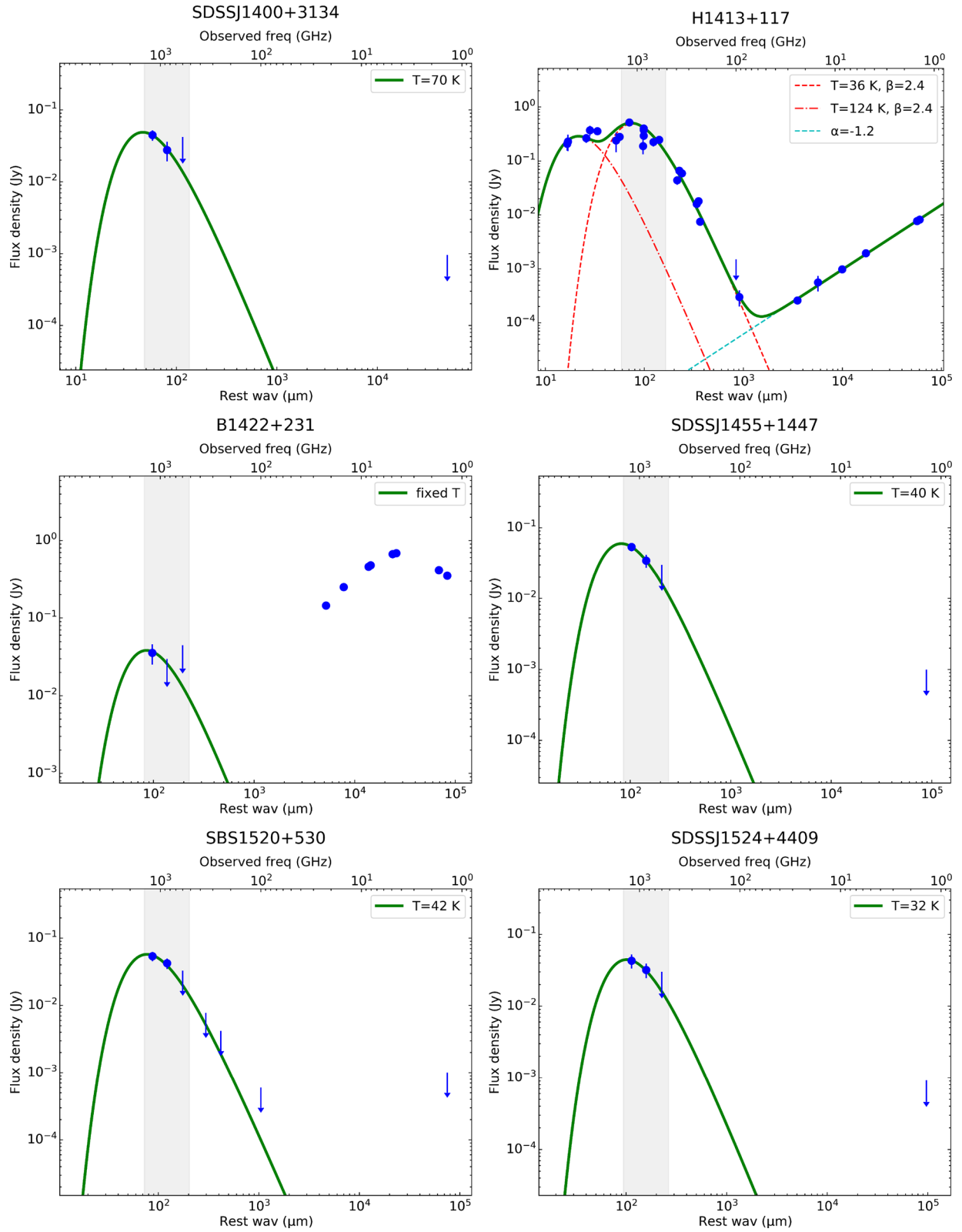
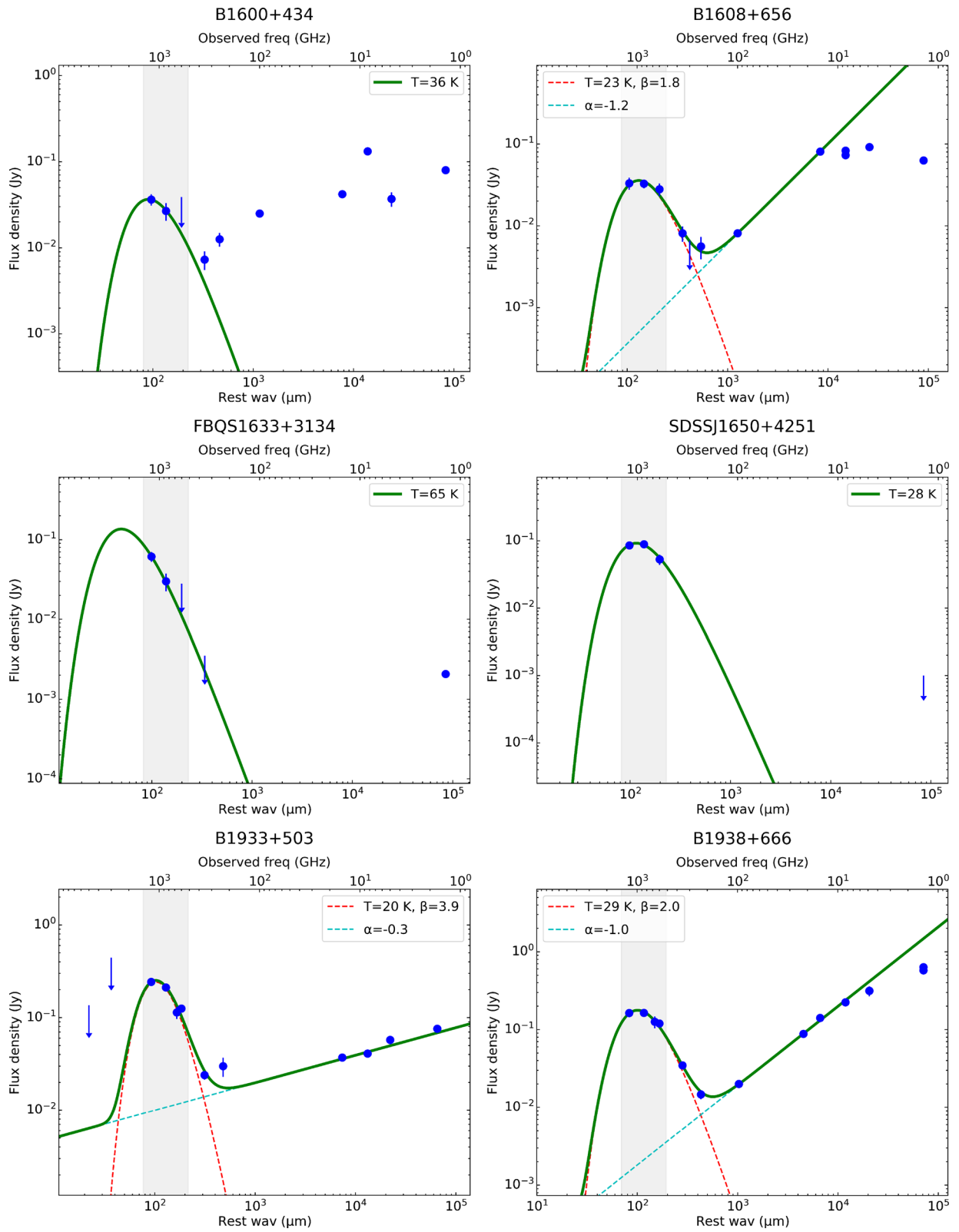


Figure B1 – continued

Figure B1 – *continued*

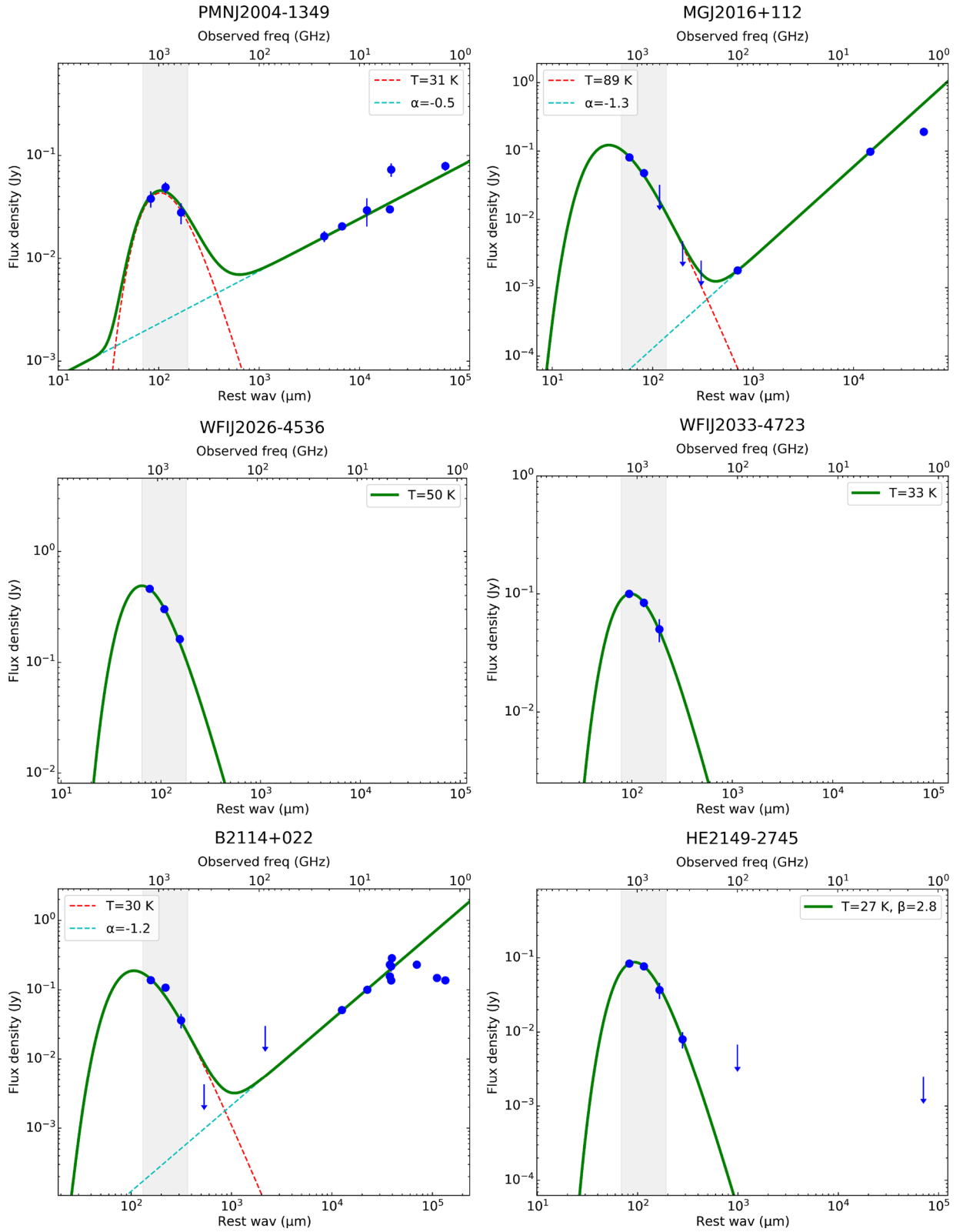
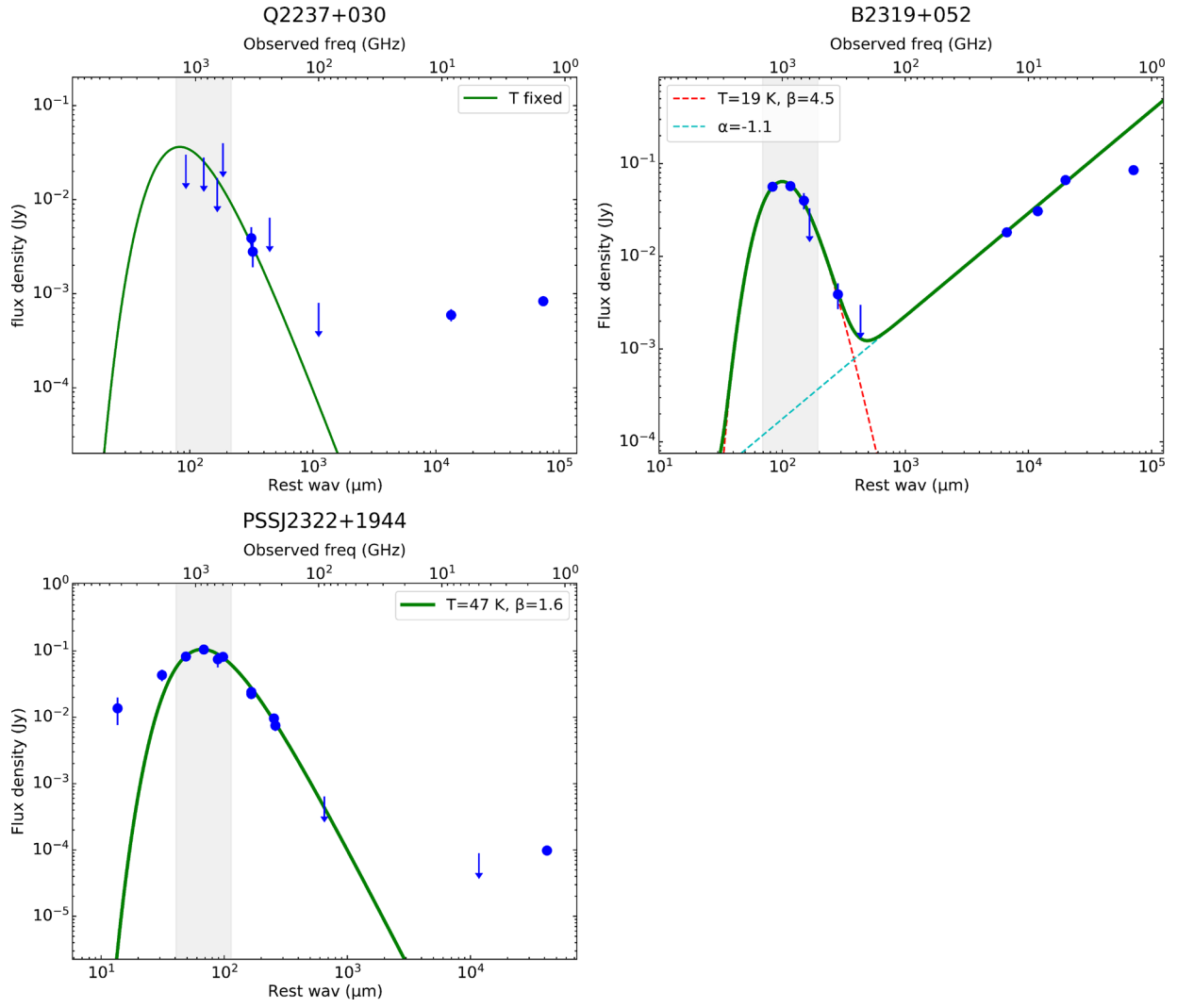


Figure B1 – continued

Figure B1 – *continued*

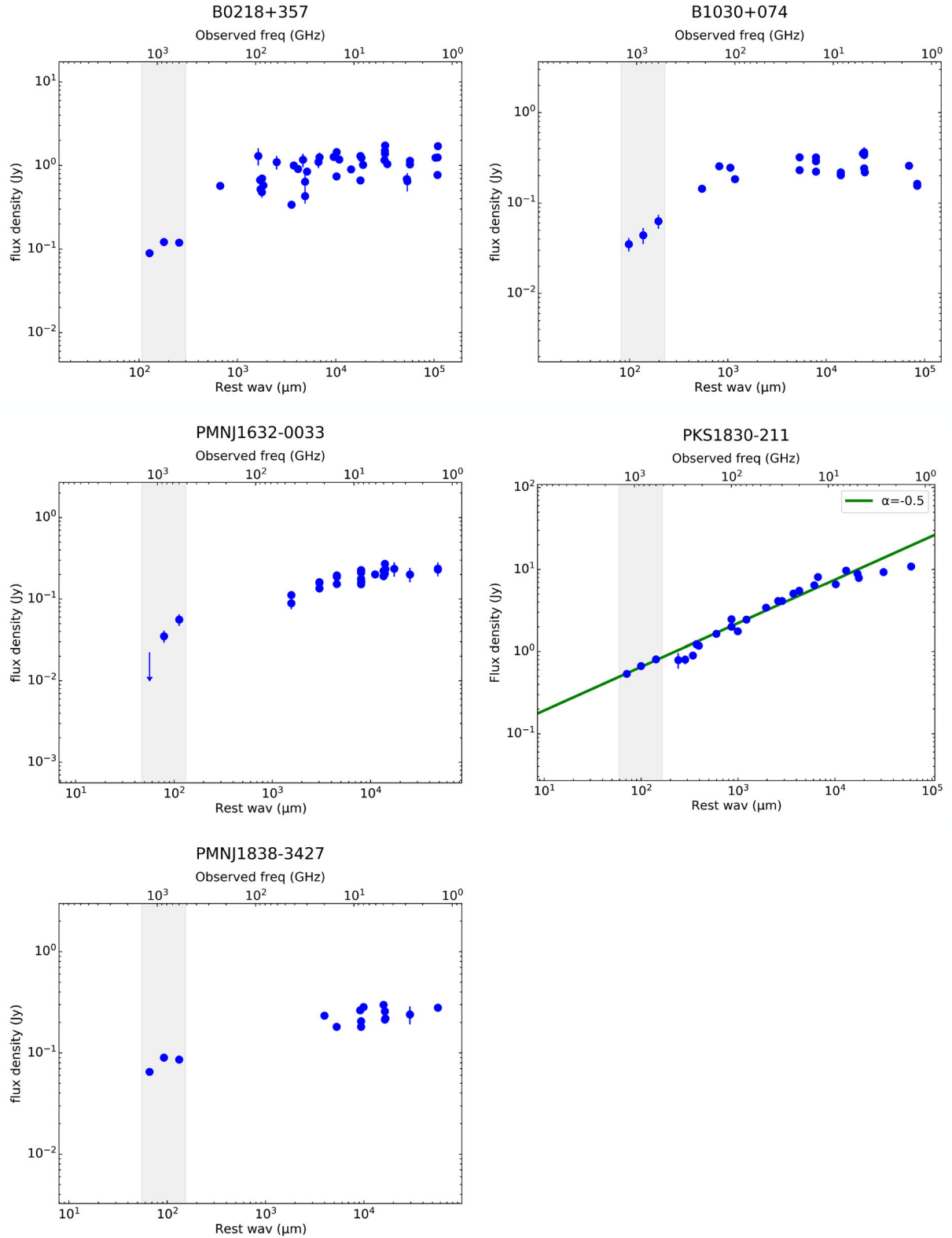


Figure B2. Included here are SEDs for 5 quasars that appear to have synchrotron-dominated emission in the FIR.

APPENDIX C: NOTES ON INDIVIDUAL SOURCES

We discuss the results for a few individual sources of note.

C1 HS 0810+2554

HS 0810+2554 is an outlier of our sample in several respects. It has a high effective dust temperature ($T_d = 89.0^{+6.5}_{-6.0}$ K) and the lowest dust emissivity index ($\beta = 1.0 \pm 0.2$) of our sample. Such a high dust temperature is more consistent with dust heated by the AGN than star formation (e.g. APM 08279+5255; Weiß et al. 2007).

Despite its weak radio flux density ($\sim \mu\text{Jy}$), this source has a mas-scale radio jet that is the major contributor to its radio emission (Hartley et al., in preparation). Radio observations with the VLA and the Multi-Element Radio Linked Interferometer Network (e-MERLIN) also imply a compact radio-emitting region with a scale of 70 pc (Jackson et al. 2015). However, we find this source falls below the radio–infrared correlation, rather than above as would be expected for a source with a radio excess.

The emissivity index of 1.0 is lower than the typically observed values of $\beta = 1.5$ –2 for star-forming galaxies. These properties could be a result of composite dust emission from both AGN and star formation heating, similar to that observed in APM 08279+5255, IRAS F10214+4724 and the Cloverleaf (Beelen et al. 2006). Differential magnification could be responsible for a boosting of the more compact AGN-heated component.

The measured L_{FIR} and SFR given in Table A2 of this Appendix A are from a single-temperature model, and likely an overestimate the actual properties of this quasar. Additional data, taken at millimetre and submillimetre wavelengths will be needed to properly separate the two components of the true SED.

We try fitting a two-component dust model, one of fixed temperature 38 K (the median-fitted dust temperature of the sample) and leave the second temperature as a free parameter, both with fixed $\beta = 1.5$. We find a fit for the warmer component of $84.4^{+6.5}_{-5.7}$ K and derive a FIR luminosity of $3.7^{+1.7}_{-2.3} \times 10^{12} L_\odot$ for the cold component: almost an order of magnitude lower than that from the single-temperature model. The result is consistent within 2σ with the radio–infrared correlation and falls amid non-jetted sources.

C2 RX J1131–1231

RX J1131–1231 is one of the lower redshift quasars in our sample, at $z = 0.67$. The quasar is known to have a radio jet based on JVLA observations (Wucknitz & Volino 2008). Leung et al. (2017) observed RX J1131–1231 with the Plateau de Bure Interferometer (PdBI) and Combined Array for Research in Millimeter-wave Astronomy (CARMA) at 2.2 and 3 mm, respectively. They derive star-forming properties of this source by fitting an SED, assuming both the PdBI measurement and CARMA upper-limit describe the Rayleigh–Jeans slope of the modified blackbody. We consider these data and also include a recent ALMA observation at 2.1 mm (Paraficz et al. 2017); however, we find significant differences between the ALMA 2.1 mm and PdBI 2.2 mm measurements. Paraficz et al. (2017) propose the difference is due to a contribution from synchrotron emission at the base of the jet associated with the AGN, which could be either highly variable, or so compact ($\sim 10^{-4}$ pc) that micro-lensing may be changing the flux density over time-scales of months (the observations were performed 5–7 months apart: PdBI between 2014 December and 2015 February; ALMA in 2015 July). We include only the *Herschel*/SPIRE measurements

and the CARMA upper limit to constrain the thermal dust emission, finding a relatively low dust temperature of $T_d = 21^{+6}_{-4}$ K and a high emissivity index of $\beta = 2.7^{+1.0}_{-0.7}$; however, this is not robust as the peak is poorly constrained. We do not attempt to fit a synchrotron component due to the uncertainties discussed here and by Paraficz et al. (2017). Further, high- and low-wavelength data are needed to better constrain the dust temperature and L_{FIR} and characterize the millimetre emission for this object.

C3 H 1413+117

The *Cloverleaf* quasar (H 1413+117) has been studied extensively over the past ~ 20 yr, as it is one of the most FIR-luminous gravitationally lensed quasars known, and so there are many measurements in the literature that cover the full infrared SED. The SED can be resolved into two dust peaks that are presumably due to heating by both star formation ($T_d = 35.6 \pm 0.6$ K) and AGN activity ($T_d = 125.6^{+10.6}_{-8.9}$ K). We find this quasar falls above the radio–infrared correlation with a q_{IR} value of 1.42 ± 0.01 , consistent with jet emission known to exist in this source based on radio observations with e-MEJIN (Stacey et al., in preparation).

With the addition of the *Herschel*/SPIRE data, we see clear differences in the measurements around the lower-temperature peak over a period of ~ 20 yr. This is most obvious with the four measurements around 350 μm , which have increased intermittently to a factor of 2 relative to the first measurement by Barvainis et al. (1992). This is likely the effect of calibration errors in previous measurements rather than intrinsic variability.

The spread of the data causes uncertainty in the SED fitting, as the AGN contribution and Rayleigh–Jeans slope are not well constrained. However, the new *Herschel*/SPIRE data constrain the peak of the SF-heated dust component accurately; thus, assuming the contribution from the AGN component is small at those wavelengths, the effect of the uncertainty on the derived FIR luminosity due to star formation is not significant.

C4 PKS 1830-211

PKS 1830–211 is a radio-powerful gravitationally lensed blazar; a radio source that is being viewed directly down the line of sight of the relativistic jet (Martí-Vidal et al. 2013). The SED appears to be dominated by synchrotron emission from the radio through to the FIR measured with *Herschel*/SPIRE. While there appears to be a tentative suggestion that the synchrotron component begins to fall off towards the FIR, this is far from clear because this source is highly variable. We fit the data > 10 GHz with both a simple power law, typical of optically thin synchrotron emission, and a function that includes a modified blackbody (leaving the temperature as a free parameter) to account for the possibility of a contribution from thermal dust. We assume this grey body represents an upper limit on the star-forming properties, meaning there could be underlying dust-obscured star formation in the host galaxy at a rate as high as $\sim 1000 M_\odot \text{ yr}^{-1}$. The data are also consistent with a simple power law of $\alpha = -0.5$ that begins to flatten ~ 10 GHz.

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